Theories of dark matter

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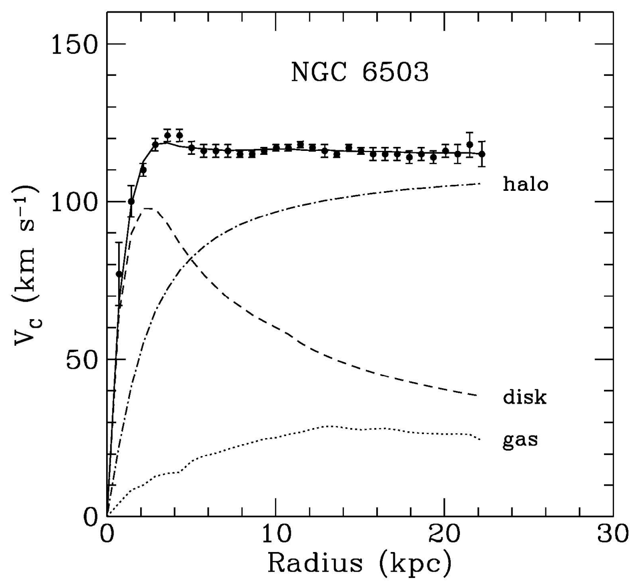
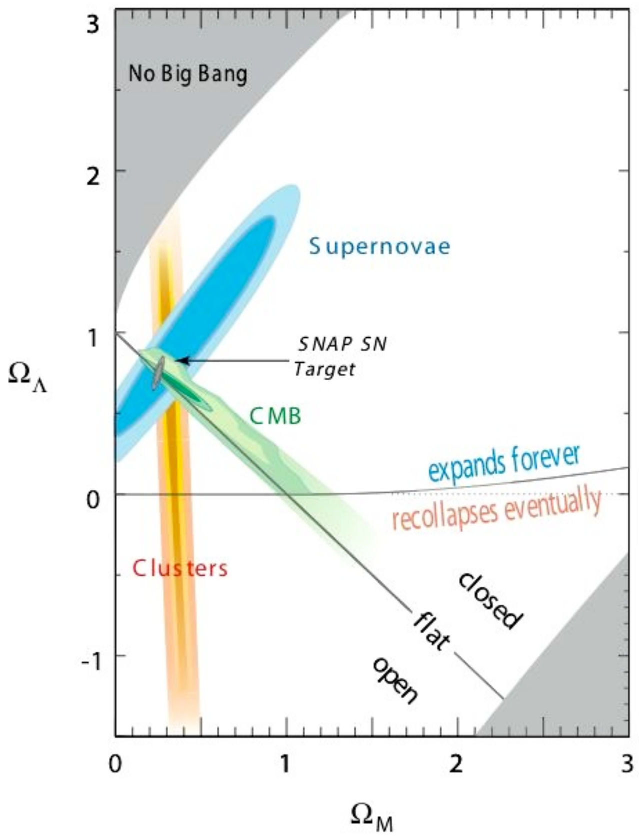
University of Chicago

PreSUSY 2021. August 20, 2021

We have solid evidence for dark matter:





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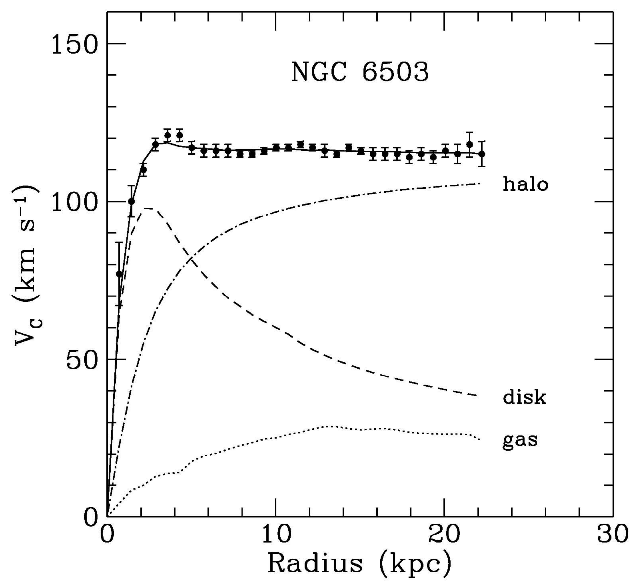

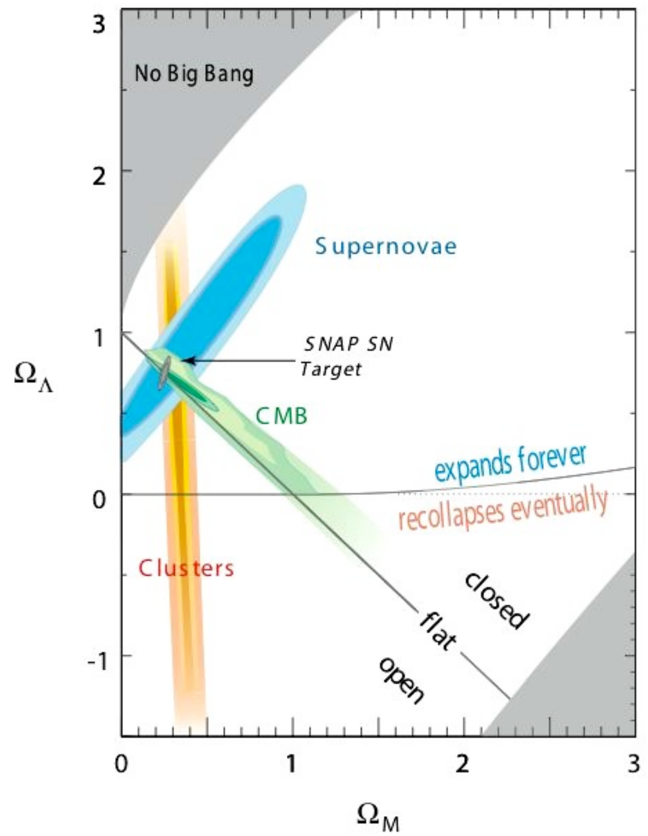

Our goal:
Understand the properties of dark matter.

This talk

- Going over our basic understanding and main scenarios of dark matter.
- Focusing on theoretical aspects of dark matter models
- It is a HUGE subject.
  - This talk is only an overview and glossing over many details.

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- Stable.
  - If it decays, lifetime much longer than the age of universe $\approx 10^{17}$ sec.

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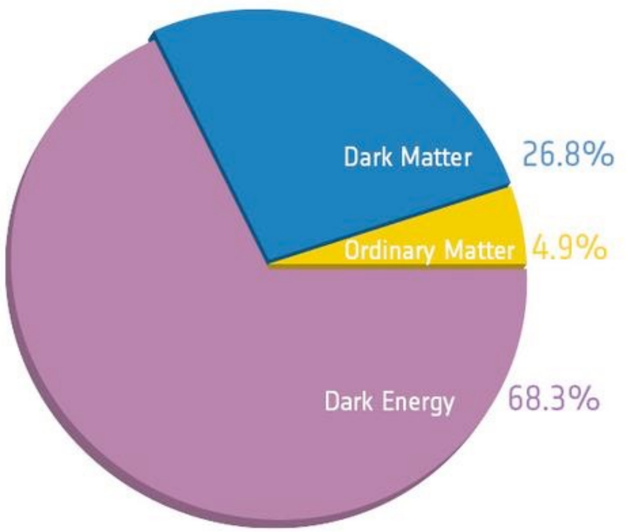
- Stable.
  - If it decays, lifetime much longer than the age of universe ≈ 10^{17} sec.
- Dark. Does not emit/absorb/reflect light.
  - Does not have electric charge.
- Produced in the early universe with the right amount. Right "relic abundance"



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- Seed structures in the universe.

We begin with quantum fluctuations in early universe

$z \simeq 10^7$, $T \sim 2$ keV

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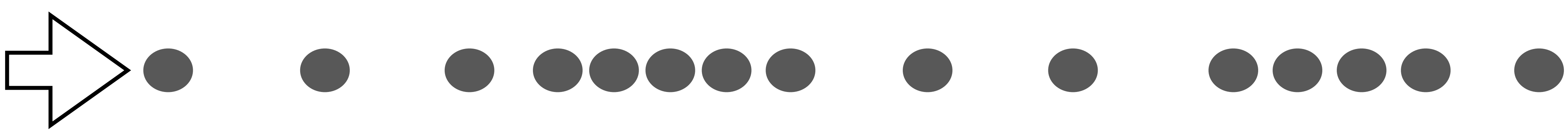

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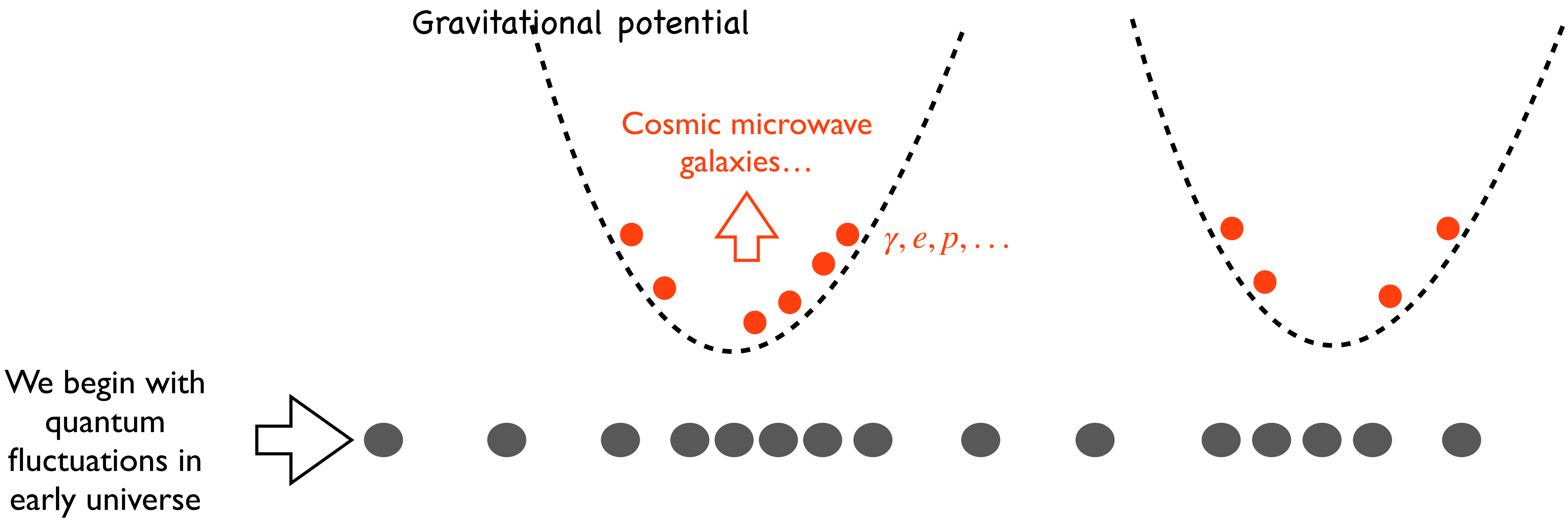

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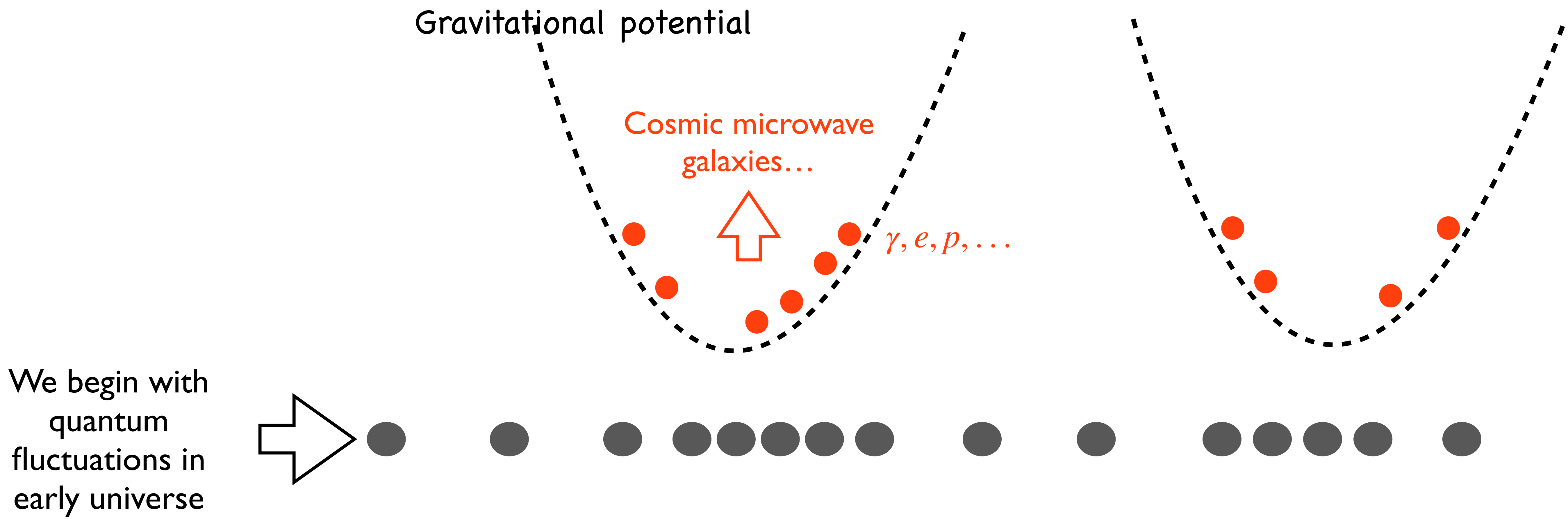
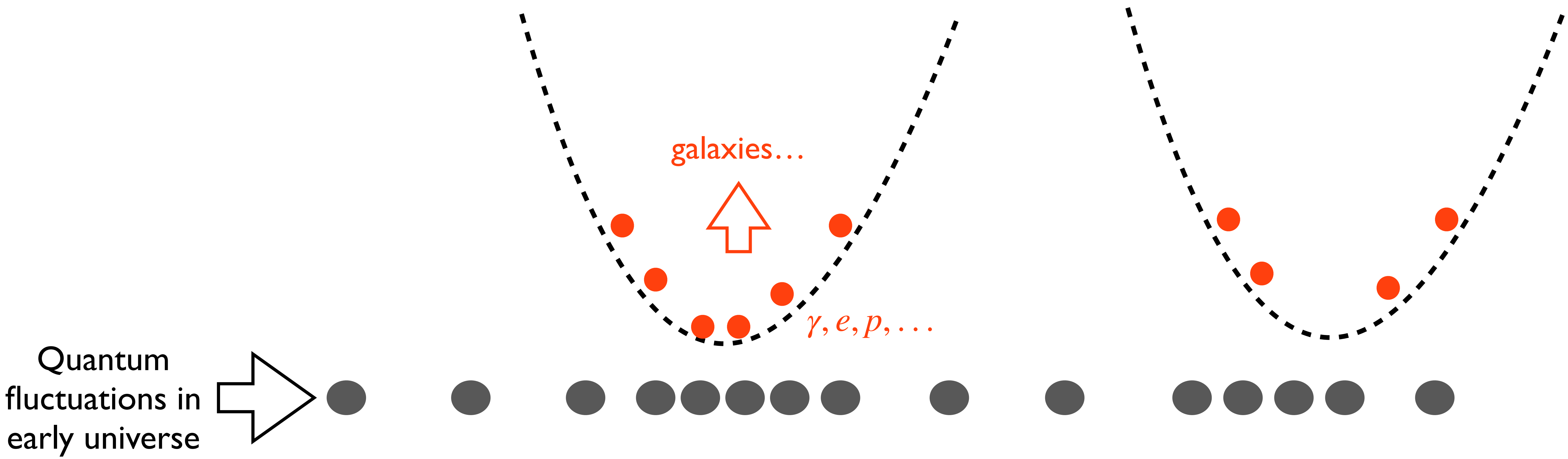


Dark matter follow these inhomogeneities

Dark matter needs to be “primordial”, be there in early universe.

$z \simeq 10^7,\ T \sim 2$ keV

What do we know about dark matter



- "Collisionless". No long range interaction, except gravity.
- Cold. Non-relativistic: kinetic energy $\ll$ mass

Mass of dark matter

Mass of dark matter

DM needs to seed structures

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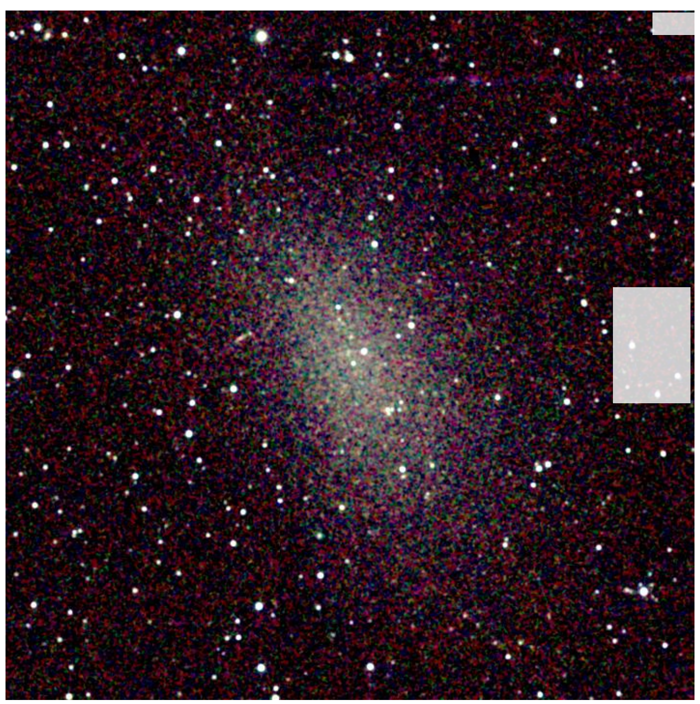
DM needs to seed structures



NGC147

Smallest structure DM seeded:
Dwarf spheroidal galaxy, size ≈ 1 kpc (3000 lyr)

Dark matter particle wave packet must be smaller,
Lightest ⇔ largest de Broglie wave length

$$\lambda_{\rm dB} = \frac{2\pi}{m_{\rm DM} v} \approx 0.4 \text{ kpc} \left(\frac{10^{-22}\text{eV}}{m_{\rm DM}} \right)$$

"Fuzzy dark matter"

Mass of dark matter

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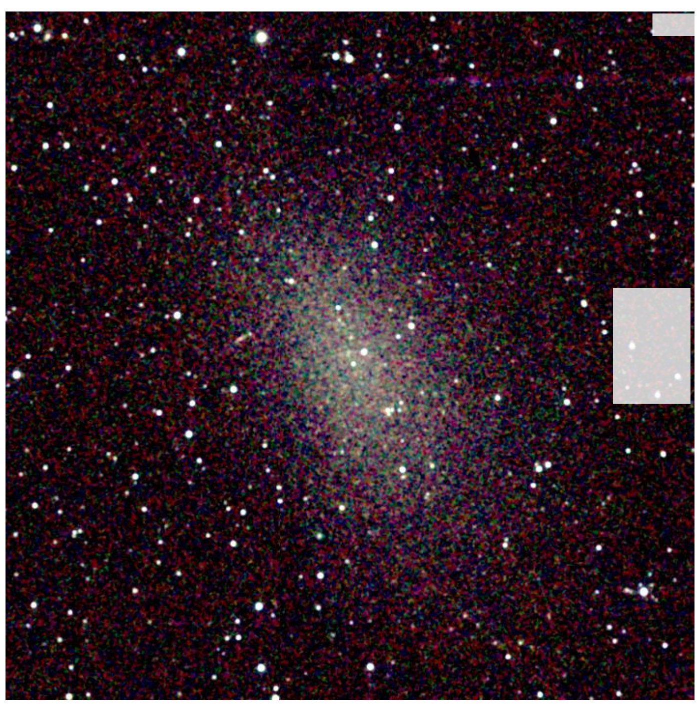

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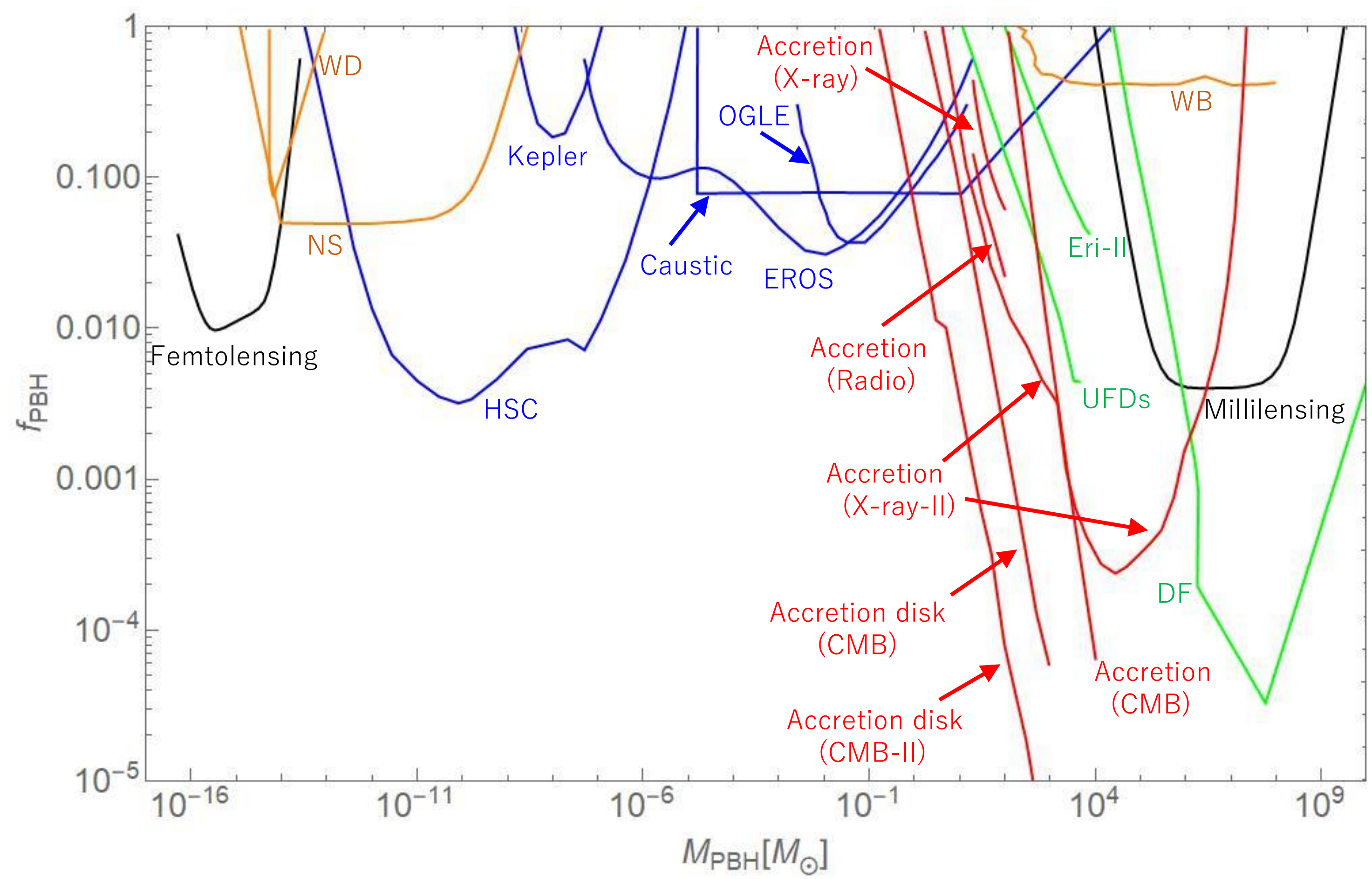
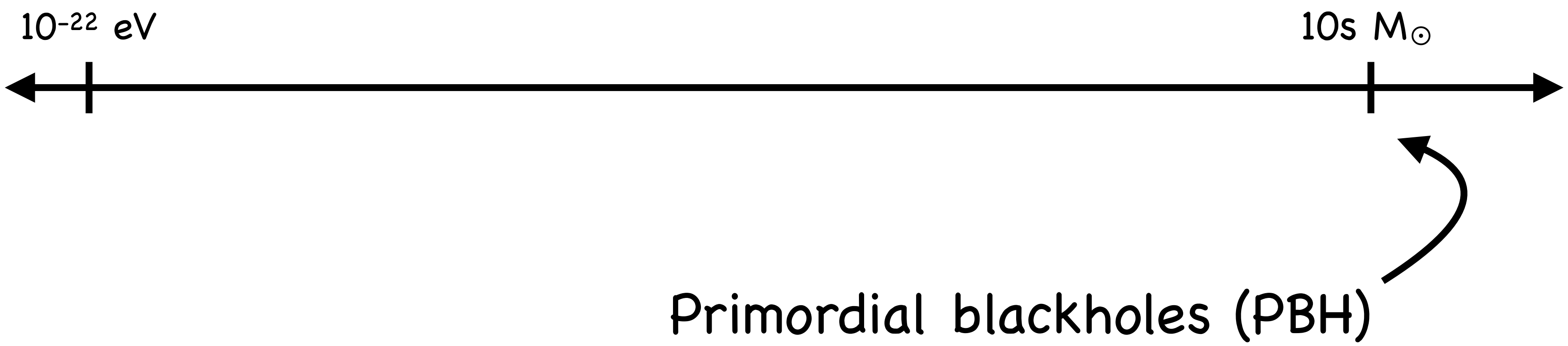
Very heavy BH accrete matter, too much ionizing radiation, CMB constraints $M_{PBH} < 10s\ M_\odot$

Blackhole lighter than $10^{-17}\ M_\odot$ will evaporate in the age of universe, not dark.

Other searches...



Mass of dark matter



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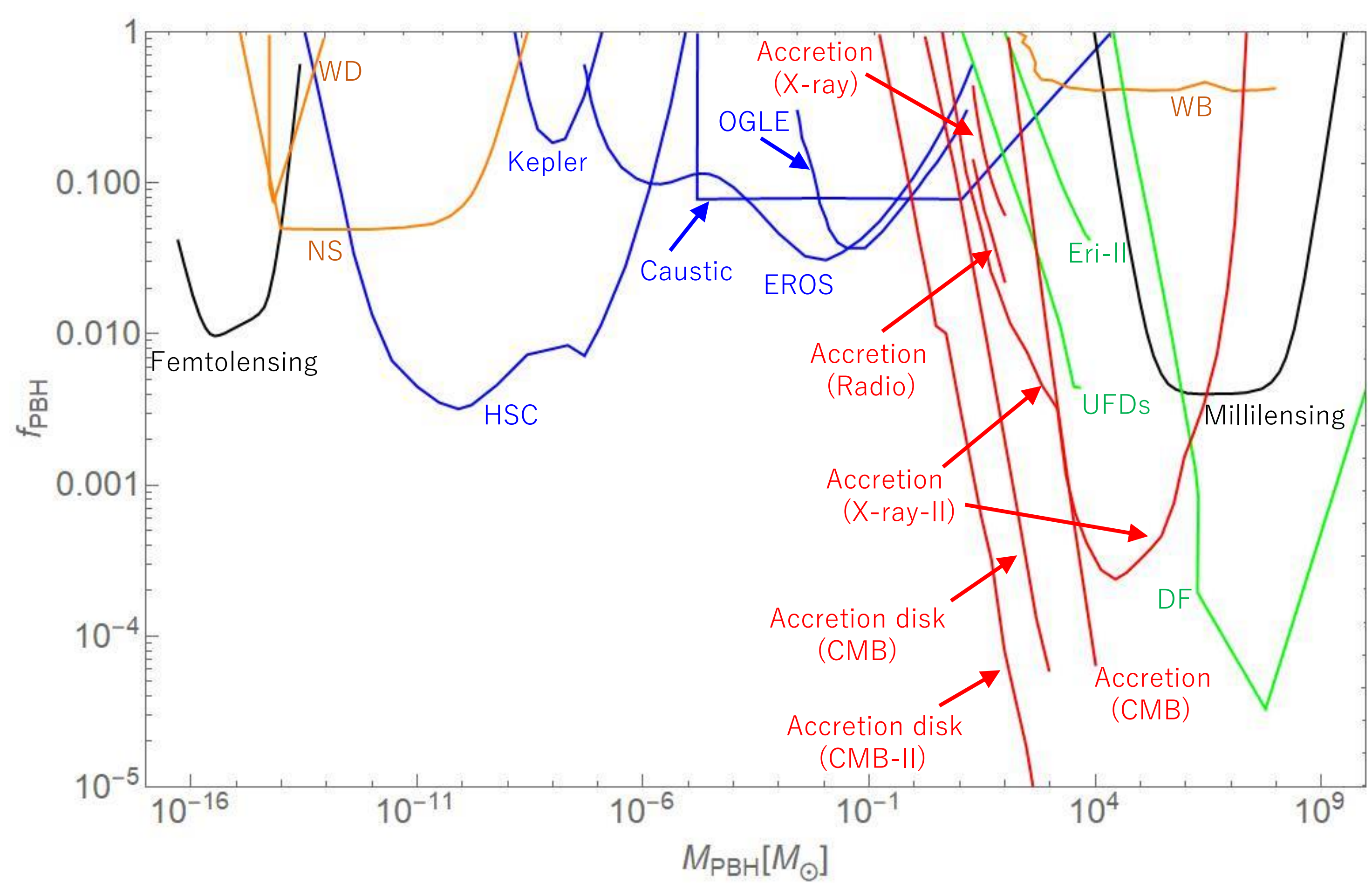
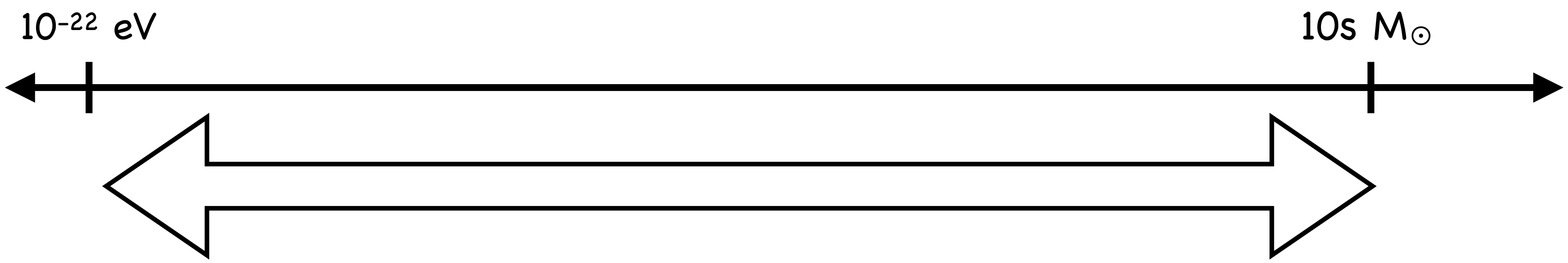


Mass of dark matter

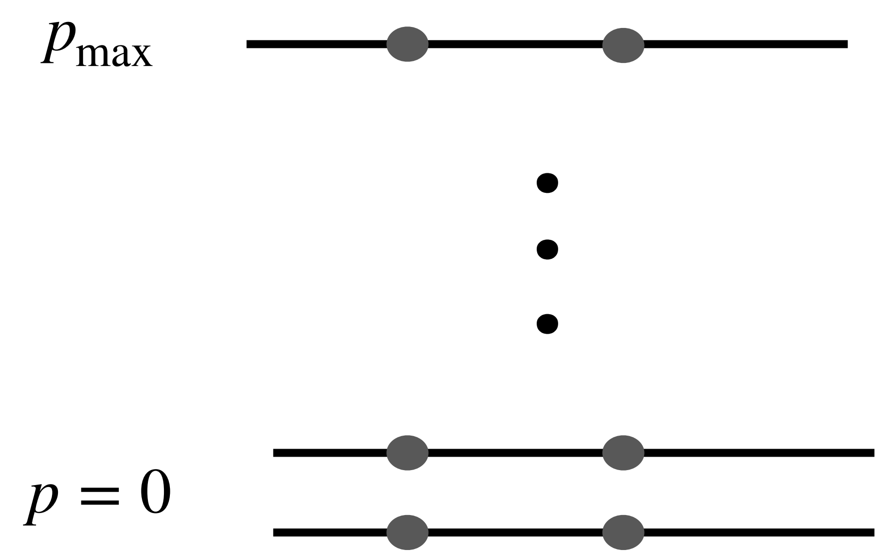


> 80 order of magnitudes!

What else can we say?

Mass of dark matter

Pauli exclusion principle



Fermionic dark matter

Mass of dark matter

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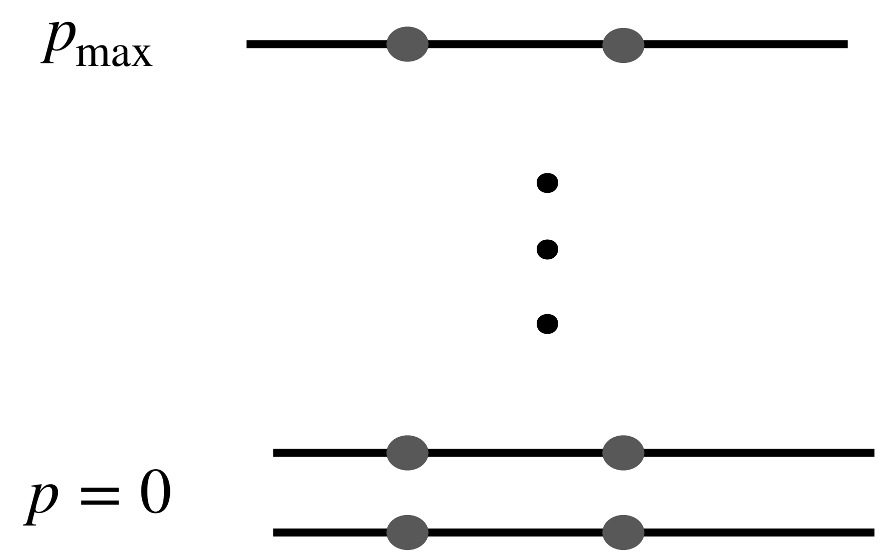


Fermionic dark matter

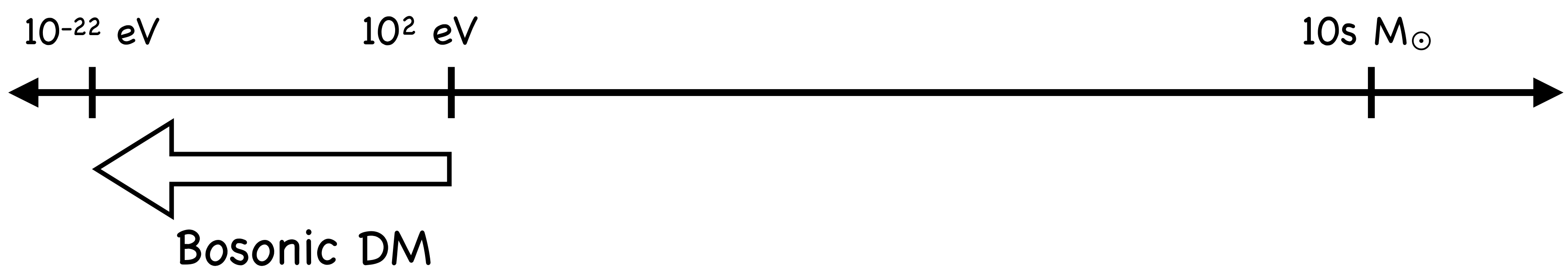
Since dark matter forms local bound structures such as galaxies

$$p_{\max} \simeq m_{\rm DM} \times v_{\rm esc}$$

$v_{\rm esc}$: escape velocity

$$m_{\rm DM} > 10s \text{ eV}$$

Mass of dark matter



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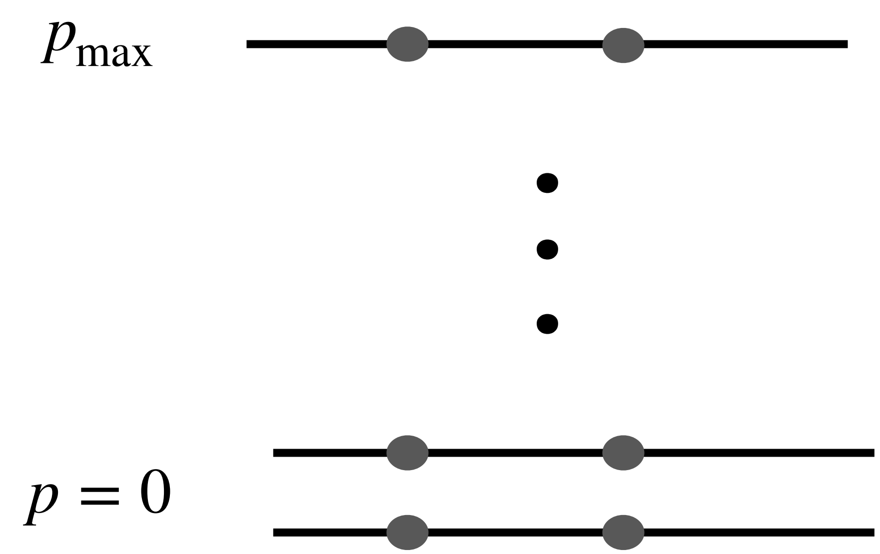

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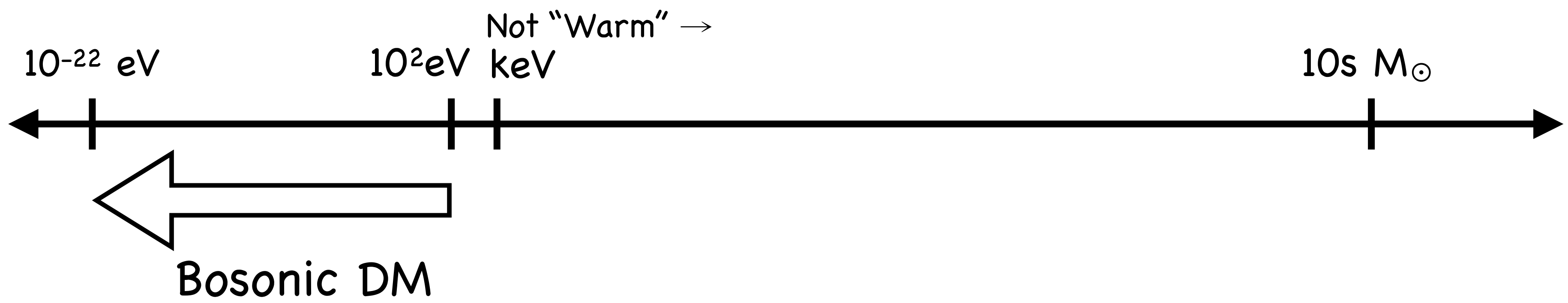
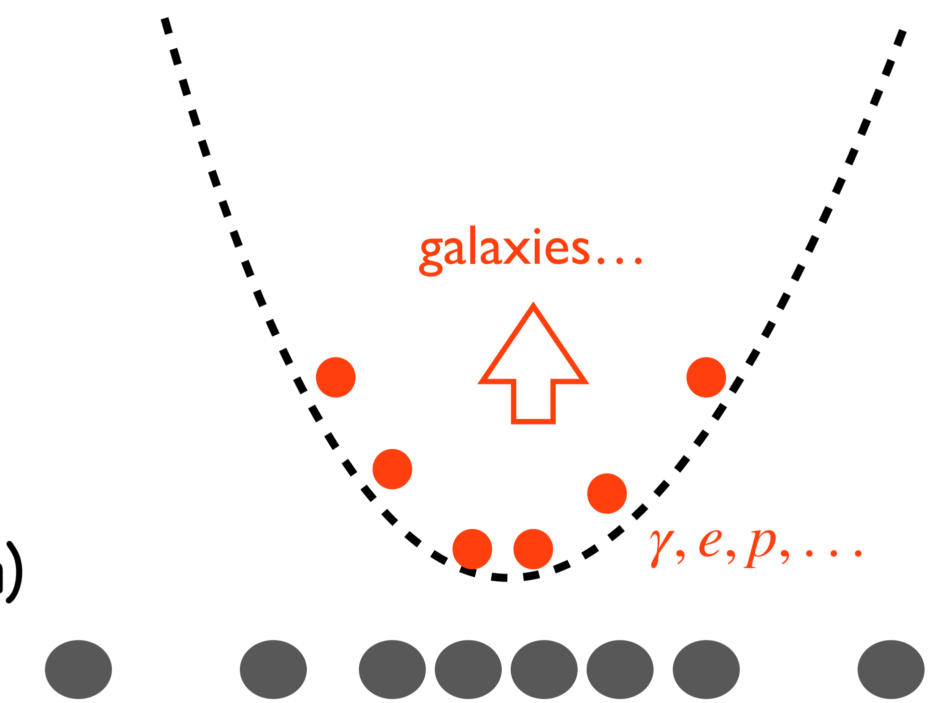


Warm dark matter limit:

Dark matter needs to be cold (non-relativistic) for the smallest structure it can seed.

For dark matter particle (in thermal equilibrium) m_{DM} > keV (10^3 eV)

Could be ways of getting around this if DM is not in thermal eq., or there is some ways they can cool.



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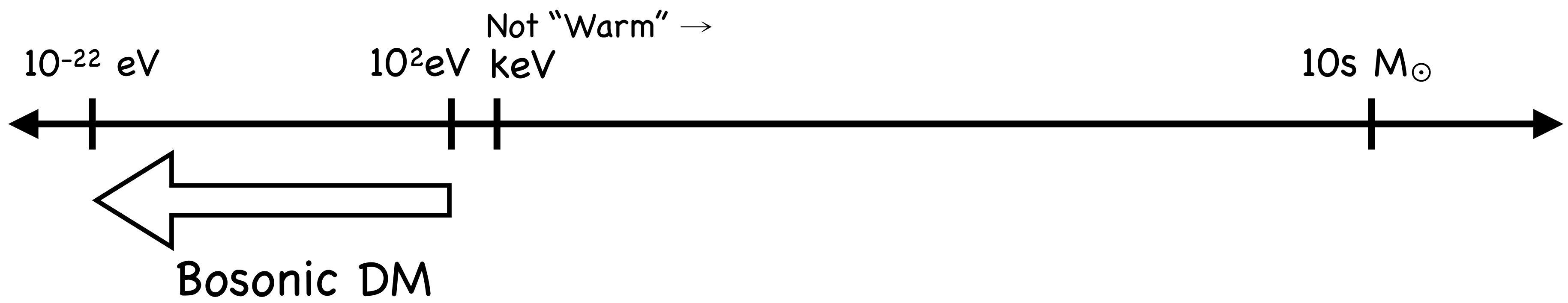

How do we guide our searches? We need theories (stories).

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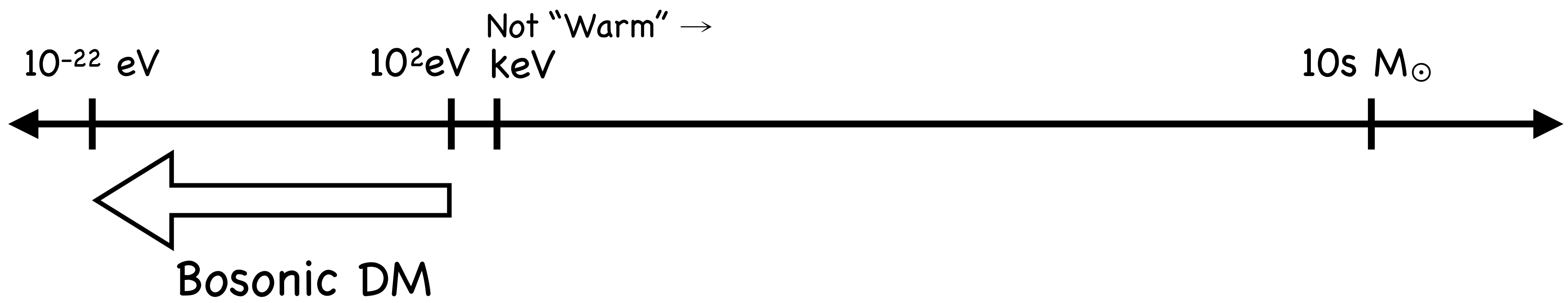

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2) How is dark matter produced in the early universe?

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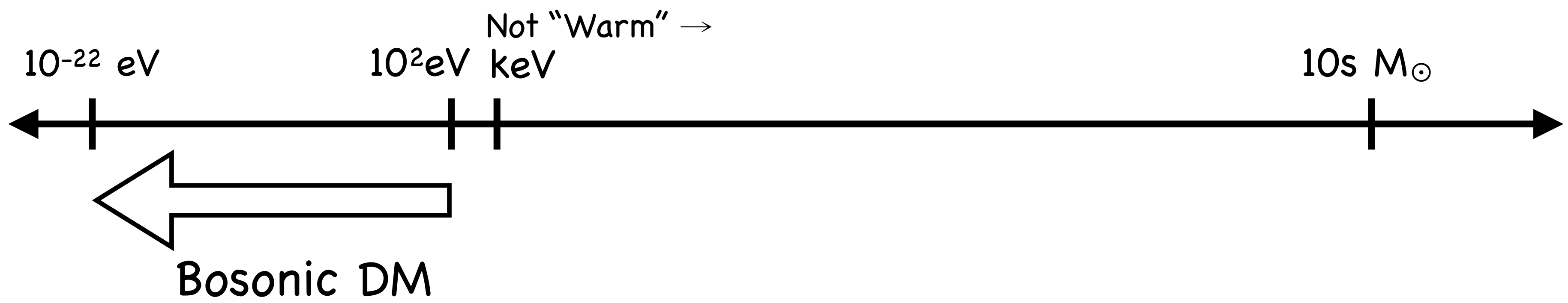

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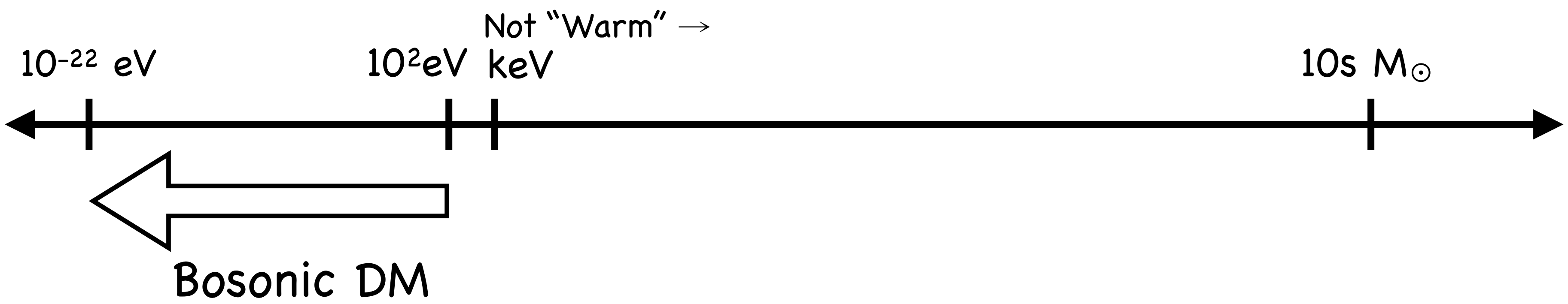

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Vast number of models, only a few good theories.

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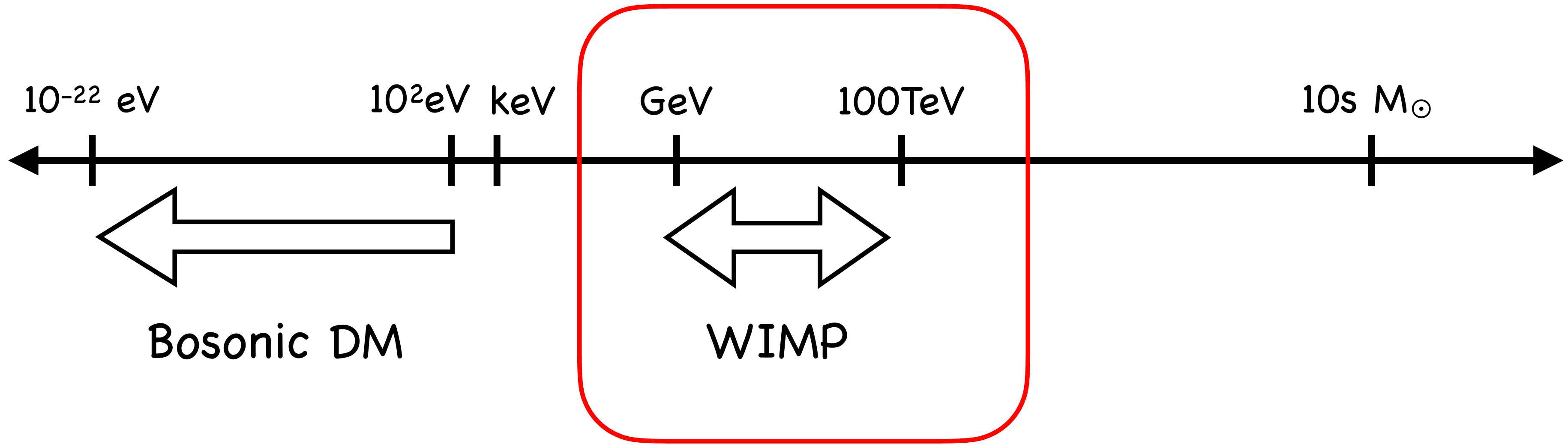

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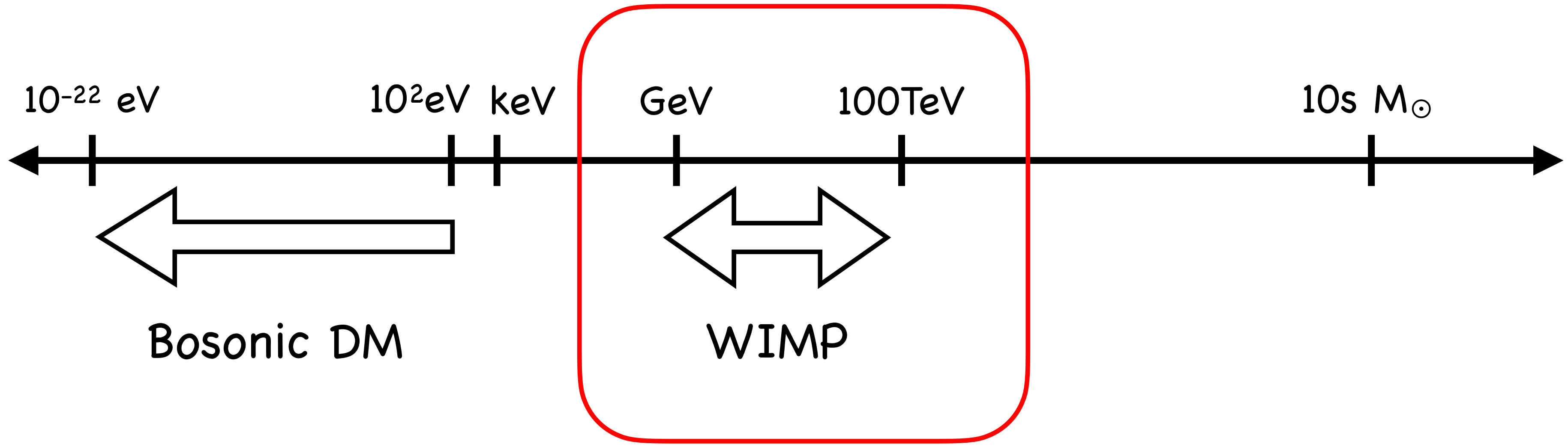

A lamppost.
A tiny window in the full mass range.
But, it is a good lamppost.

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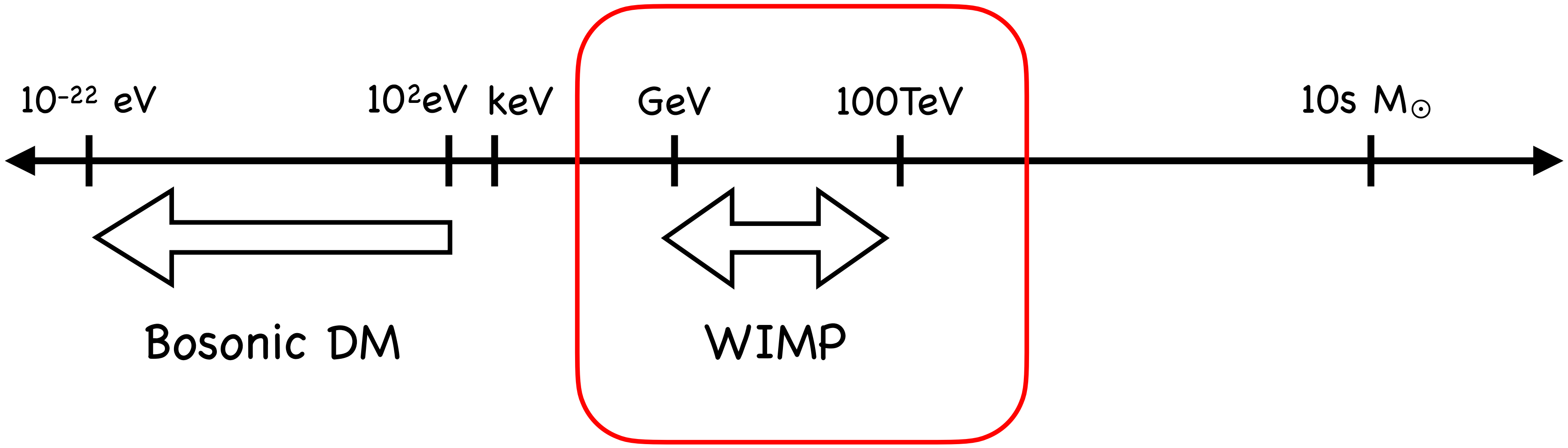

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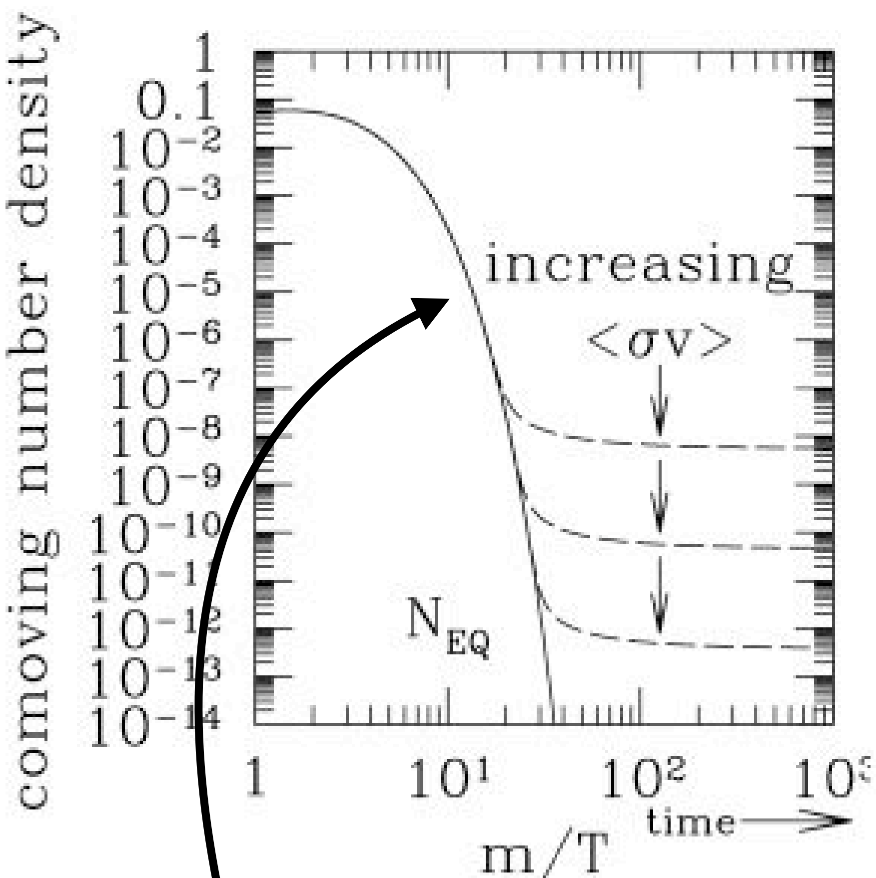
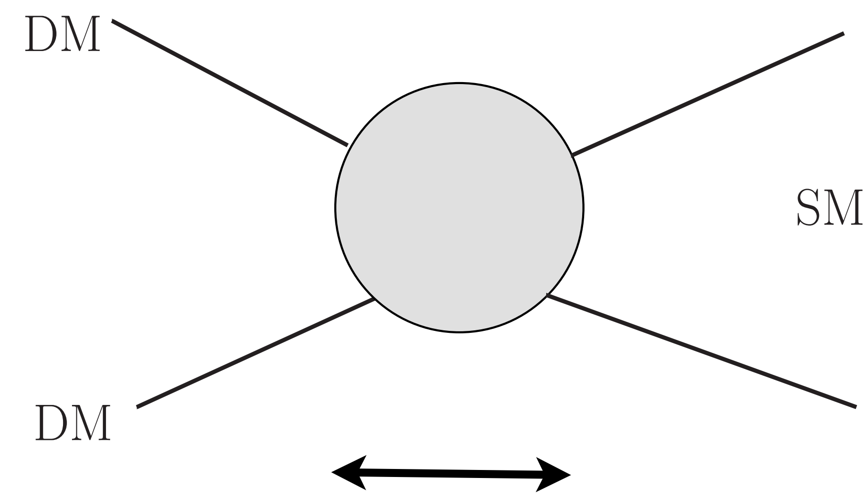
WIMP (weakly interacting massive particle)



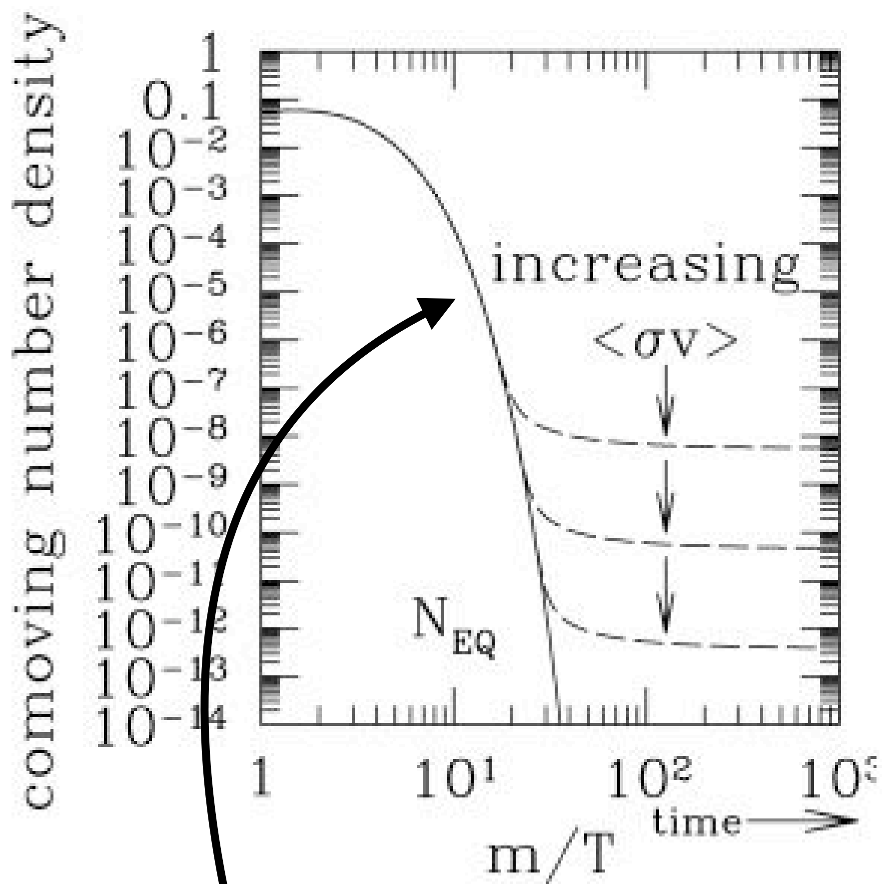


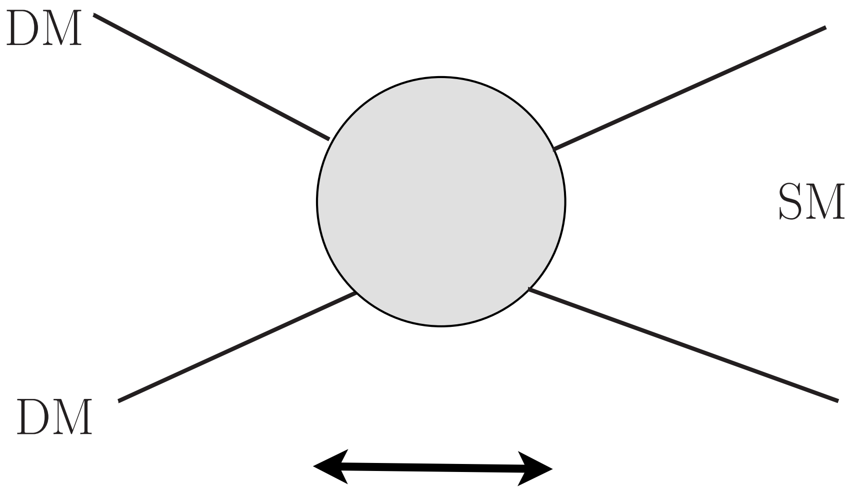
$$\langle \sigma v \rangle \sim \frac{g_{\rm D}^4}{m_{\rm DM}^2}$$

Dark matter in thermal equilibrium with the known (Standard Model) particles in the early universe.

Interaction rate faster than the expansion of the universe

WIMP (weakly interacting massive particle)



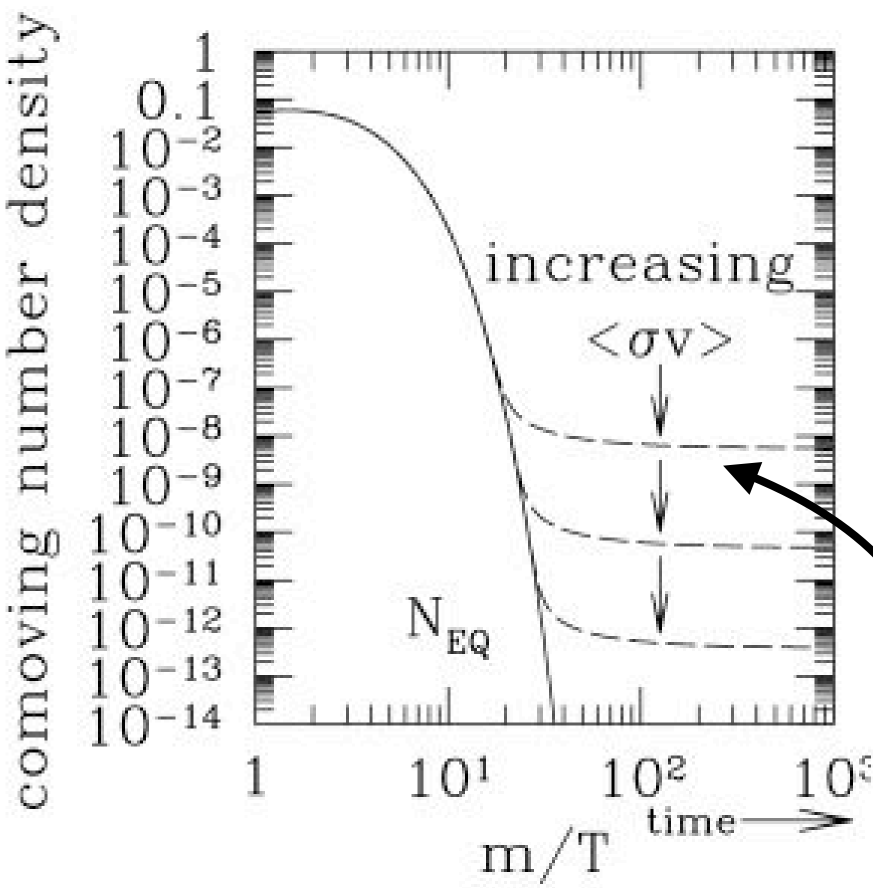


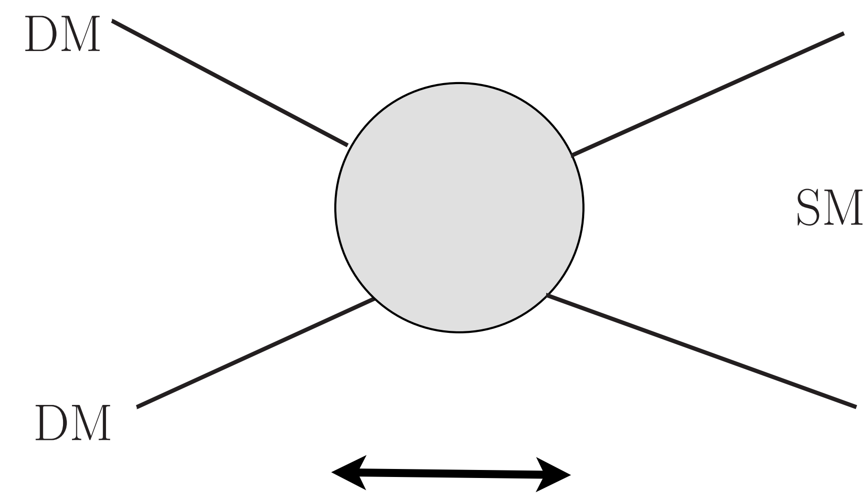
$$\langle \sigma v \rangle \sim \frac{g_D^4}{m_{DM}^2}$$

Dark matter in thermal equilibrium with the known (Standard Model) particles in the early universe.

Dark matter number density predicted by thermal eq: N_{EQ}

WIMP (weakly interacting massive particle)



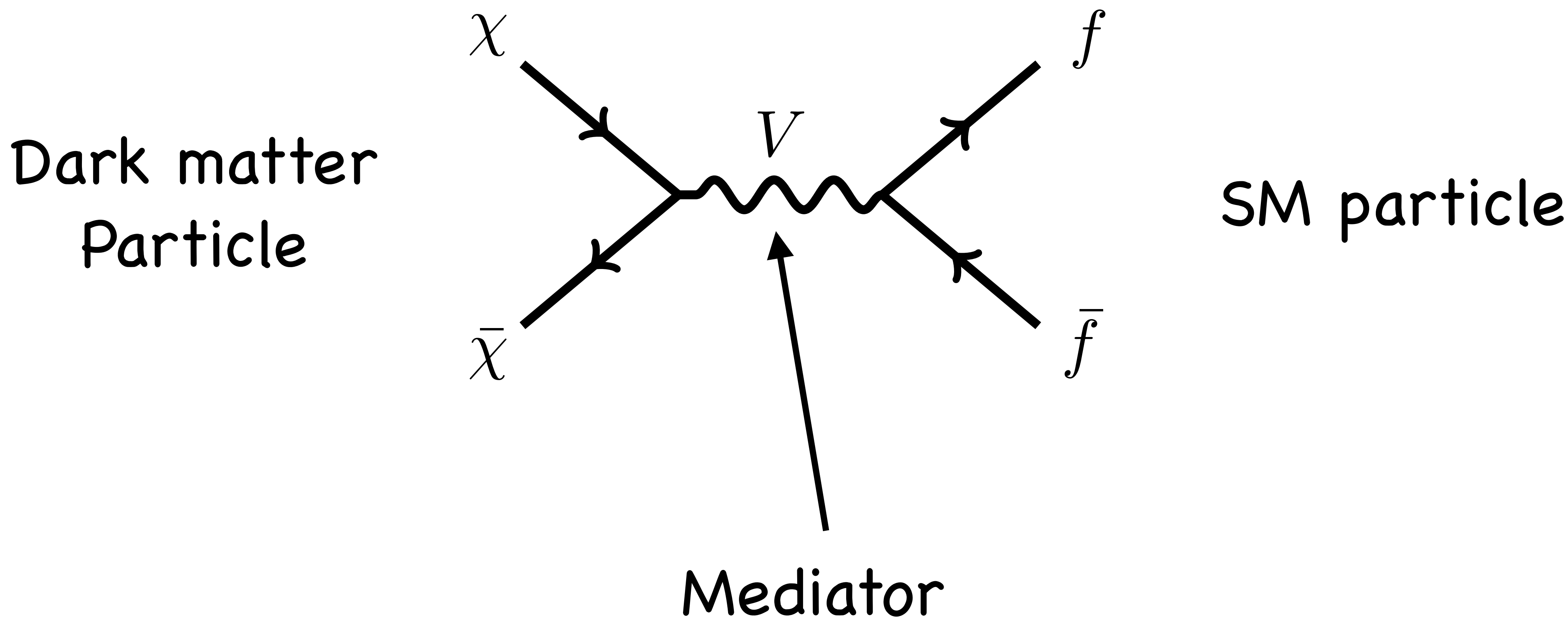


$$\langle \sigma v \rangle \sim \frac{g_D^4}{m_{DM}^2}$$

As universe expands, dark matter become rare. The DM-SM interaction rate can't keep up. DM drops out thermal eq.

Dark matter density become fixed, "Freeze-out"

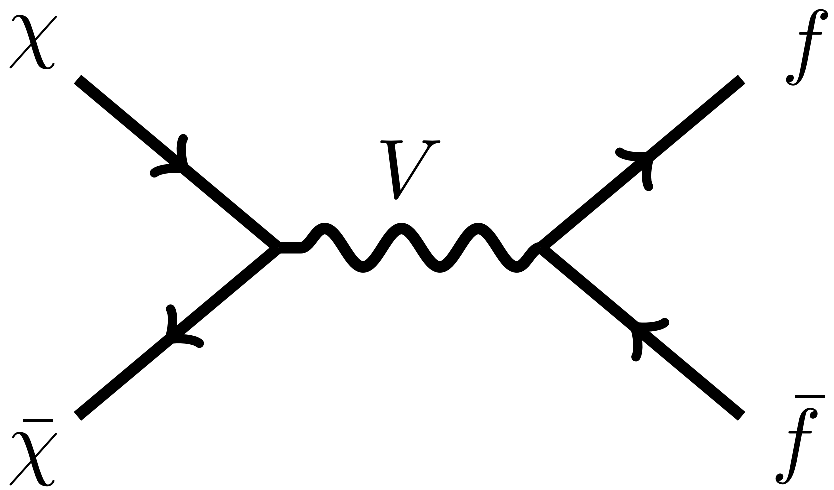
A simple picture of interaction



To get the correct relic abundance:

$$\langle \sigma v \rangle \approx 2 \times 10^{-26} \mathrm{cm}^3/\mathrm{s}$$

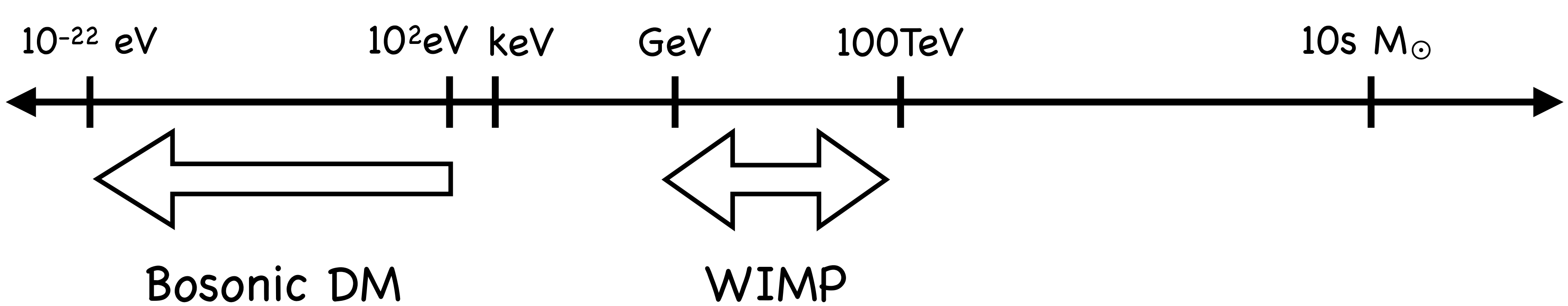
Two limits



$$m_{\mathrm{DM}} > m_V, \quad \sigma v \sim \frac{g^4}{4\pi} \frac{1}{m_{\mathrm{DM}}^2}$$

g: coupling

Limit on coupling: $g < 4\pi \Rightarrow M_{DM} < 100\text{s TeV}$



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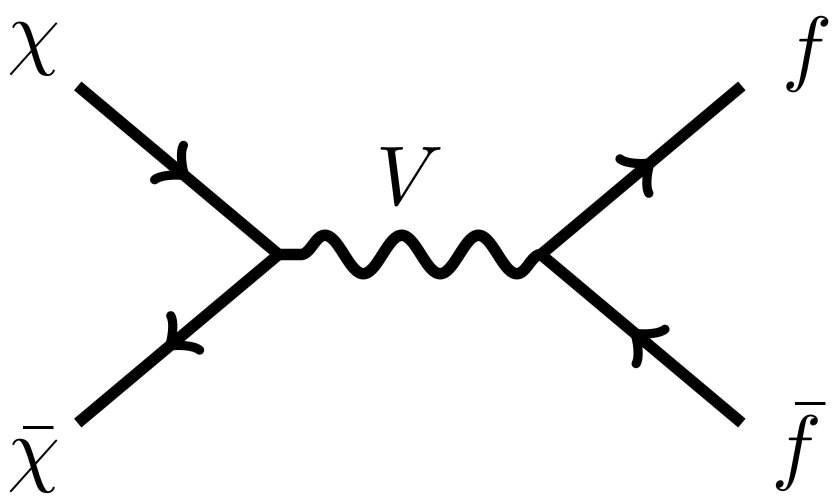
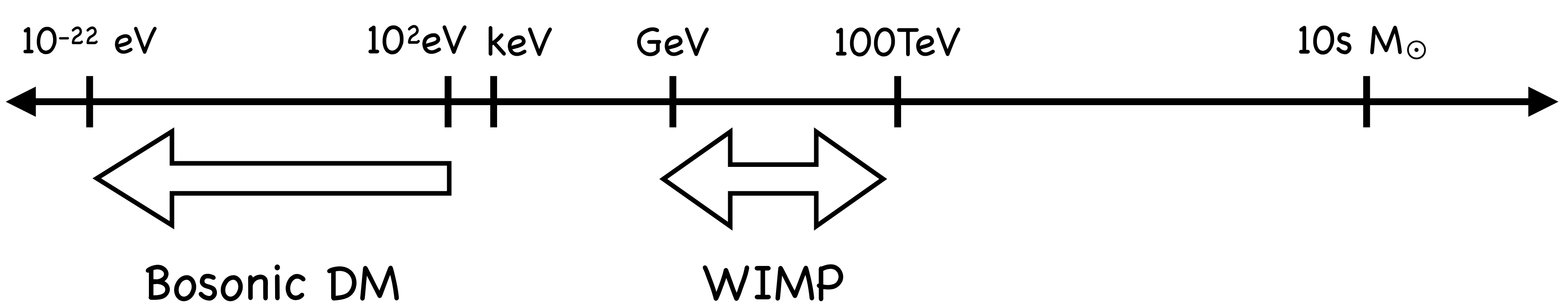


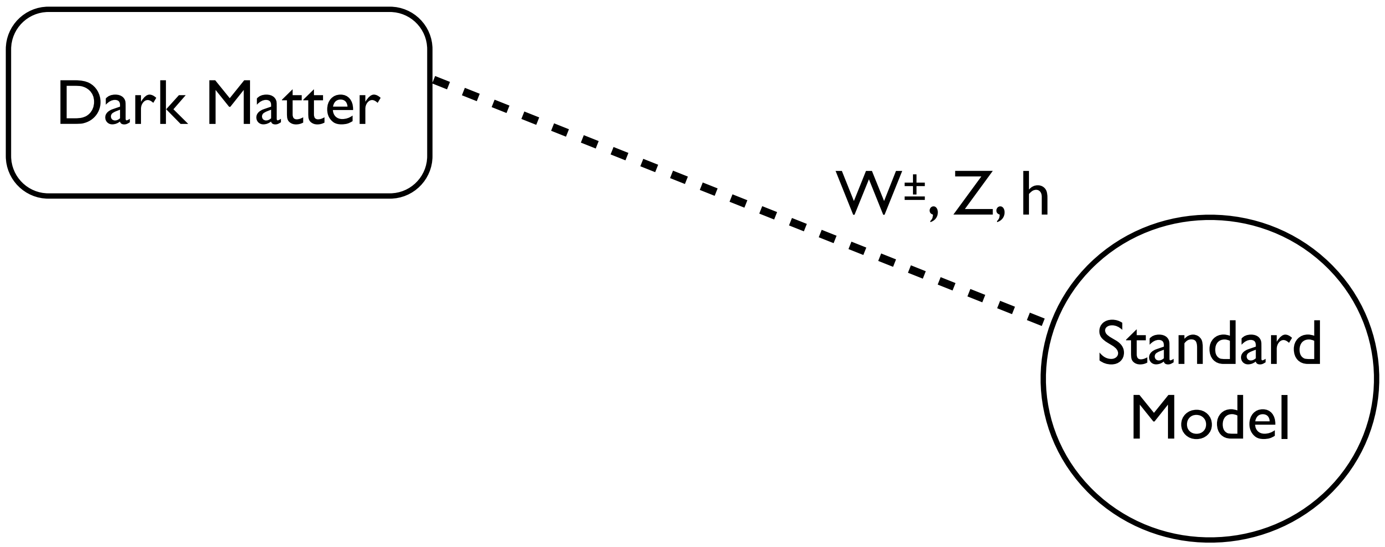
$$m_{\rm DM} < m_V, \quad \sigma v \sim \frac{g^4}{4\pi}\frac{m_{\rm DM}^2}{m_V^4}$$

g: coupling

$m_V \approx 10^2$ GeV $\Rightarrow$ m_{DM} > GeV



Simple WIMP model

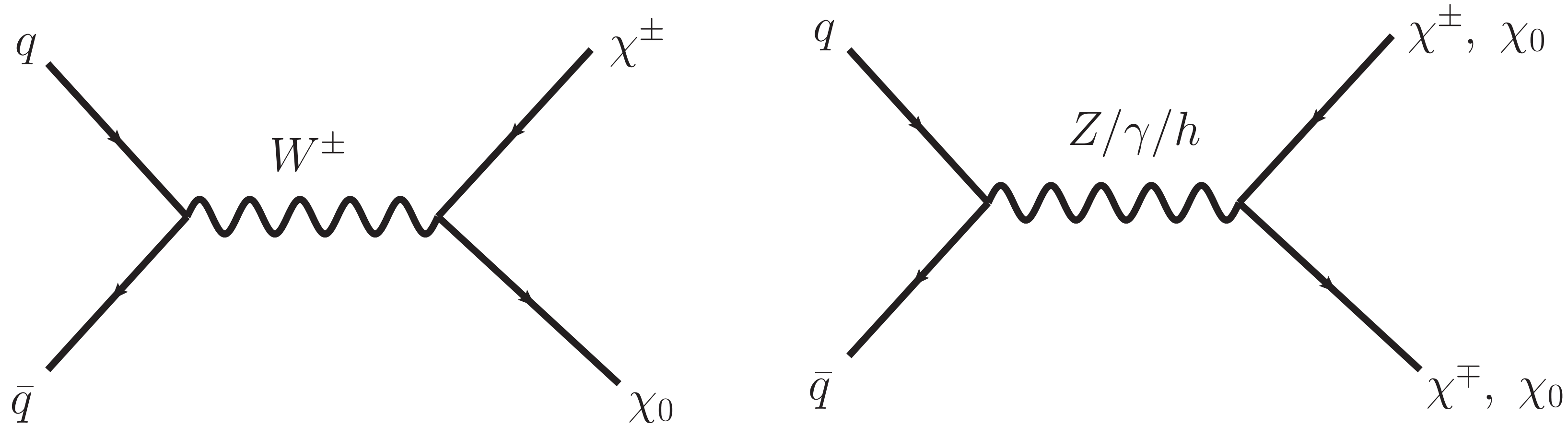


Mediated by a known interaction:

The weak interaction in the Standard Model

Mediator mass: 10^2 GeV

DM part of a EW multiplet



- Simplicity: there is no additional new mediator.
  - Mediated by W/Z/h.

- In SUSY, there are two such examples
  - Higgsino: doublet. Wino: triplet.

More generally.

n: $SU(2)_L$, Y: Hyper charge

Model (color, n, Y)	
(1,2,1/2)	Dirac
(1,3,0)	Majorana
(1,3,ϵ)	Dirac
(1,5,0)	Majorana
(1,5,ϵ)	Dirac
(1,7,0)	Majorana
(1,7,ϵ)	Dirac

Fermionic dark matter:

Renormalizable interaction completely determined by SM gate interactions.

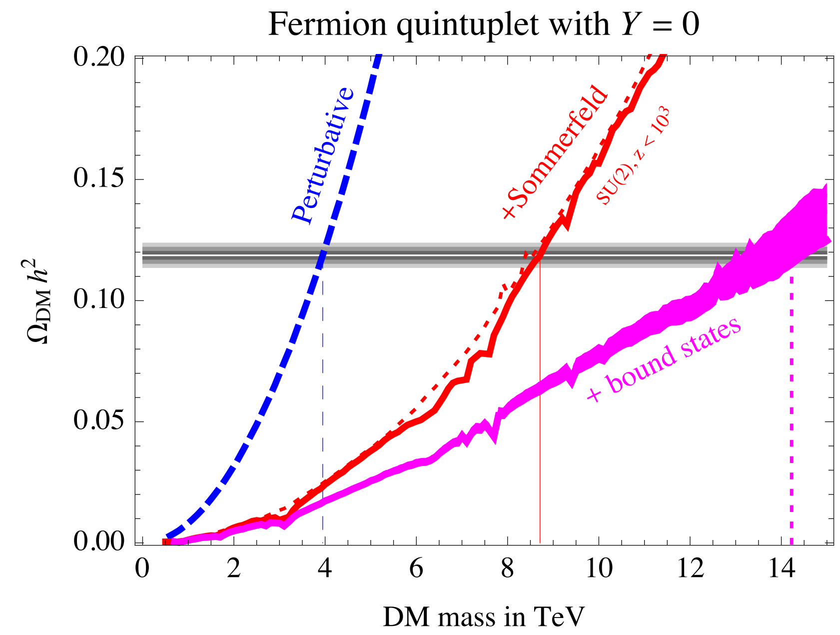
Very predictive: abundance only depends on the mass of the DM

M_{DM} with the right DM abundance = Thermal target

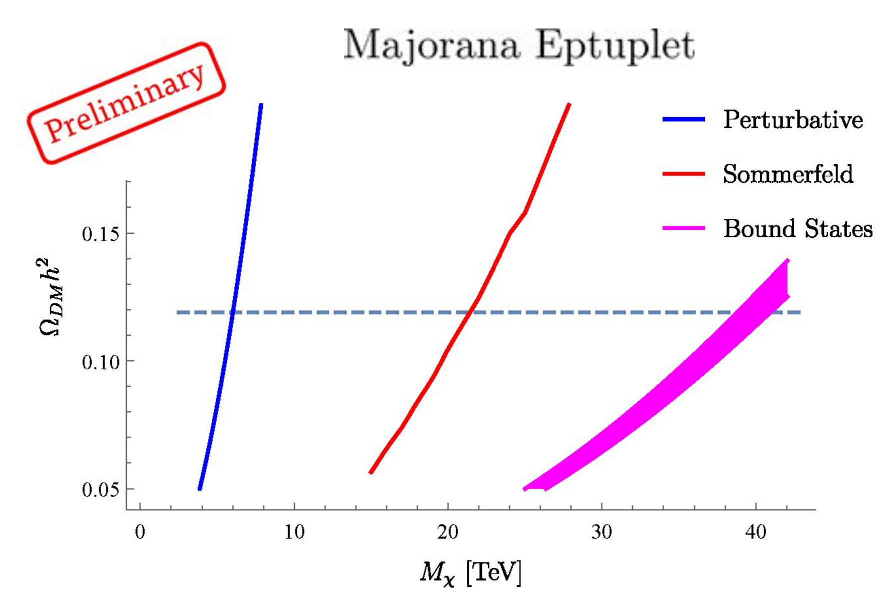
More generally.

n: $SU(2)_L$, Y: Hyper charge

Model (color, n, Y)		Therm. target
(1,2,1/2)	Dirac	1.1 TeV
(1,3,0)	Majorana	2.8 TeV
(1,3,ϵ)	Dirac	2.0 TeV
(1,5,0)	Majorana	14 TeV
(1,5,ϵ)	Dirac	6.6 TeV
(1,7,0)	Majorana	23 TeV $\rightarrow$ 40 TeV
(1,7,ϵ)	Dirac	16 TeV



Mitridate, Redi, Smirnov, Strumia, 1702.01141



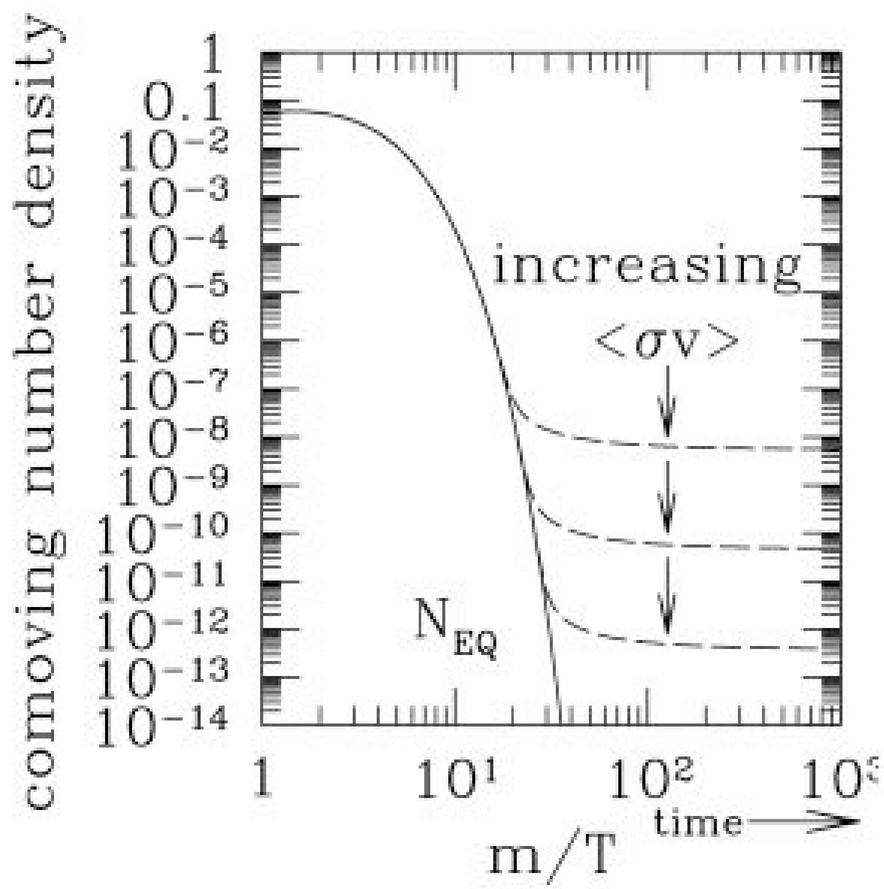
Bottaro, 2nd muon collider physics potential meeting

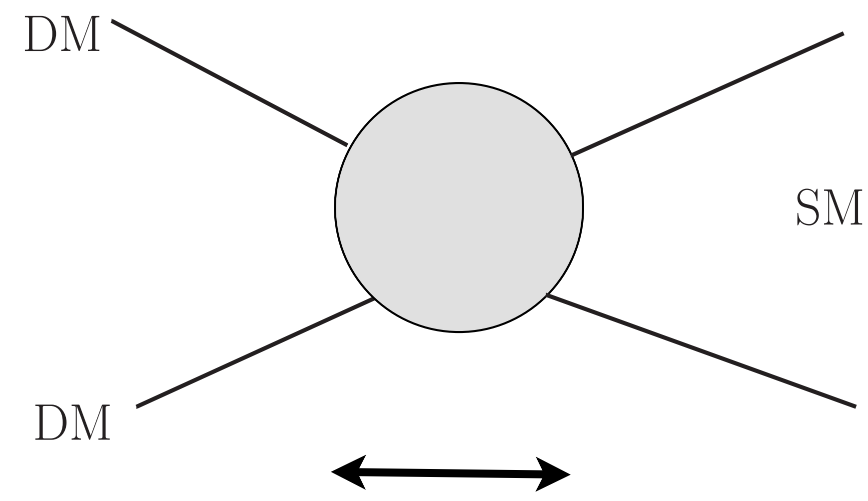
Reach up to thermal target

≈

complete coverage for WIMP candidate

WIMP "miracle"

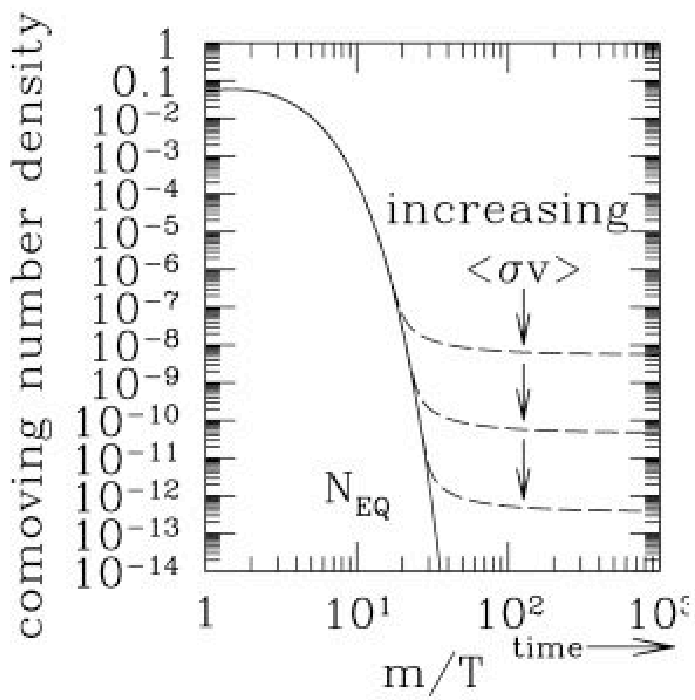


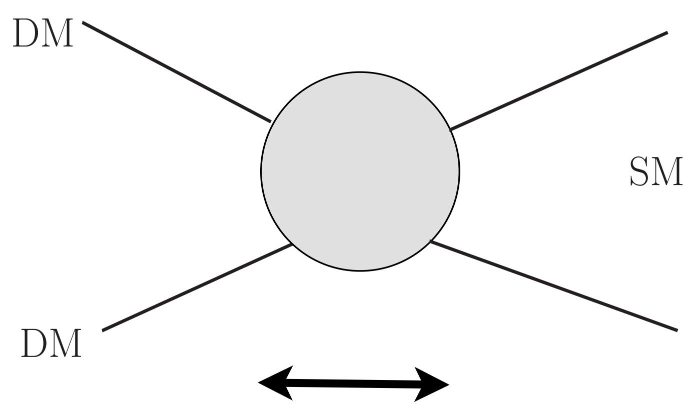


$$\langle \sigma v \rangle \sim \frac{g_D^4}{m_{DM}^2}$$

- If $g_D \sim 0.1$ $M_D \sim$ 10s GeV - TeV
  - We get the right relic abundance of dark matter.
- Coincide with our expectation for weak($\pm$) scale new physics!

Why is WIMP a good theory?





$$\langle \sigma v \rangle \sim \frac{g_{\mathrm{D}}^4}{m_{\mathrm{DM}}^2}$$

Reasonable:

Early universe (hot) is in thermal equilibrium. Don't need to know too much detail beyond (before) that.

Can be linked to other motivations for electroweak scale new physics.

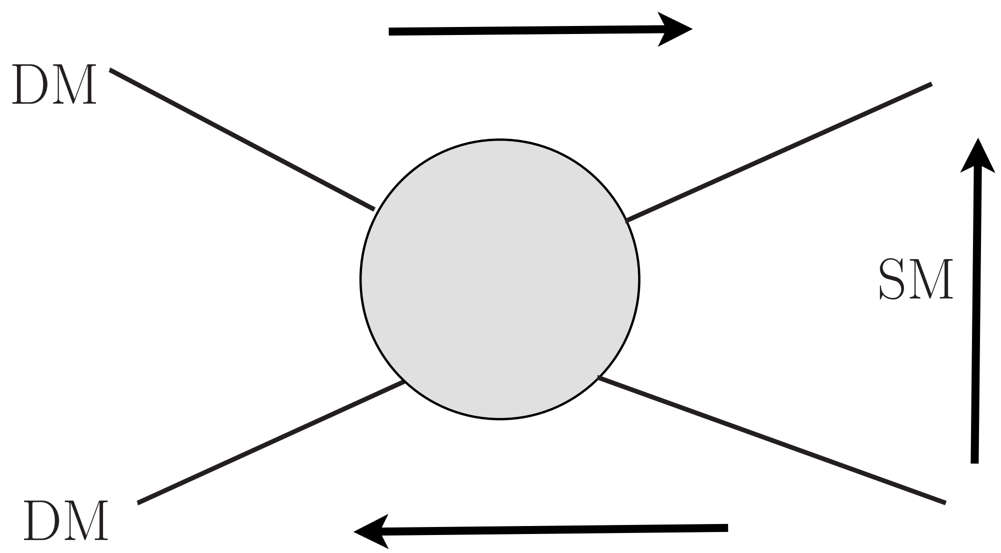
Present in many models: SUSY, extra dimension...

Testable:

With a sizable coupling to the known (SM) particle, WIMP can be searched in labs.

Looking around the lamppost

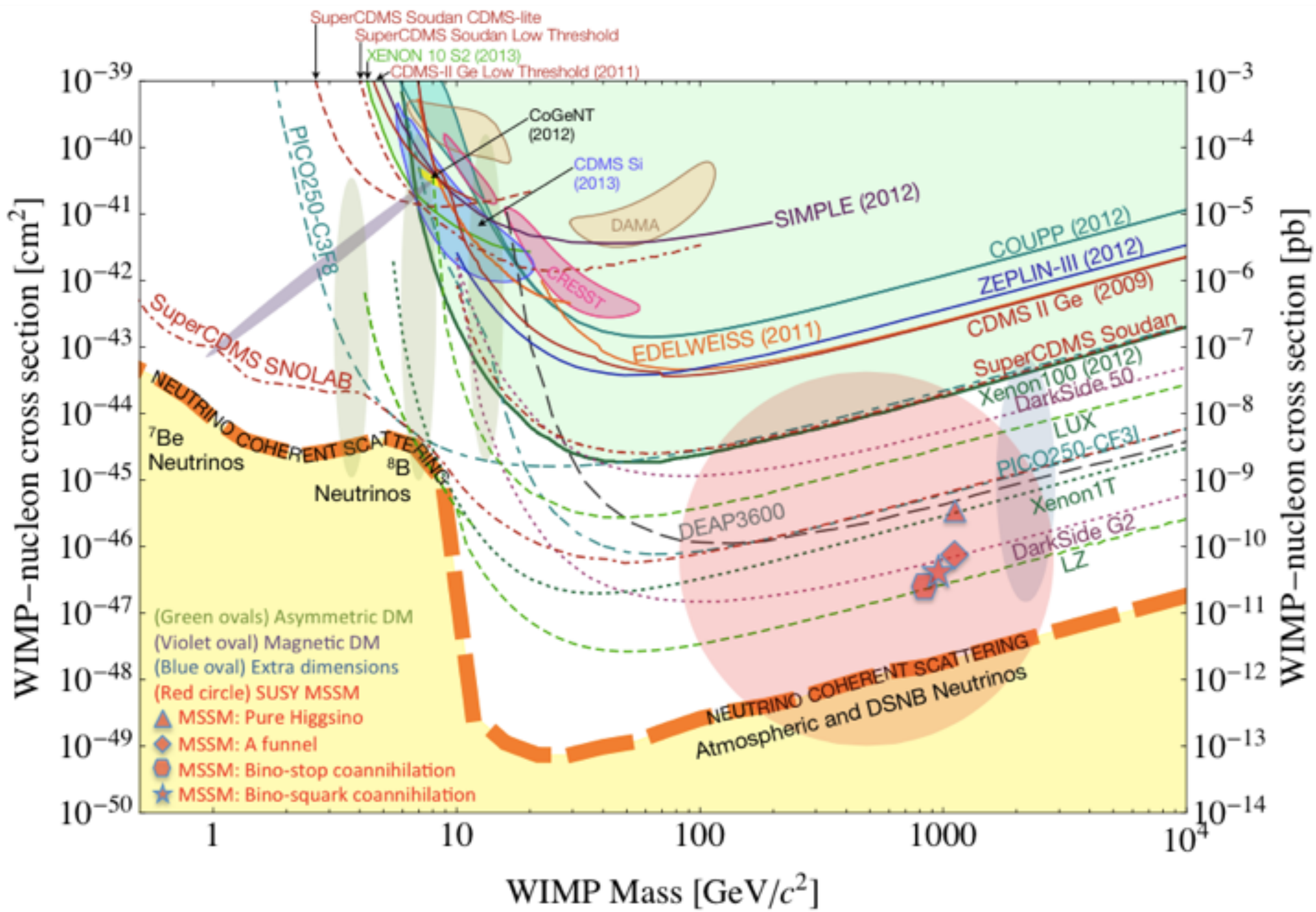
Indirect detection:
AMS2, PAMELA, Fermi-LAT
Cosmology



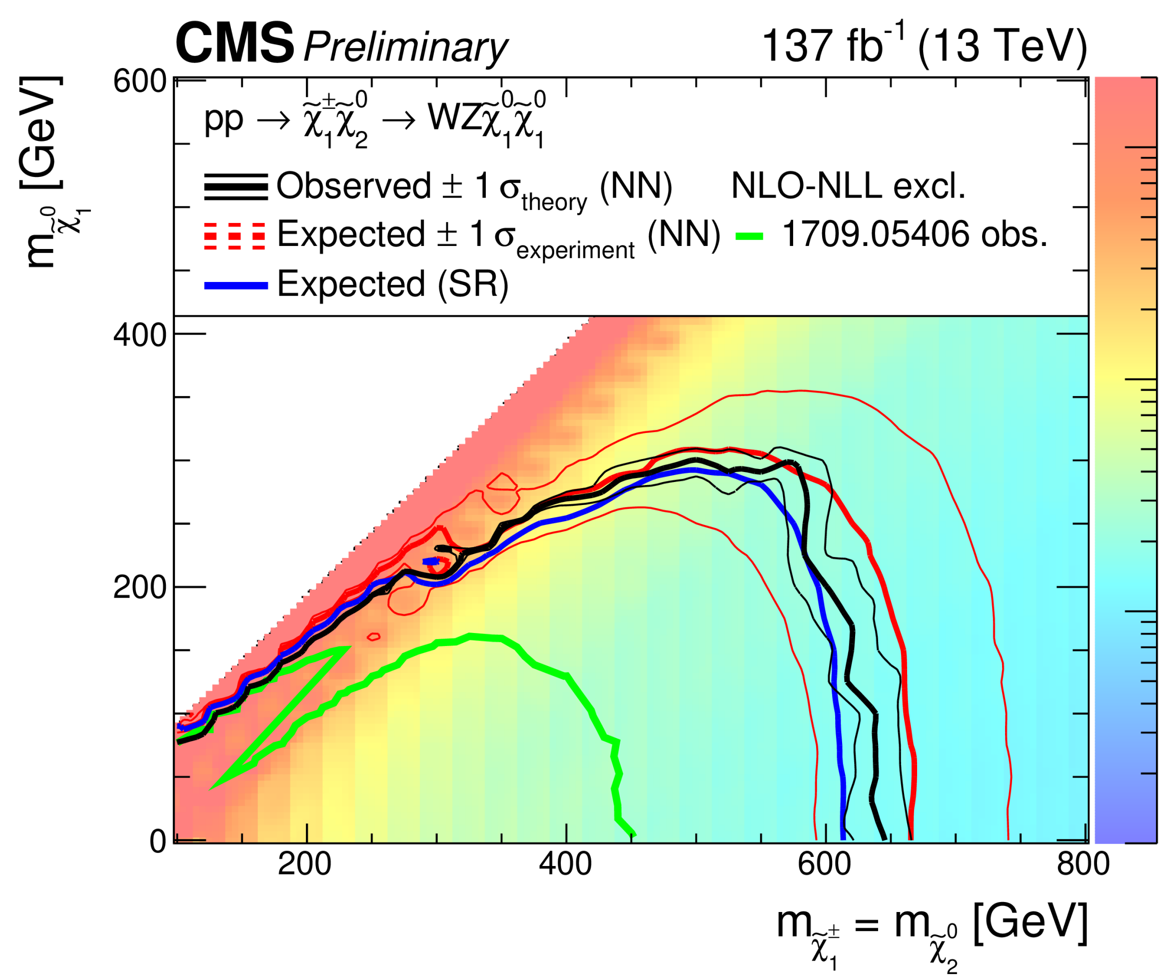
Direct detection:
CDMS
CoGeNT
COUPP
CRESST
DAMA
XENON
Lux
.....

Collider searches:
LHC, ...

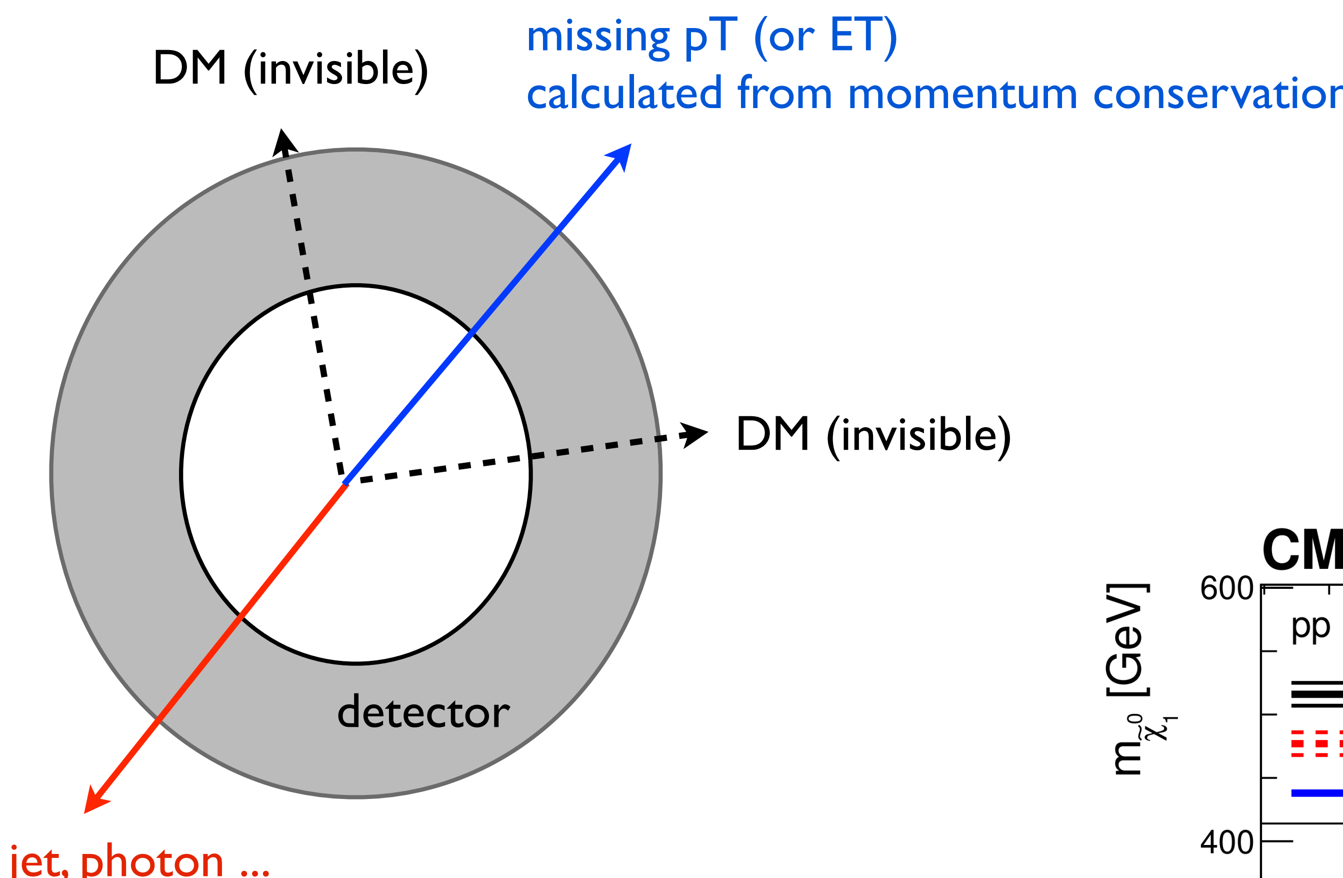
Direct detection

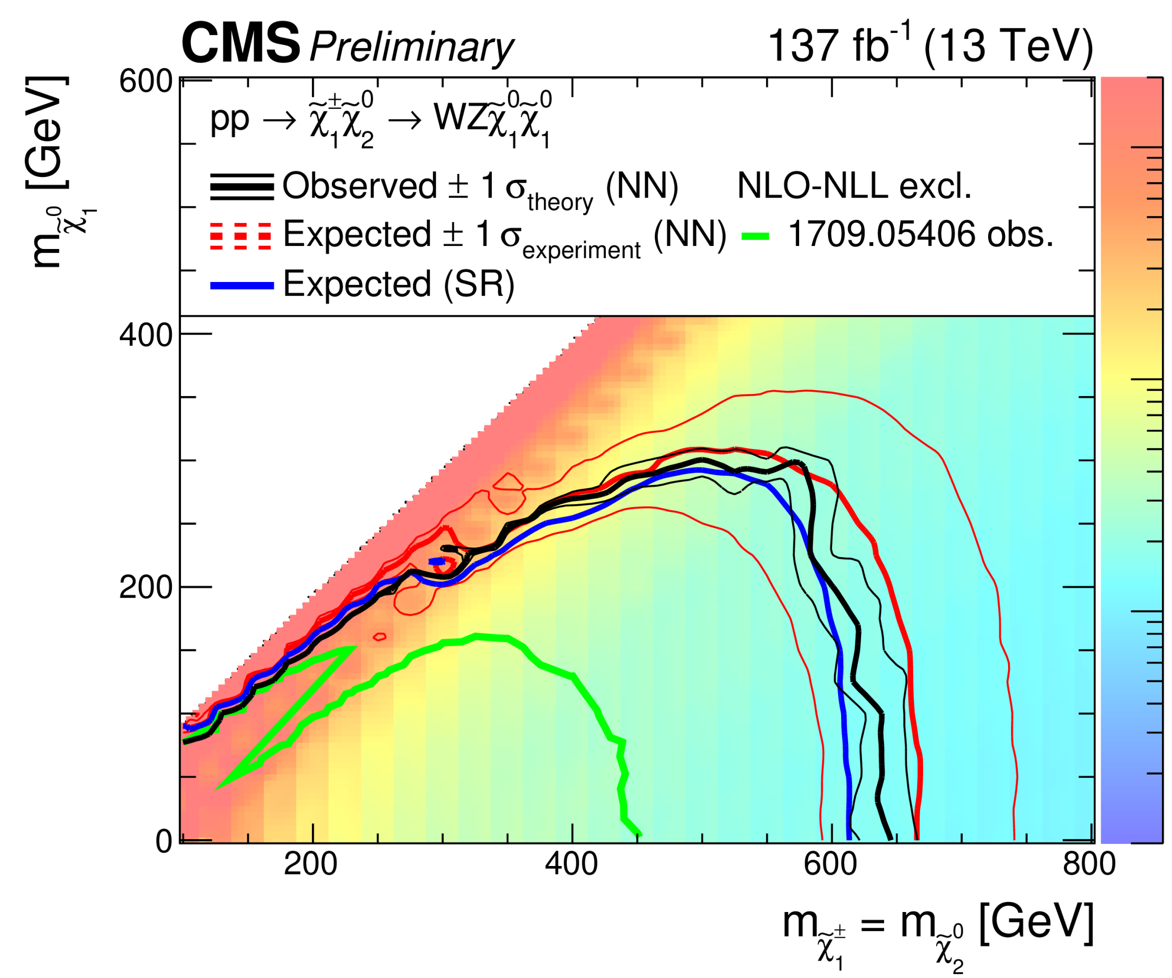


At colliders

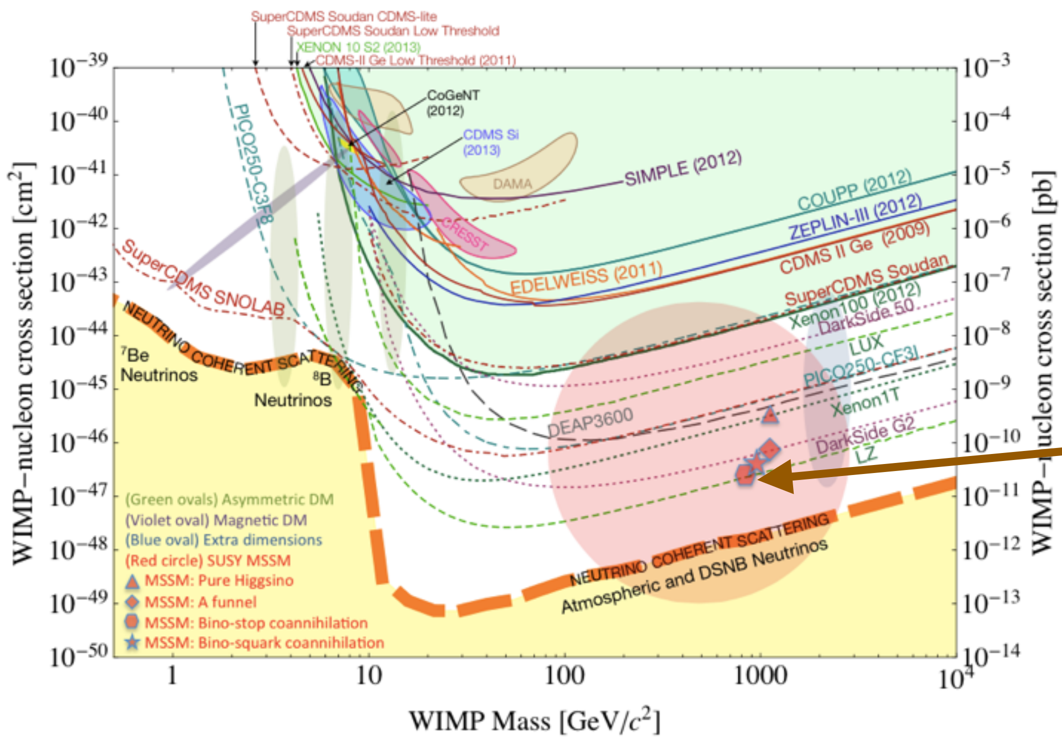


At colliders

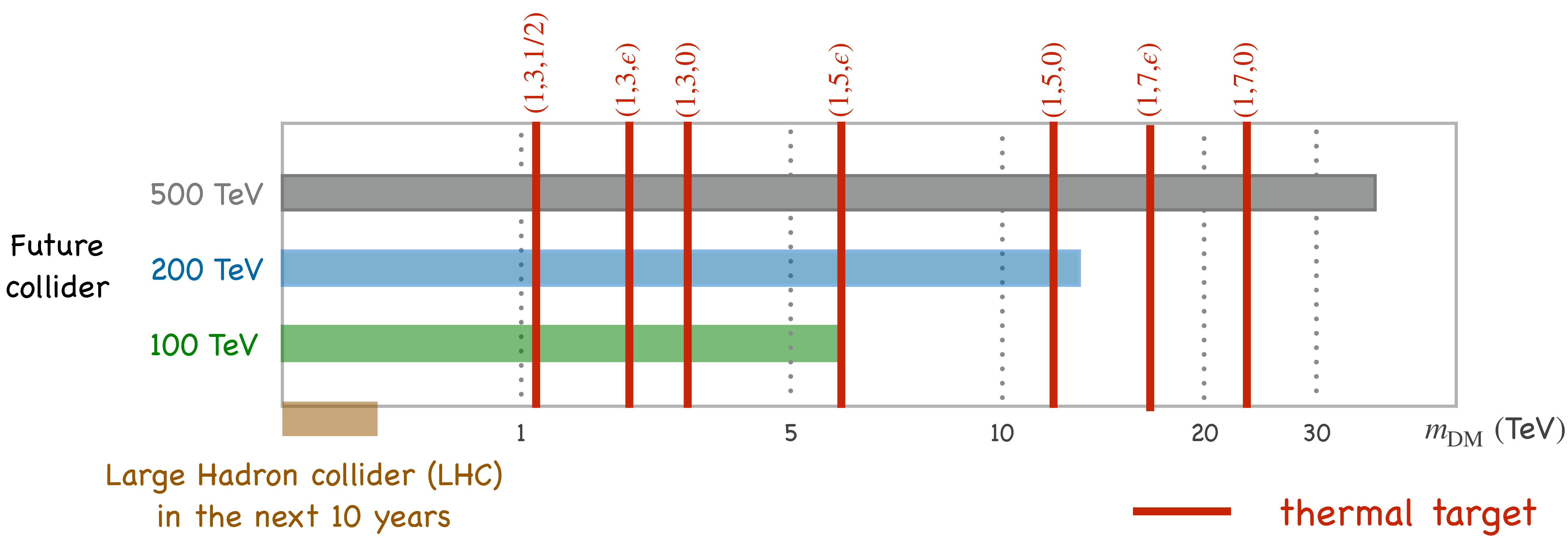




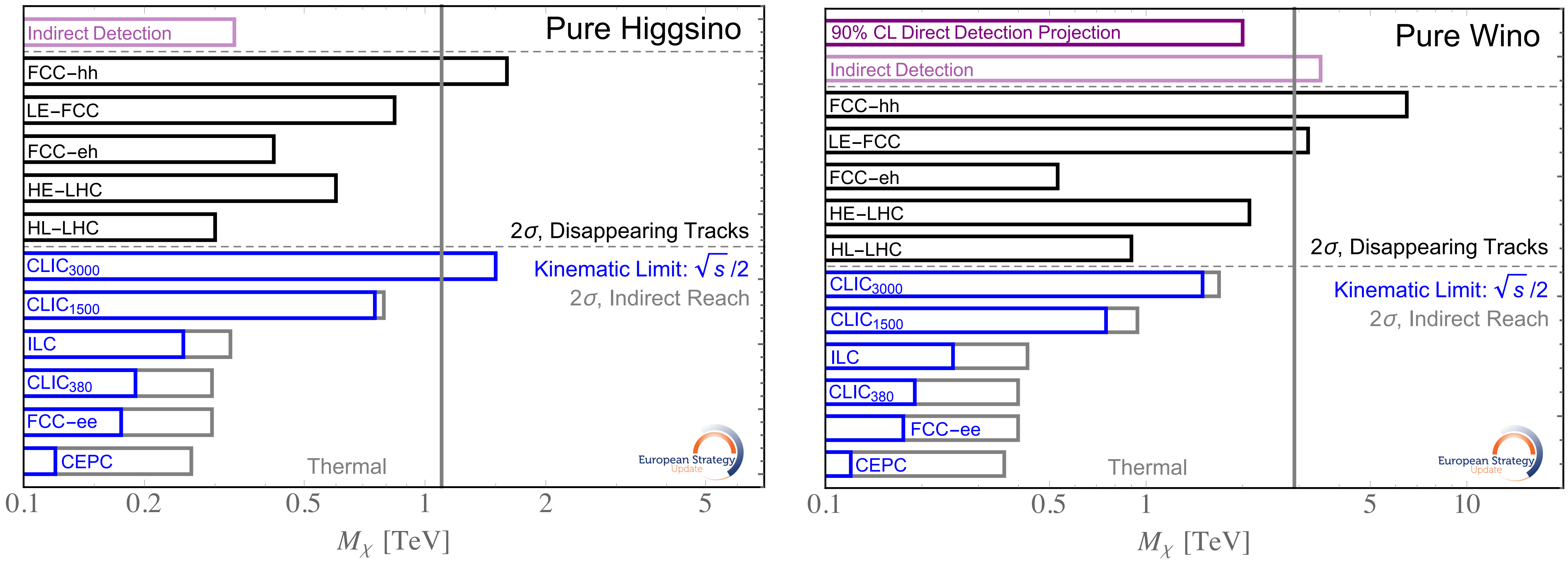
Still a lot to be done



Simplest case yet to be probed



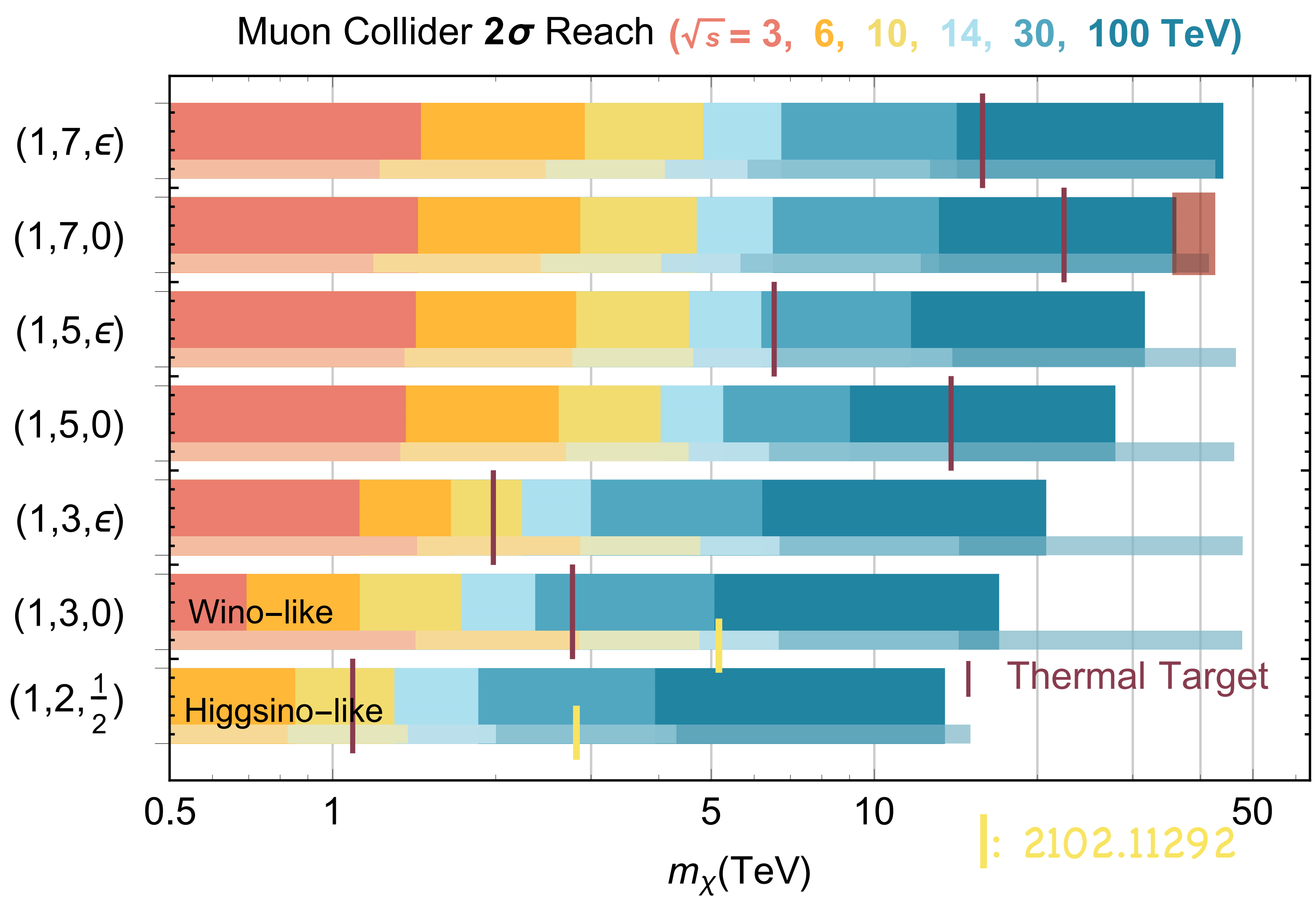
Search at future colliders



100 TeV pp collider is needed
to cover the EW doublet (Higgsino) and triplet (wino) DM.

Not enough to cover the higher dim multiplets.

Reach at muon collider

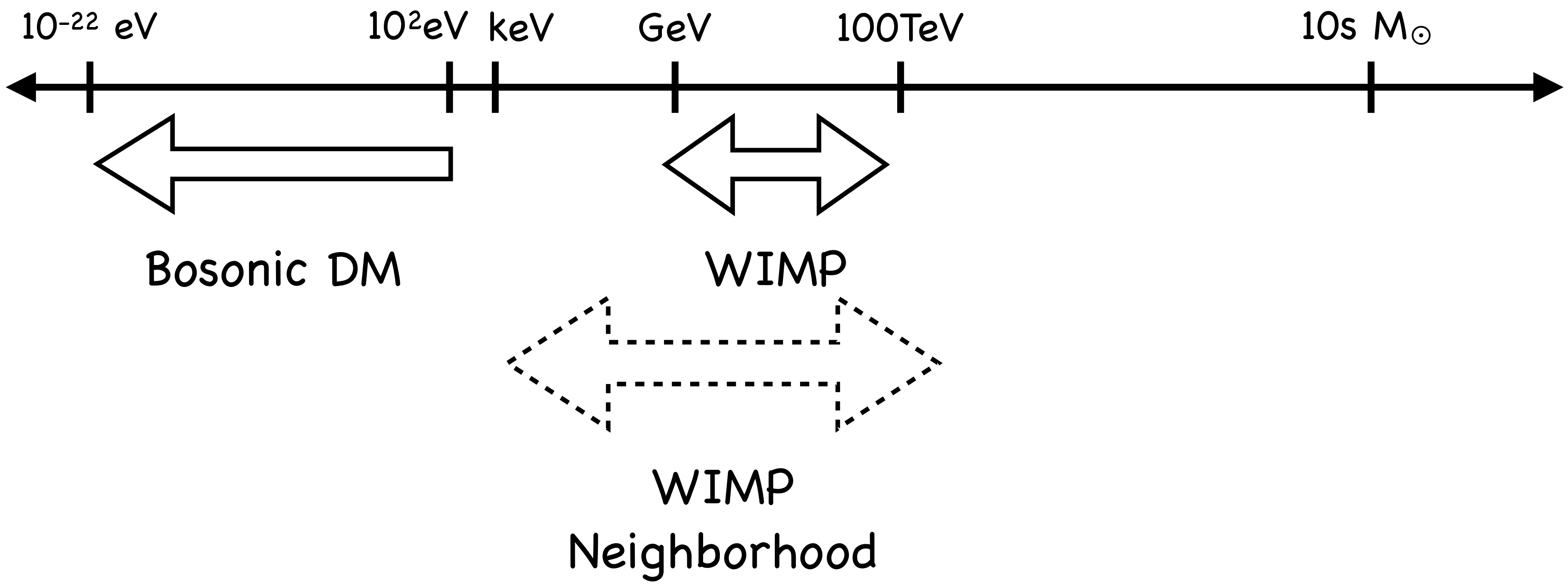


With inclusive signal: E_{CM} ≈ 14 TeV enough to cover n≤3 multiplets.

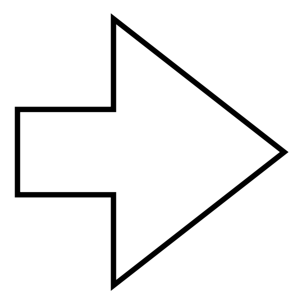
Higher energy needed to cover higher multiplets (almost reaching $m_\chi \approx 1/2\ E_{CM}$).

With disappearing track: potential to reach almost $m_\chi \approx 1/2\ E_{CM}$

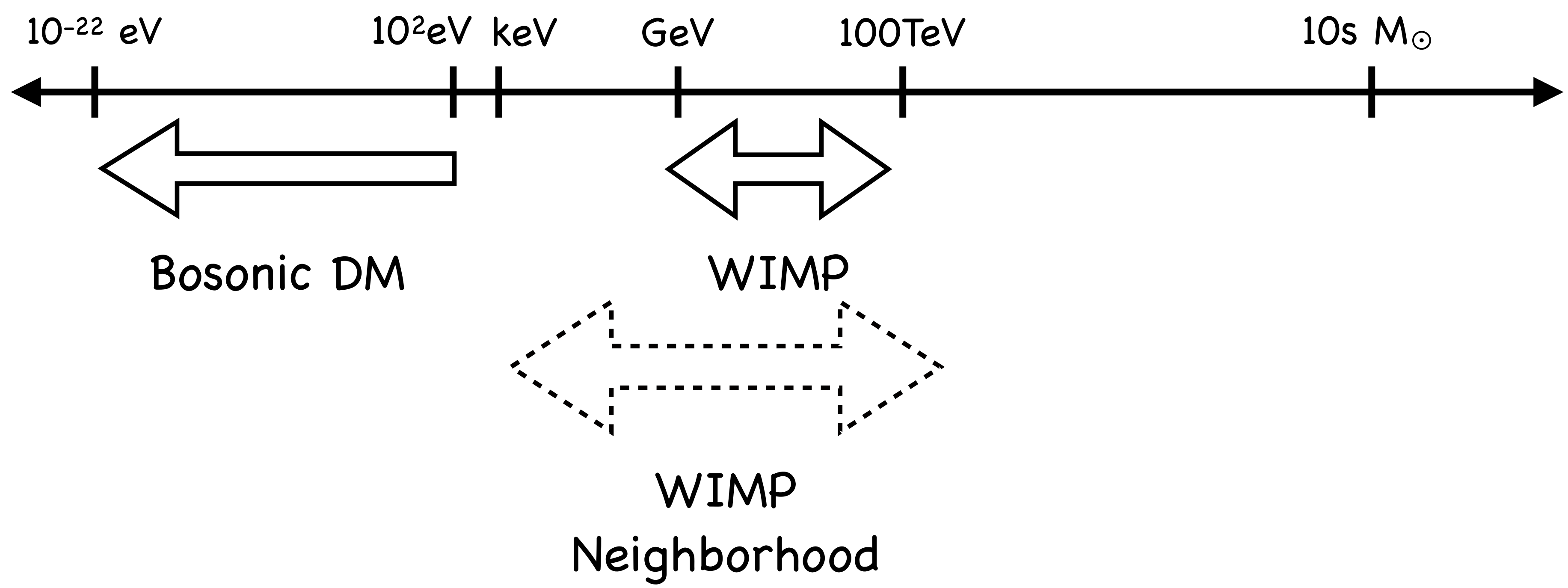
Beyond WIMP



10^{-22} eV
10^2eV
keV
GeV
100TeV
10s $M_\odot$
Bosonic DM
WIMP
WIMP Neighborhood

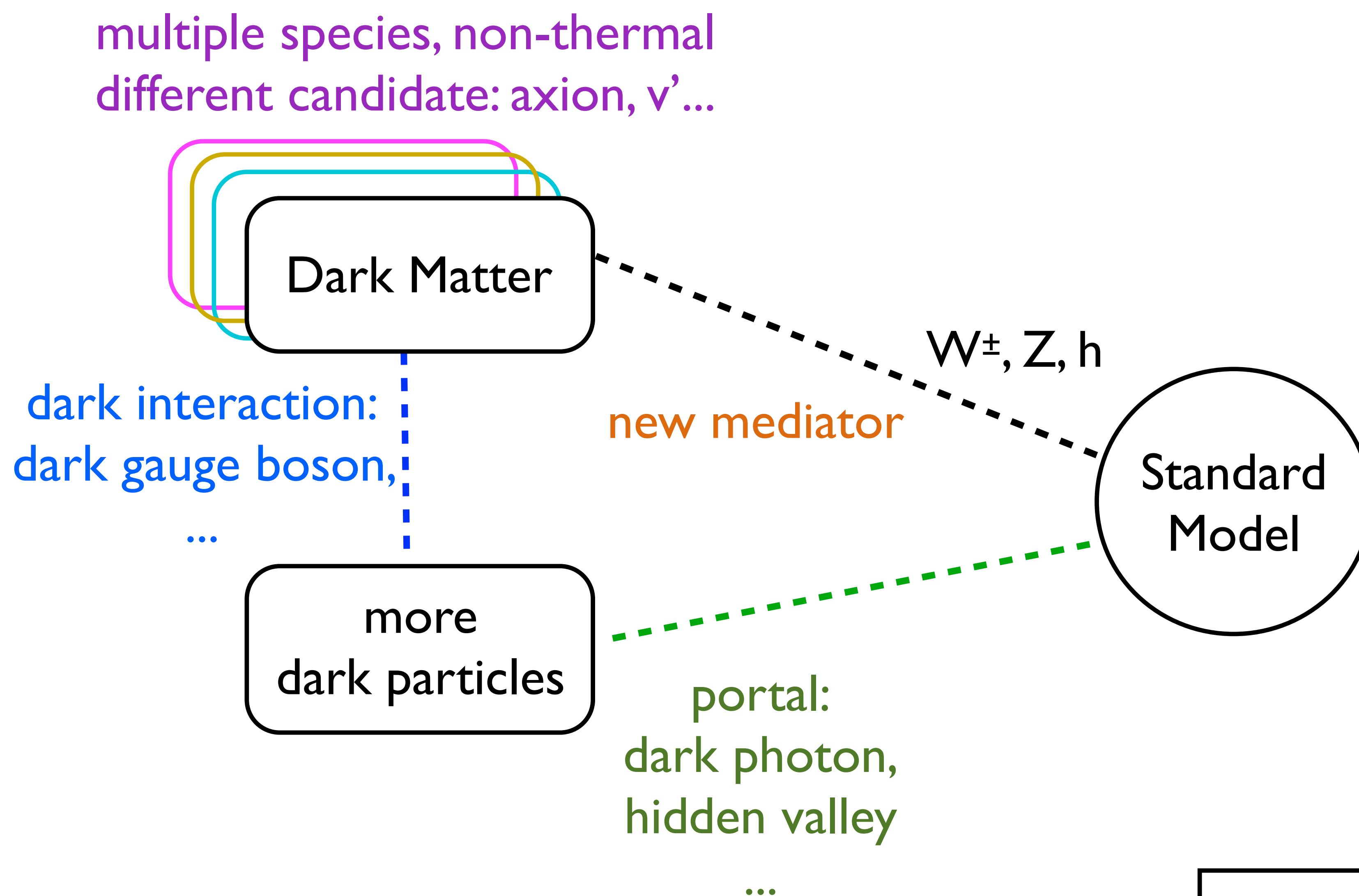


Beyond WIMP


10^{-22} eV
10^{2}eV
keV
GeV
100TeV
10s $M_{\odot}$
Bosonic DM
WIMP
WIMP
Neighborhood


chadwrite.com
chadwrite.com

Beyond the simple WIMPs



Collider searches

Beam dump

Direct detection

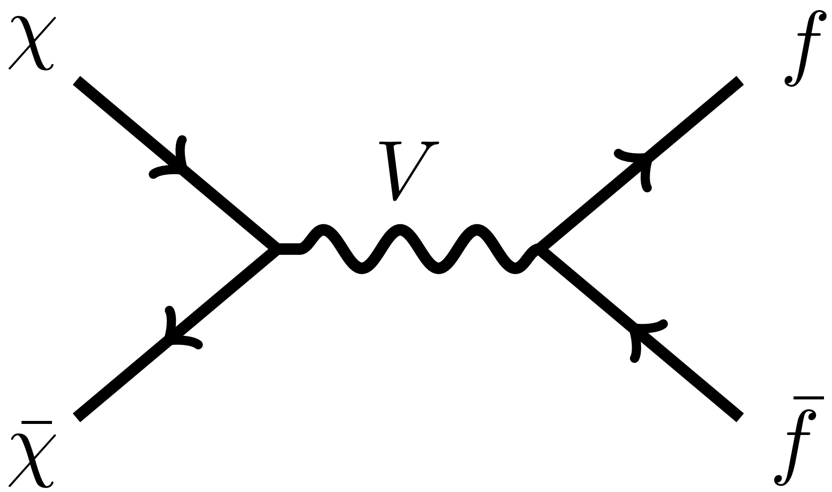
Indirect detection

Connection with early universe

. . . .

New signals. DM may not be the first dark sector discovery.

Extend the WIMP story to lower masses



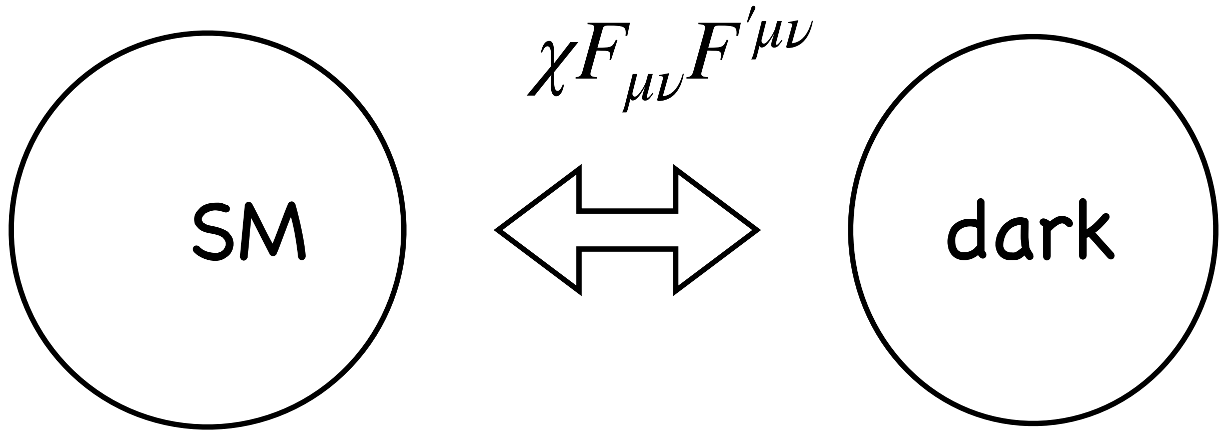
$$m_{\rm DM} < m_V, \quad \sigma v \sim \frac{g^4}{4\pi}\frac{m_{\rm DM}^2}{m_{\rm V}^4}$$

g: coupling

$m_V \approx 10^2$ GeV $\Rightarrow$ m_{DM} > GeV

We should consider lighter mediators, m_V < GeV

Dark photon



photon: $|\gamma\rangle$ photon': $|\gamma'\rangle$

dark photon: a quantum superposition of γ and γ'

$$|\gamma_{\rm dark}\rangle = |\gamma'\rangle + \chi|\gamma\rangle$$

Mediates an interaction with strength $\propto \chi$

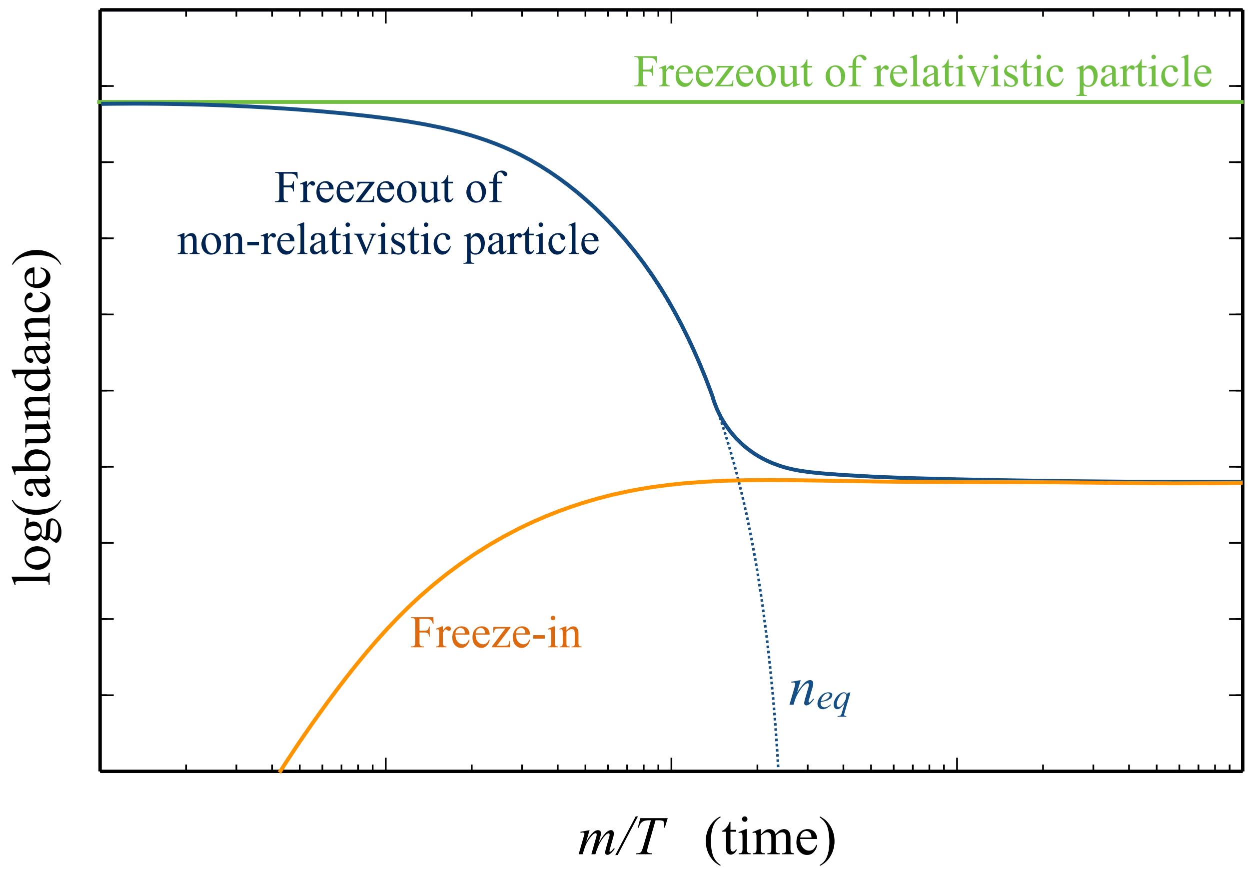
Roles of dark photon

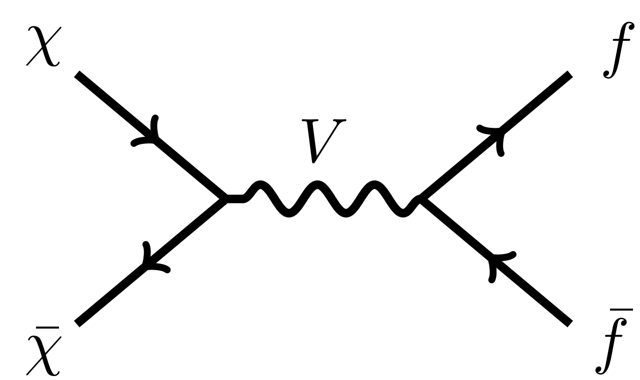
As mediator for thermal freeze out. (Discussed earlier)

Roles of dark photon

As mediator for thermal freeze out. (Discussed earlier)

Freeze-in



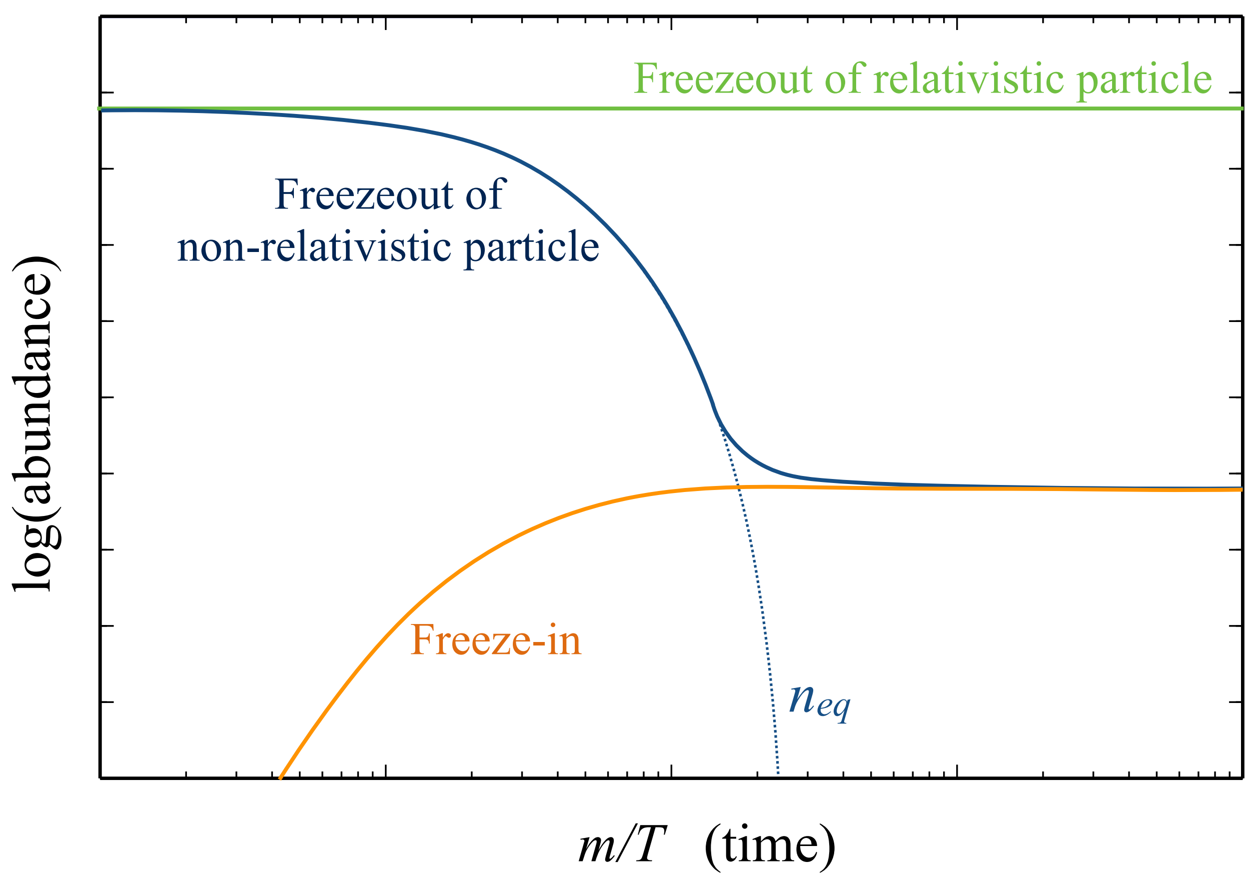


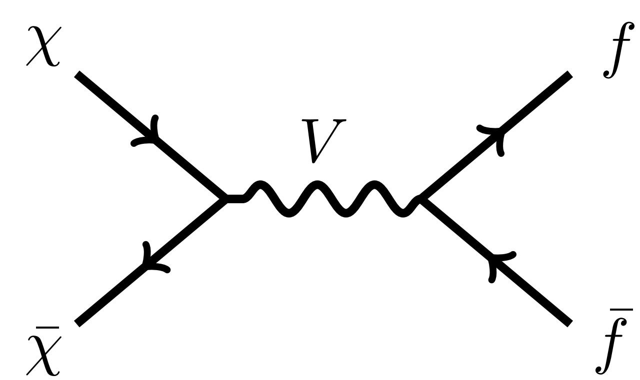
Weak coupling, dark matter not in thermal eq.
It approaches the correct relic abundance.

Roles of dark photon

As mediator for thermal freeze out. (Discussed earlier)

Freeze-in





Weak coupling, dark matter not in thermal eq.
It approaches the correct relic abundance.

Examples

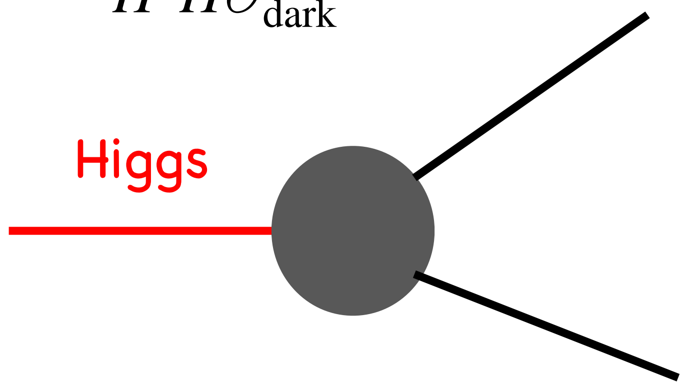
Thermal freeze out: $m_{\rm DM} = 10$ MeV, $m_V = 30$ MeV, $\chi \simeq 10^{-4}$

Freeze in: $m_{\rm DM} = 1$ MeV, $m_V = 10^{-12}$ eV, $\chi \simeq 10^{-6}$

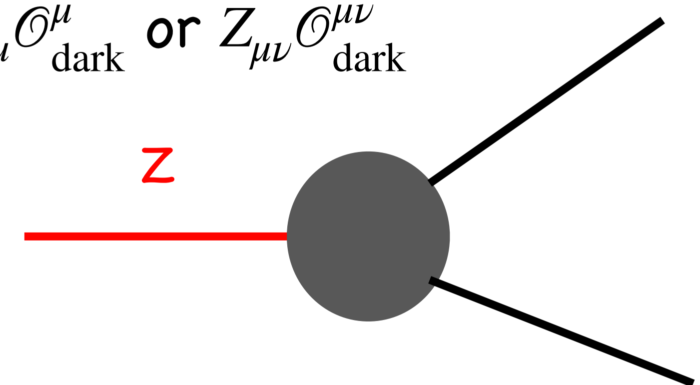
Windows into dark sector: portals

- Any known (SM) particle can in principle have small couplings to dark matter/dark sector.

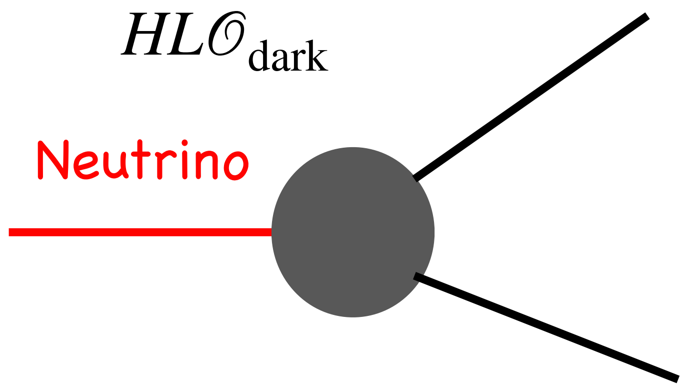
$H^\dagger H \mathcal{O}_{\rm dark}$



$Z_\mu \mathcal{O}^\mu_{\rm dark}$ or $Z_{\mu\nu} \mathcal{O}^{\mu\nu}_{\rm dark}$



$HL\mathcal{O}_{\rm dark}$

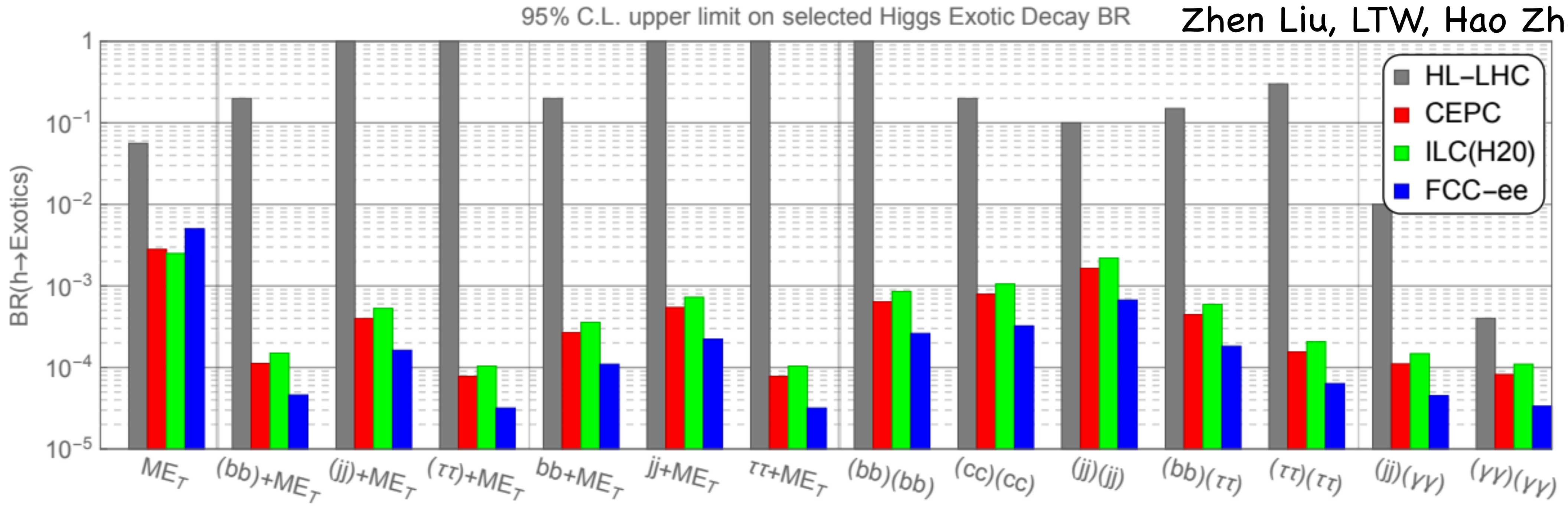


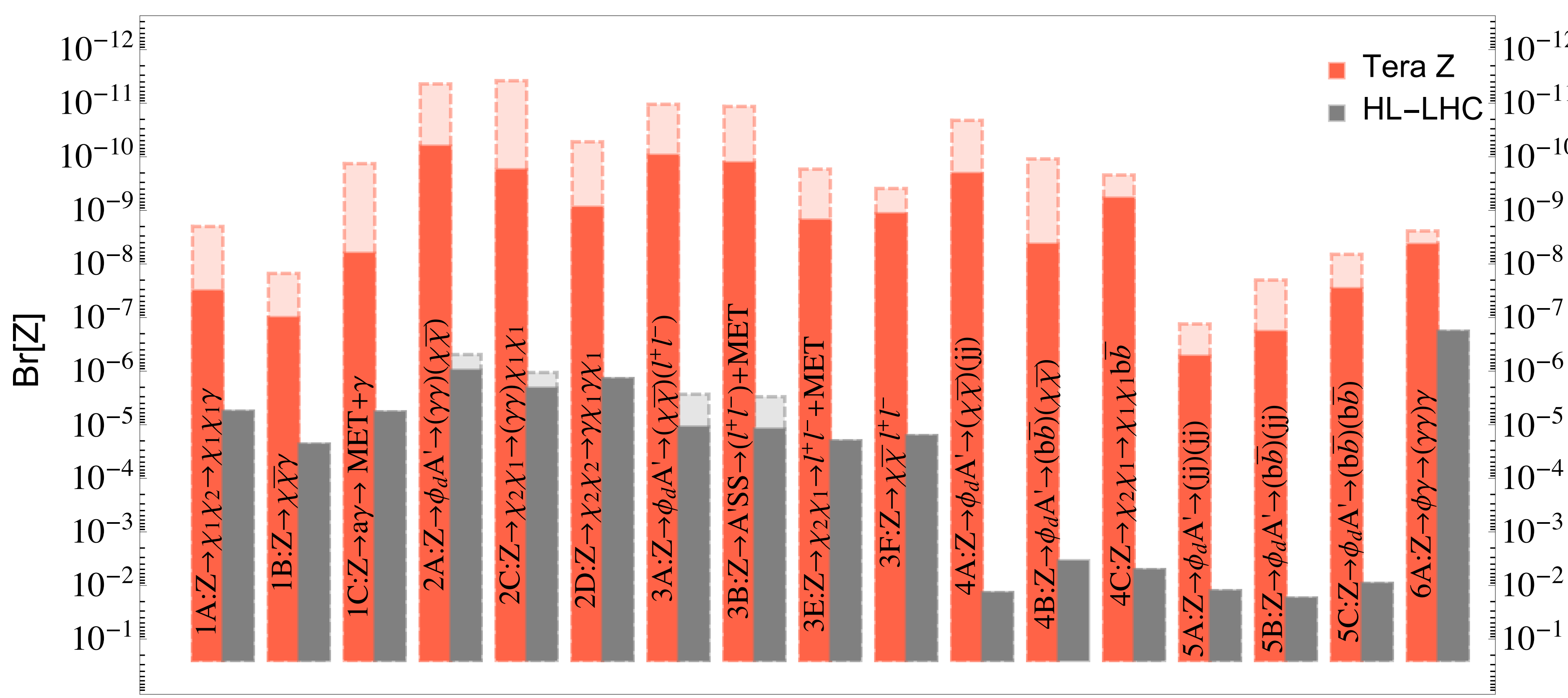
$\mathcal{O}_{\rm dark} =$ SM singlet

Higgs/Z factories, such as CEPC
Neutrino facilities, fixed target experiments…

Higgs/Z factories.

Zhen Liu, LTW, Hao Zhang





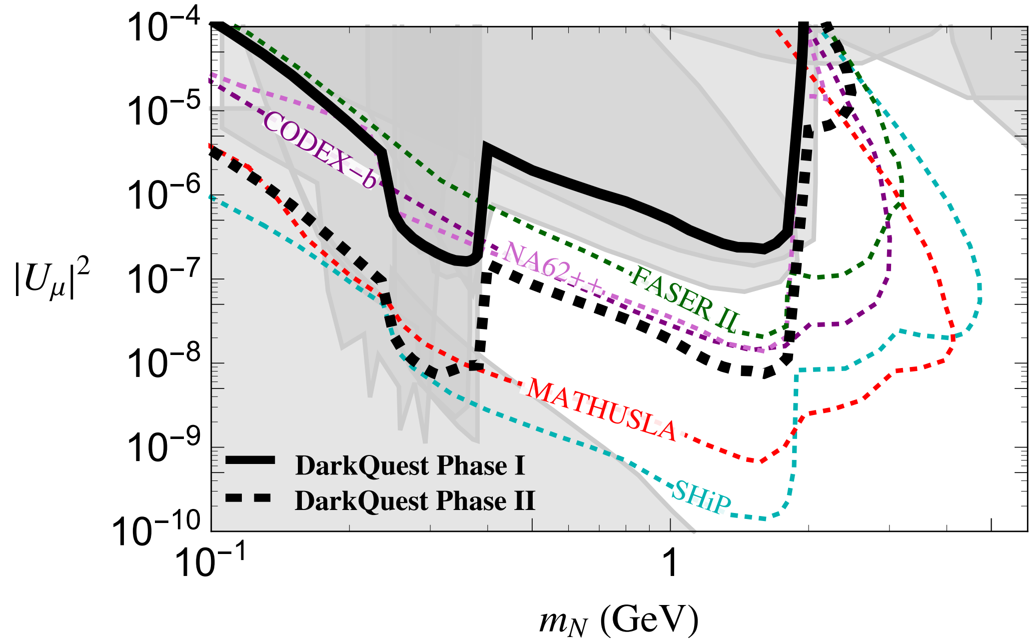
Jia Liu, Xiaoping Wang, Wei Xue, LTW

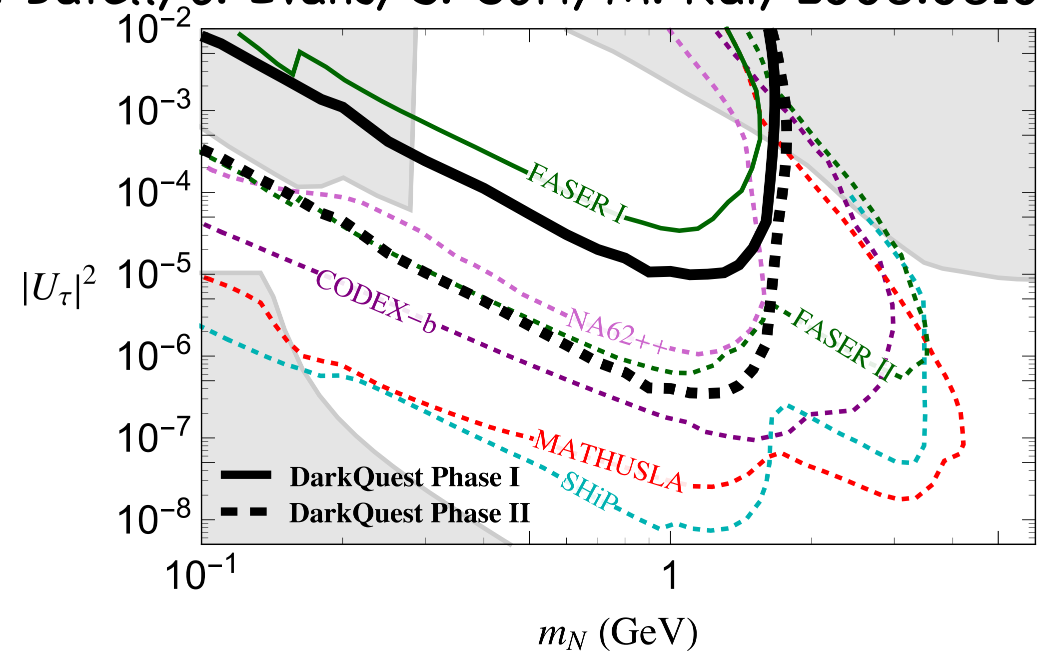
Neutrino portal example: HNL

Heavy Neutral Lepton (HNL):

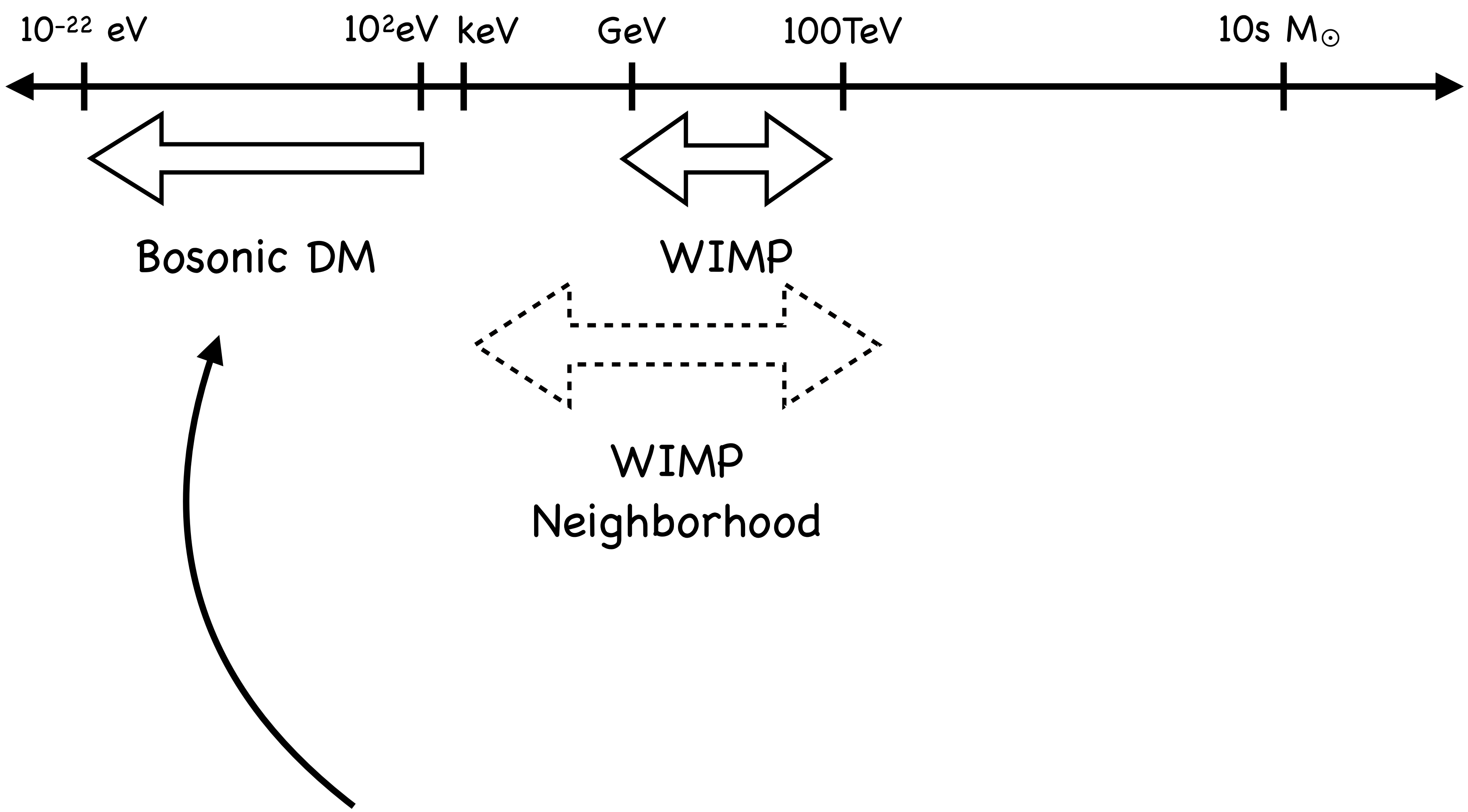
$$\mathscr{L} \supset (HL)N + h.c.$$

B. Batell, J. Evans, S. Gori, M. Rai, 2008.08108





Theories of dark matter



Axion, dark photon

Not single particle-like.

Dark matter = classical wave

$$n_{\text{occupation}} \simeq \frac{\rho_{\text{DM}}}{M_{\text{DM}}} \times \lambda_{dB}^3 = 10^{94} \left(\frac{10^{-22}\text{eV}}{M_{\text{DM}}} \right)^4$$

$$\lambda_{dB} \sim \text{kpc} \left(\frac{10^{-22}\text{eV}}{M_{\text{DM}}} \right) \qquad \rho_{\text{DM}} \simeq 0.4\ \text{GeV/cm}^3$$

- Huge occupation number within a de Broglie wavelength.
  - Collective motion $\rightarrow$classical waves, not a single particle
  - similar to sound, waves on the ocean, traveling on a string...

Classical field in expanding universe

$$\ddot{\phi} + 3H\dot{\phi} + V'(\phi) = 0$$

Expansion of universe
"Viscosity"

Hubble: $H \equiv \frac{\dot{a}}{a}$

$$V'(\phi) = m^2\phi + \cdots$$

Mass + interactions

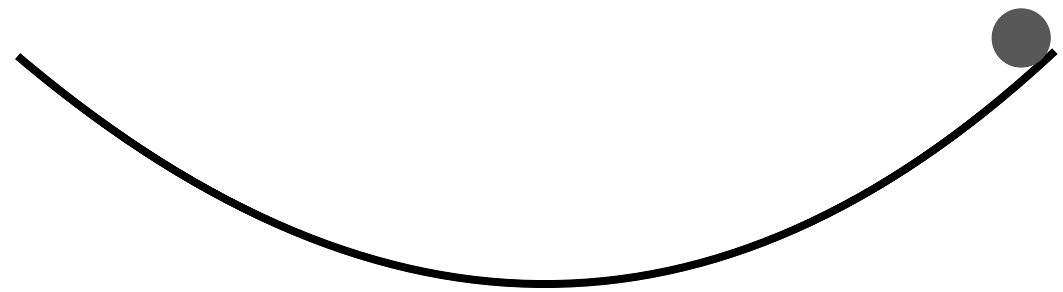
Classical field in two limits

$H > m_\phi \quad \ddot{\phi} + 3H\dot{\phi} + \cancel{V'}(\phi) = 0$

Hubble expansion more important

$H < m_\phi \quad \ddot{\phi} + 3H\dot{\phi} + \cancel{V'}(\phi) = 0$

mass more important



$$\phi(t) \propto \frac{1}{a^{3/2}(t)} \sin(m_\phi t + \phi_0)$$

Classical field in two limits

$H > m_\phi \quad \ddot{\phi} + 3H\dot{\phi} + \cancel{V'}(\phi) = 0$

Hubble expansion more important

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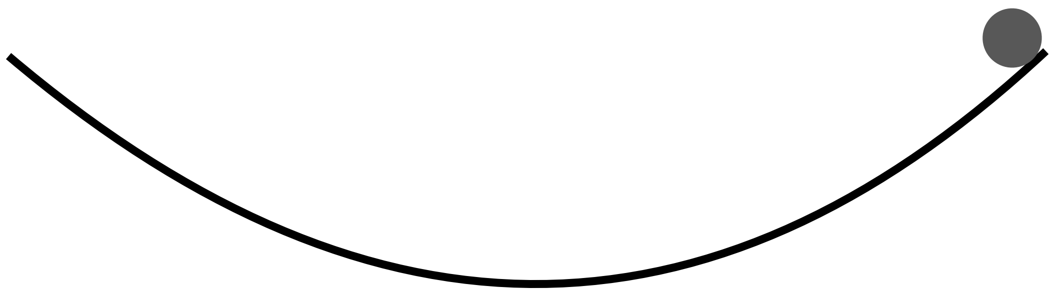
mass more important



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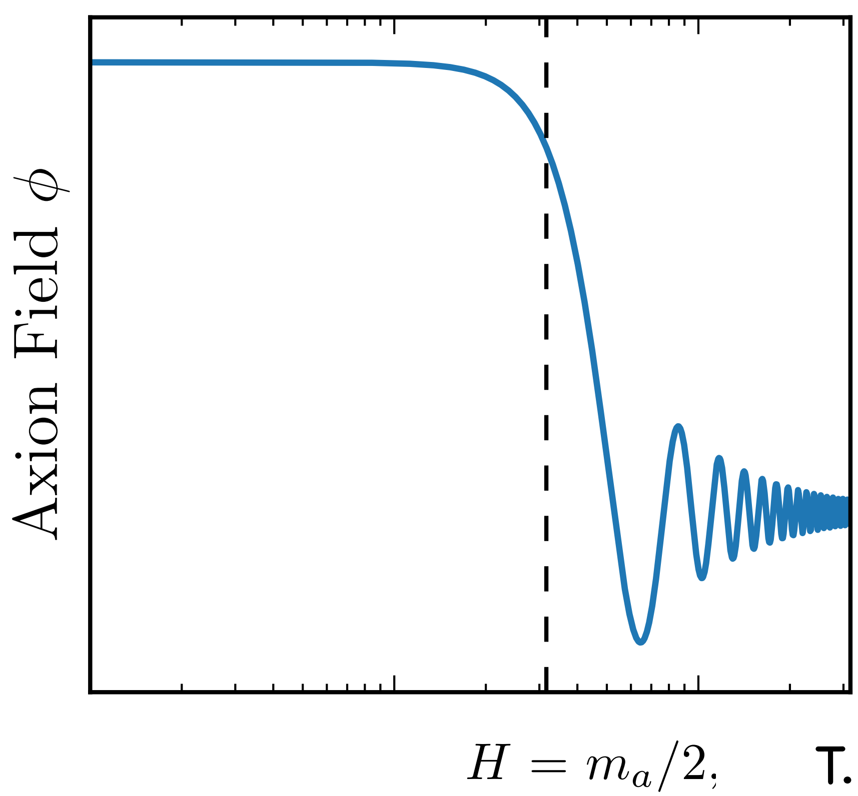


T. Lin 2018 TASI lecture

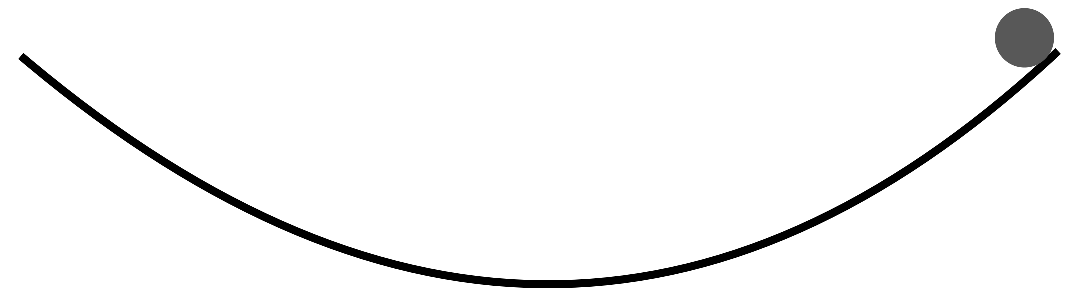
Classical field in two limits

$H > m_\phi \quad \ddot{\phi} + 3H\dot{\phi} + \cancel{V'}(\phi) = 0$

Hubble expansion more important

$H < m_\phi \quad \ddot{\phi} + 3H\dot{\phi} + \cancel{V'}(\phi) = 0$

mass more important



$$\phi(t) \propto \frac{1}{a^{3/2}(t)} \sin(m_\phi t + \phi_0)$$

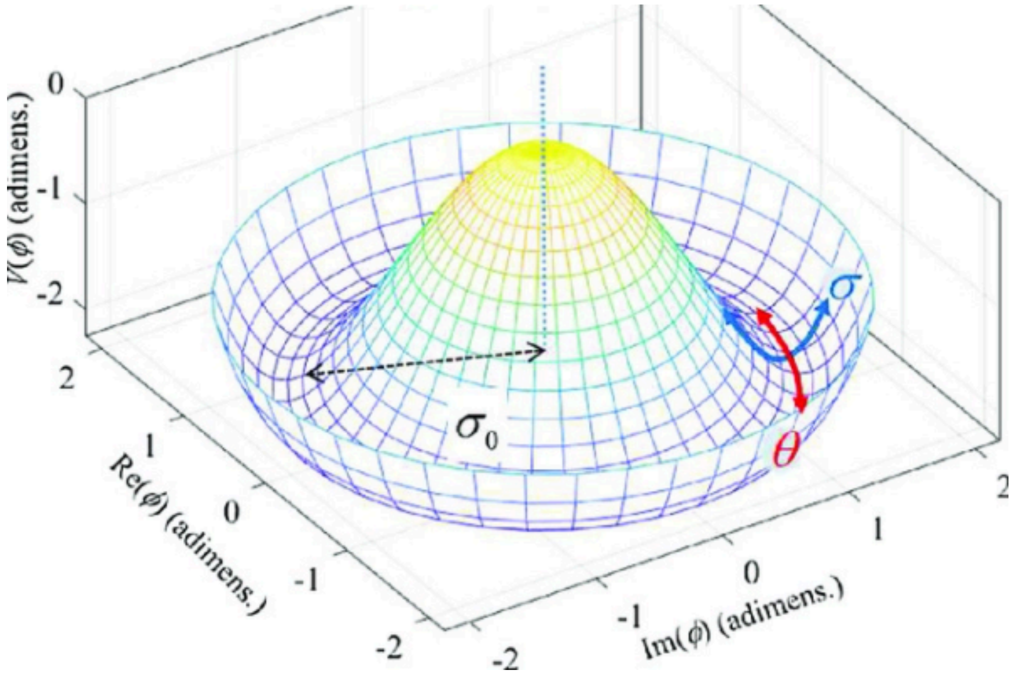
$$\rho \propto m_\phi^2 \phi^2(t) \propto \frac{1}{a^3(t)}$$

On large scales, behave same as particle-like matter.

Will similarly cluster, form structure, etc.

Why is axion light?

Potential of a symmetry breaking



A very common phenomenon:
1) Standard Model electroweak symmetry breaking. Strong interaction.
2) Condensed matter system: phonon, magnets, BCS...

Excitation in θ direction massless. "Goldstone" boson.

Symmetry $\theta \to \theta + c \Rightarrow \theta$ is massless.

Small mass can then be generated by a small coupling.

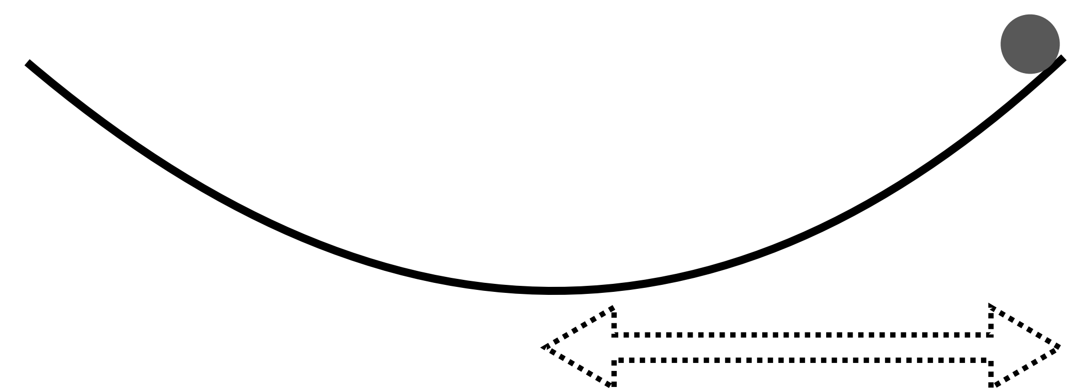
Production: misalignment

$H > m_\phi$

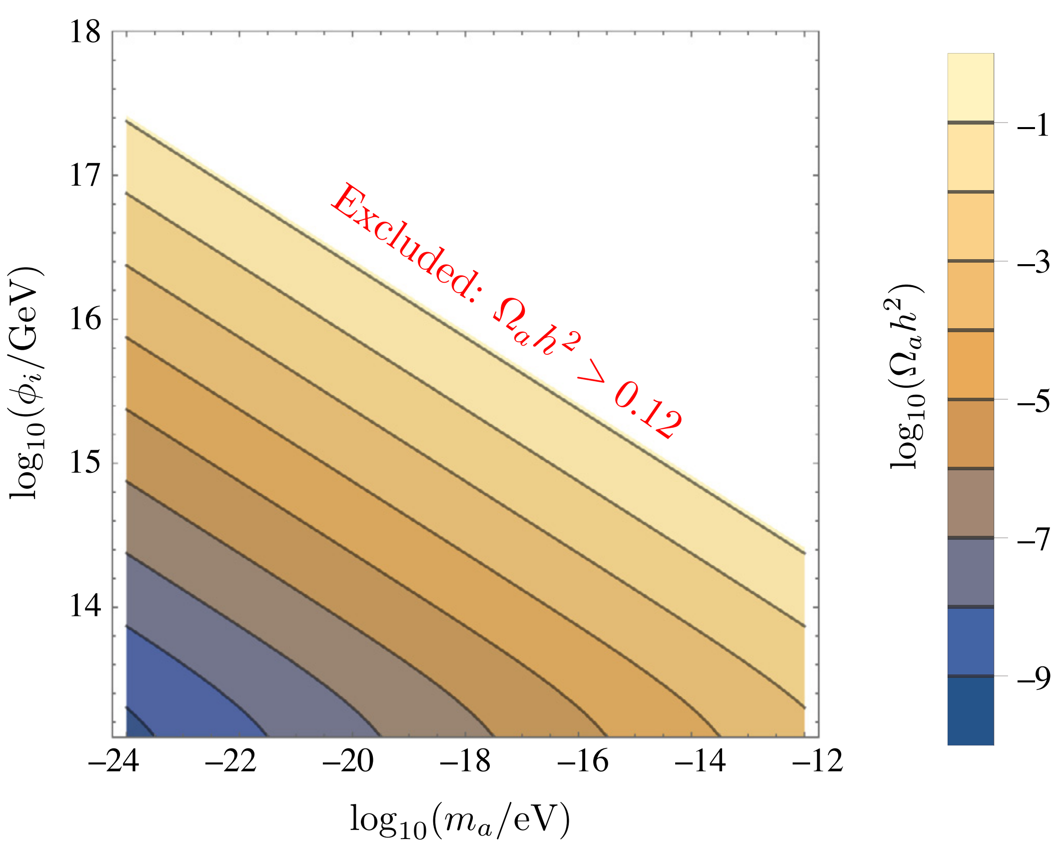
Hubble expansion more important

$H < m_\phi$

mass more important



ϕ_i



$$\phi(t) = \frac{\phi_i}{a^{3/2}(t)} \sin(m_\phi t + \phi_0)$$

$$\rho = m_a^2 \phi^2(t) \propto \phi_i^2$$

Need large initial value
Possible during inflation.

“The axion” and ALP

QCD (strong interaction) axion: the axion

Axion from breaking of a U(1) PQ global symmetry.
Axion mass generated by small non-perturbative effect of strong interaction.

Motivation: QCD strong CP problem.

The neutron electric dipole moment expected from QCD is wrong by at least 9 orders of magnitude.

Axion gives a dynamical solution to this problem.

“The axion” and ALP

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Axion gives a dynamical solution to this problem.

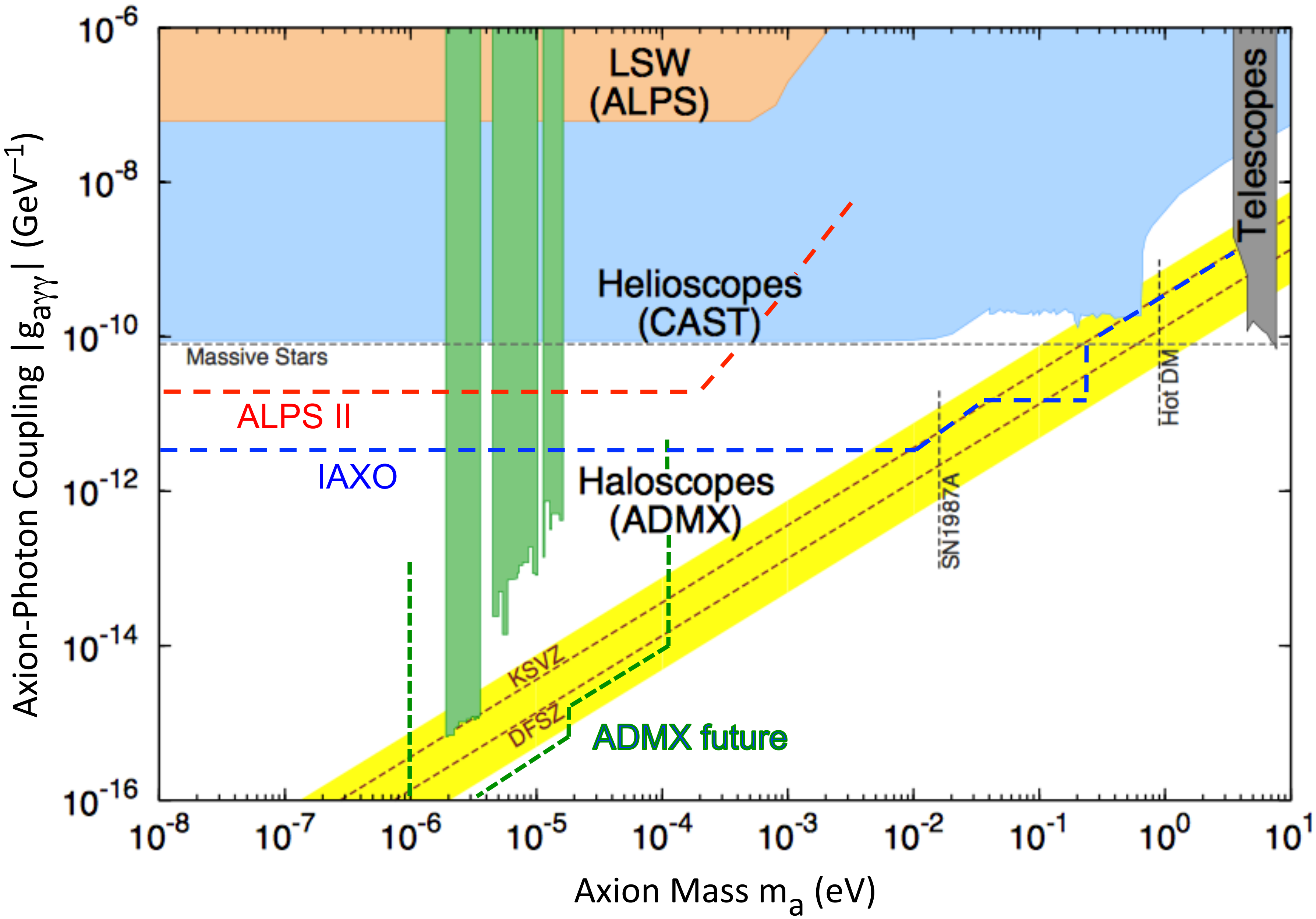
Axion like particles: ALPs

Similar light scalar particles.
The property is not dictated by the strong interaction. More general scenarios than the QCD axion.

Axion coupling to the known particles

Main detection channel relies on axion photon coupling

$$g_{a\gamma\gamma} a F_{\mu\nu} \tilde{F}^{\mu\nu}$$

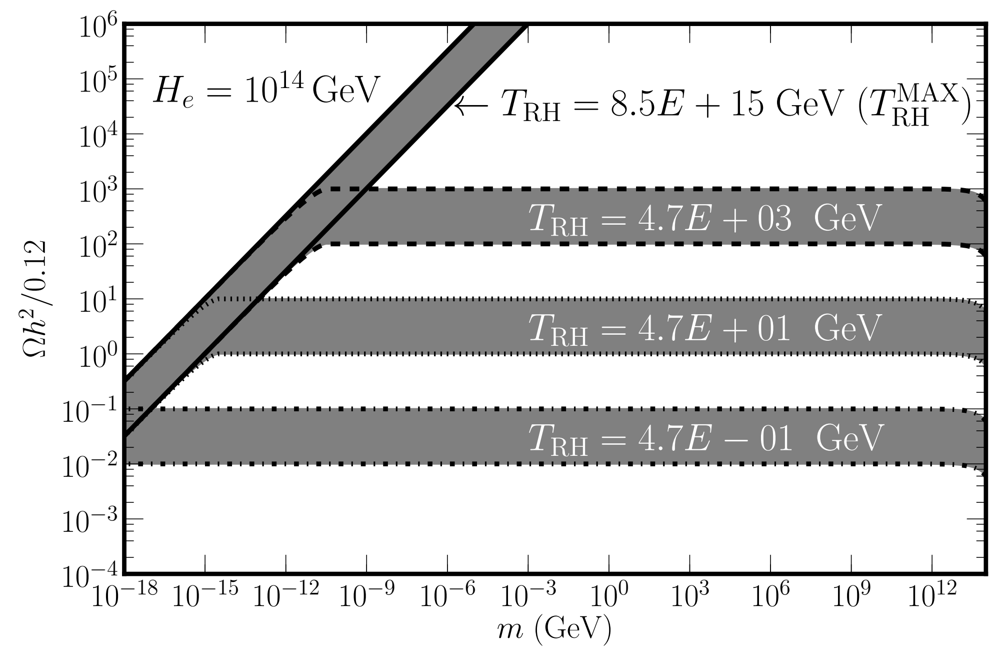


Dark photon dark matter

Multiple production mechanisms:

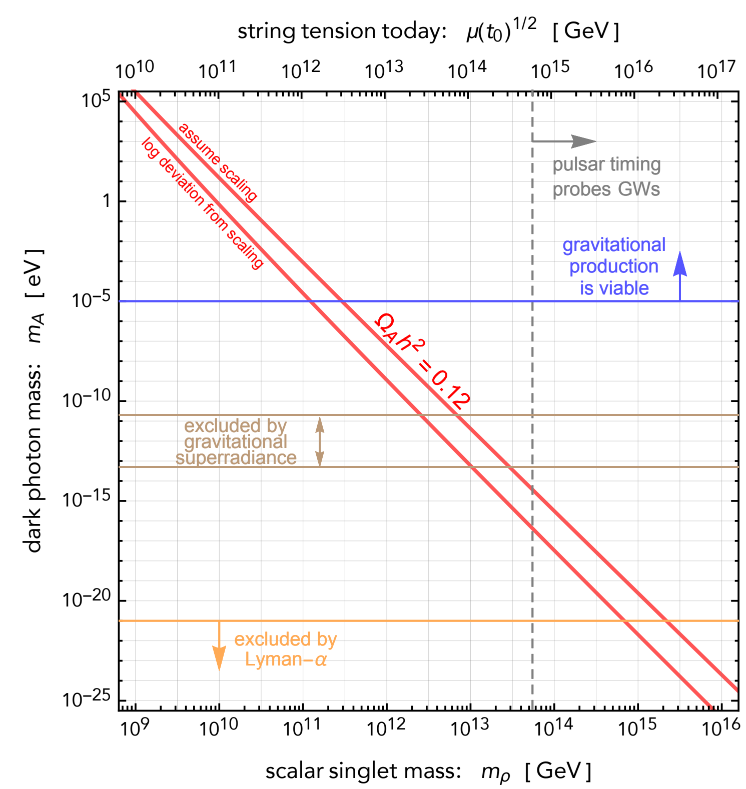
Produced gravitationally during inflation and reheating

Similar to Hawking radiation, but applied to expanding universe.

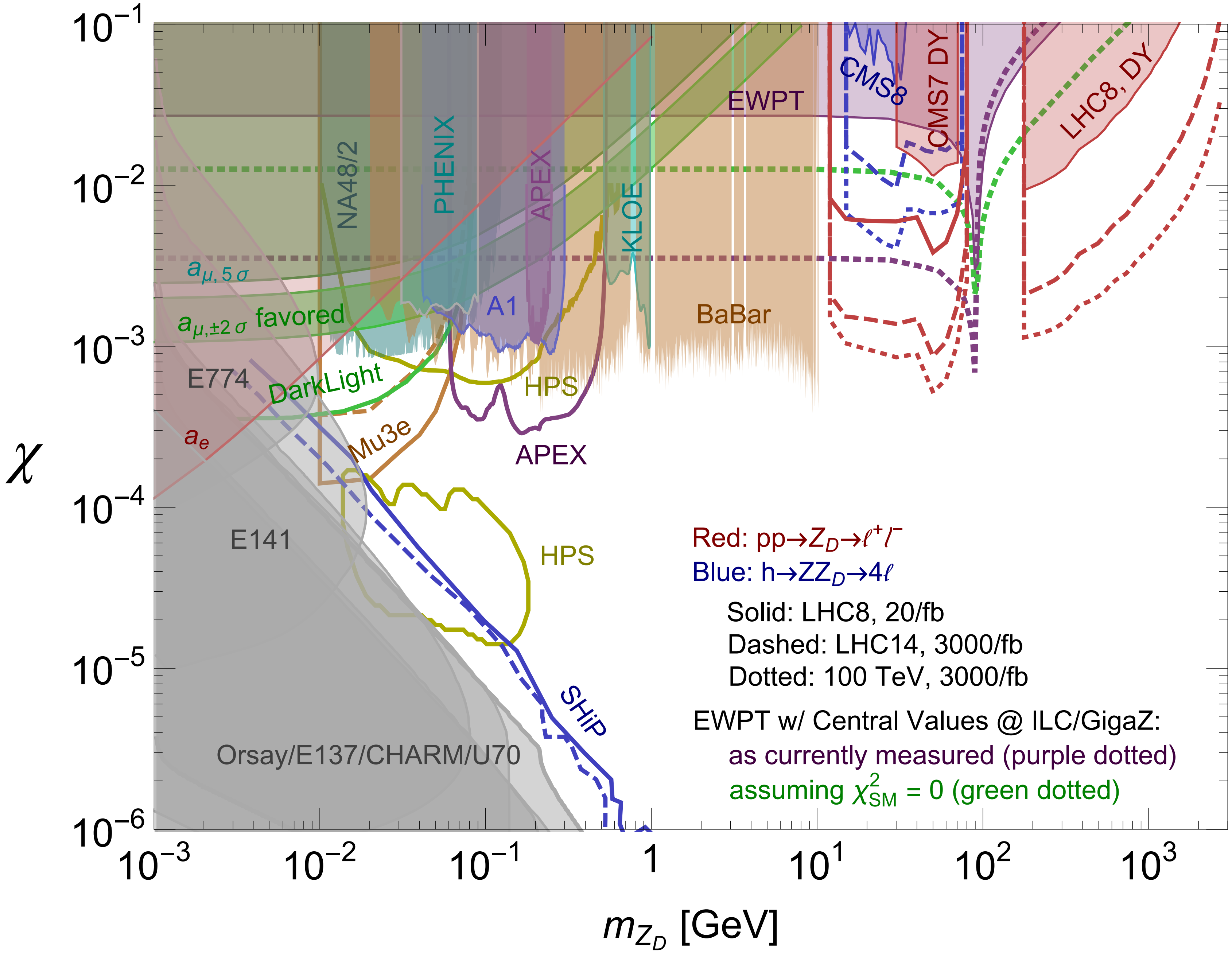


From topological defects (such as cosmic strings) radiation

+ from misalignment, coupling to axions, ...

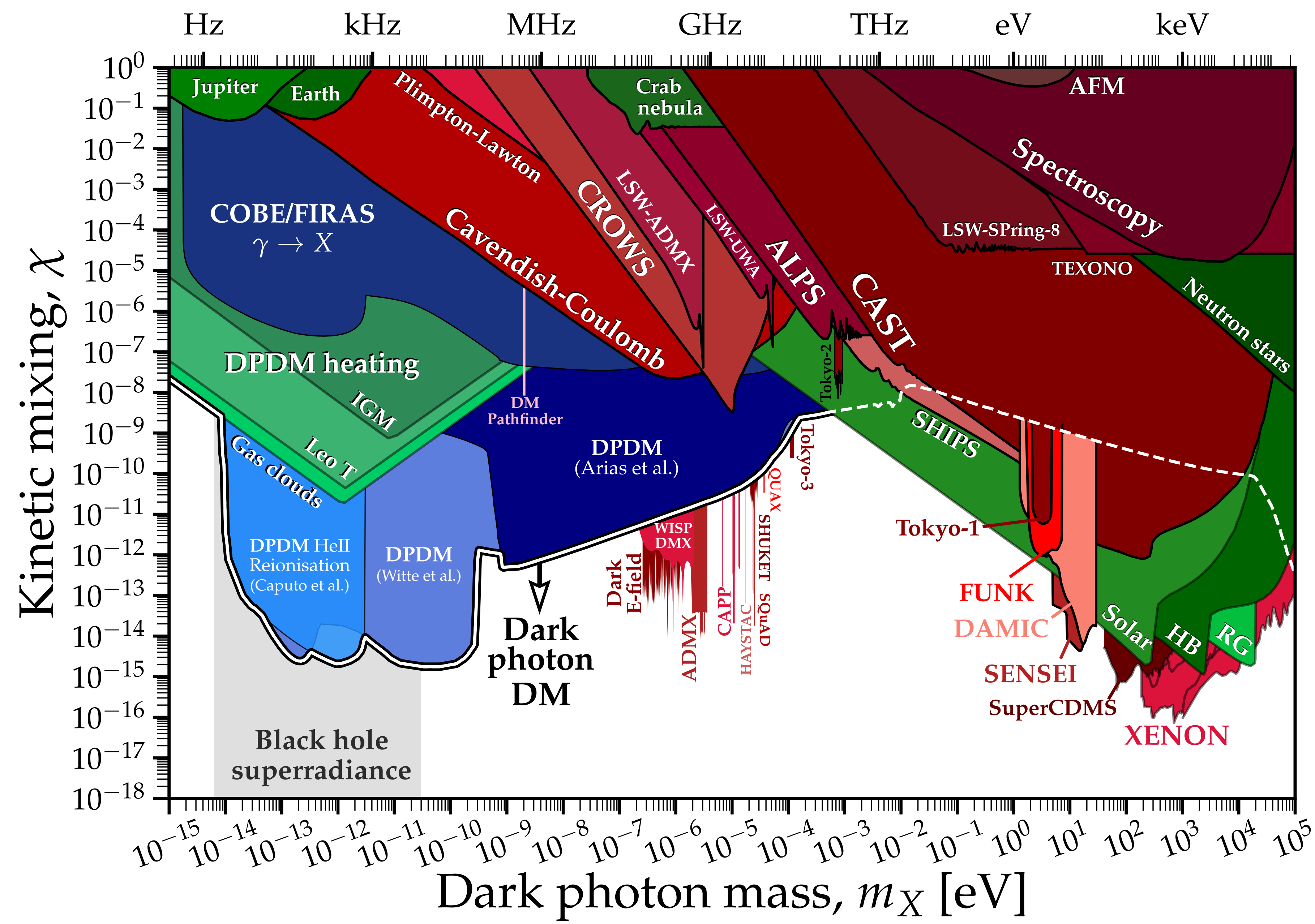


Dark photon searches



Heavier dark photon: Colliders, fixed target experiments

Dark photon searches



Lighter dark photon: Terrestrial/table top detectors, astrophysical, ...

Many new ideas still emerging: e.g. H. An, F. P. Huang, J. Liu, W. Due, 2010.15836

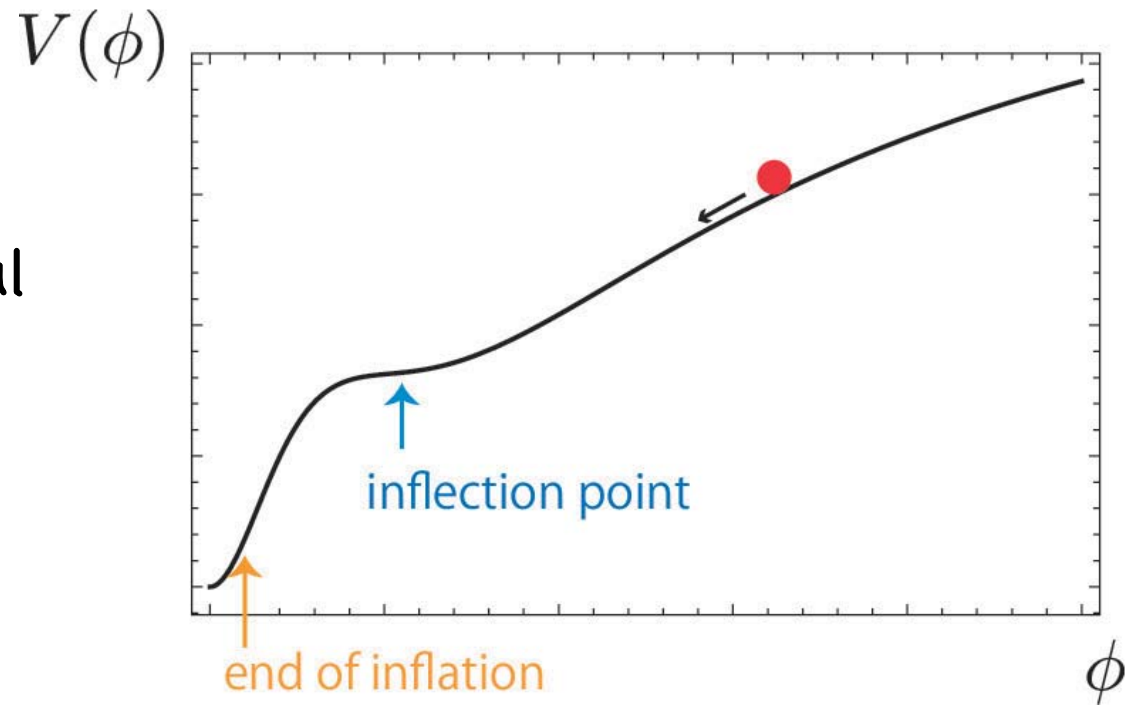
Are axion/dark photon good theories?

- Quite reasonable:
  - Based on well known physics (such as Goldstone boson and symmetry breaking).
  - QCD axion can solve strong CP problem.
- Testable:
  - Simple coupling to the Standard Model. Large possible region of coupling strength.
  - Many new development for new techniques.

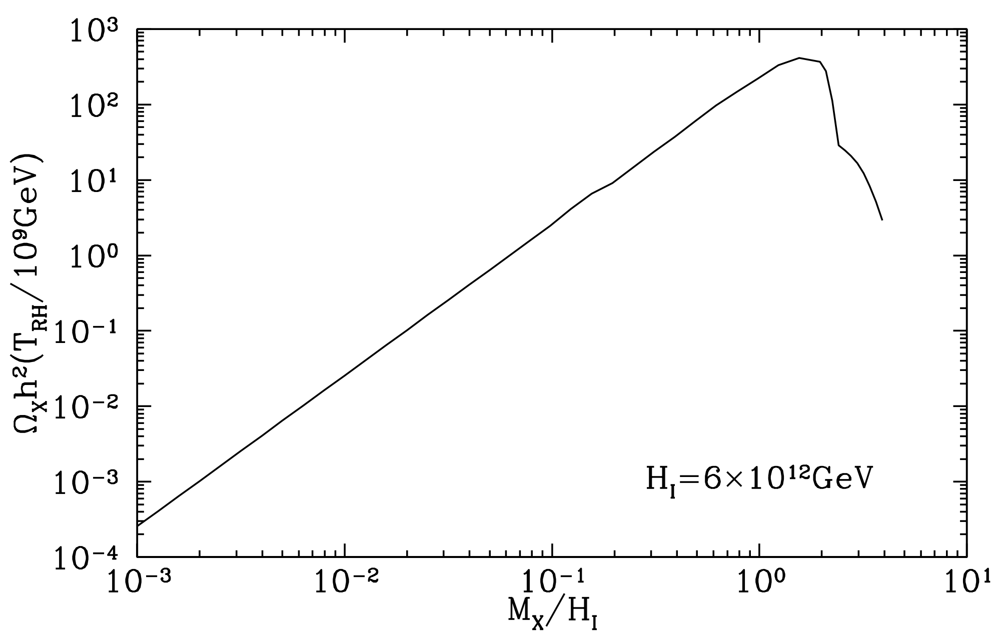
Pretty good theories. Good guide for experiments.

Other stories

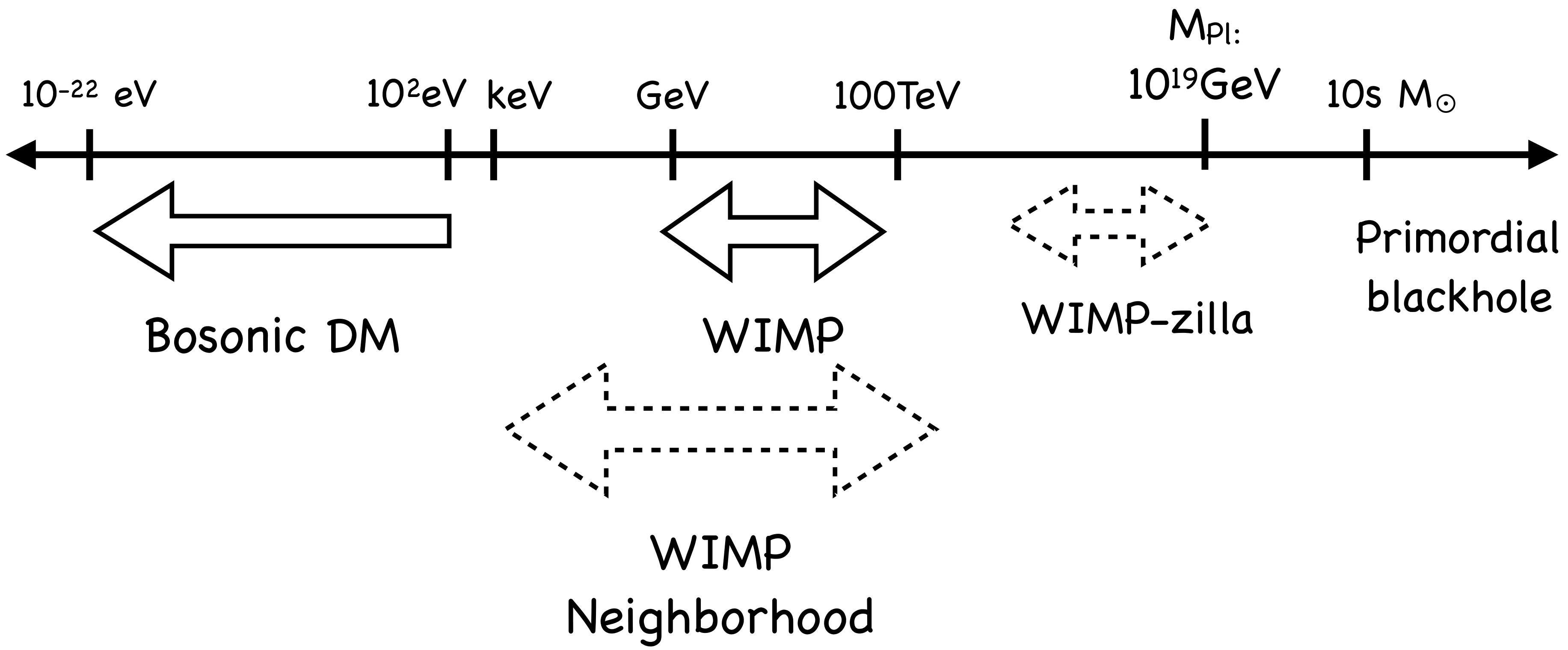
A special feature on the inflaton potential gives large fluctuations
⇒ primordial blackhole production



Gravitational effect during inflation and reheating can produce (very) heavy particles "WIMP-zillars" ($10^{12\text{-}15}$ GeV)



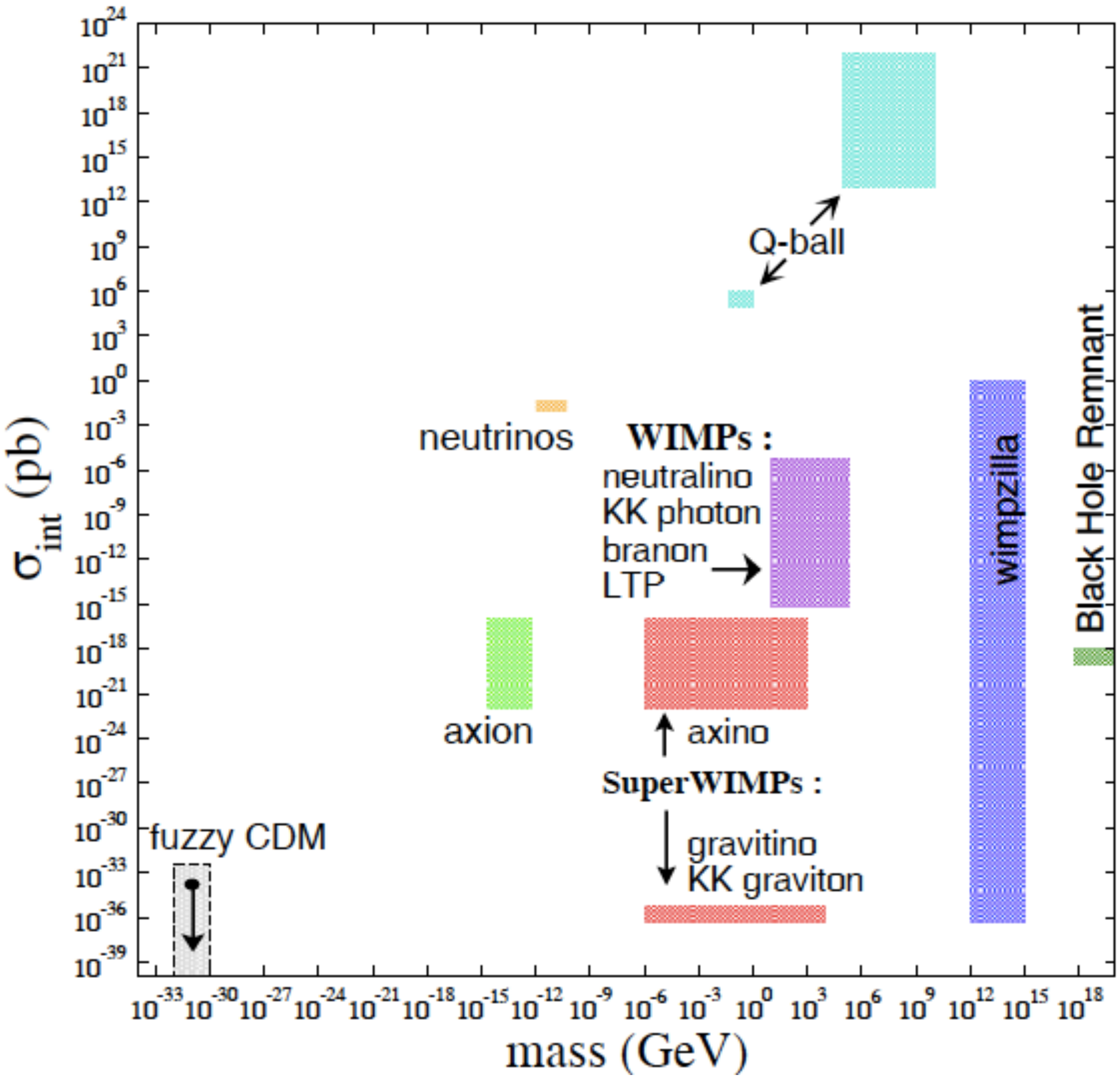
The gaps in our stories



- Still, many orders of magnitudes empty.

The gaps

J. Feng



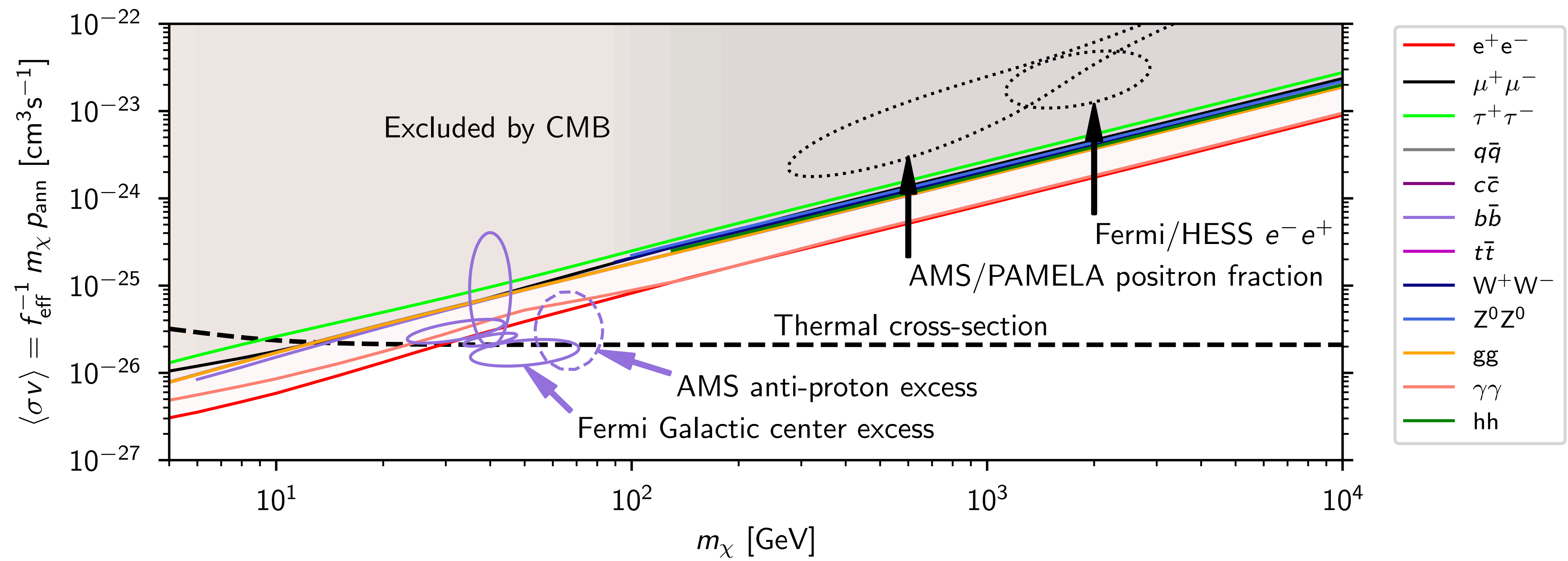
– Even more if we also take into account coupling strength

Conclusion

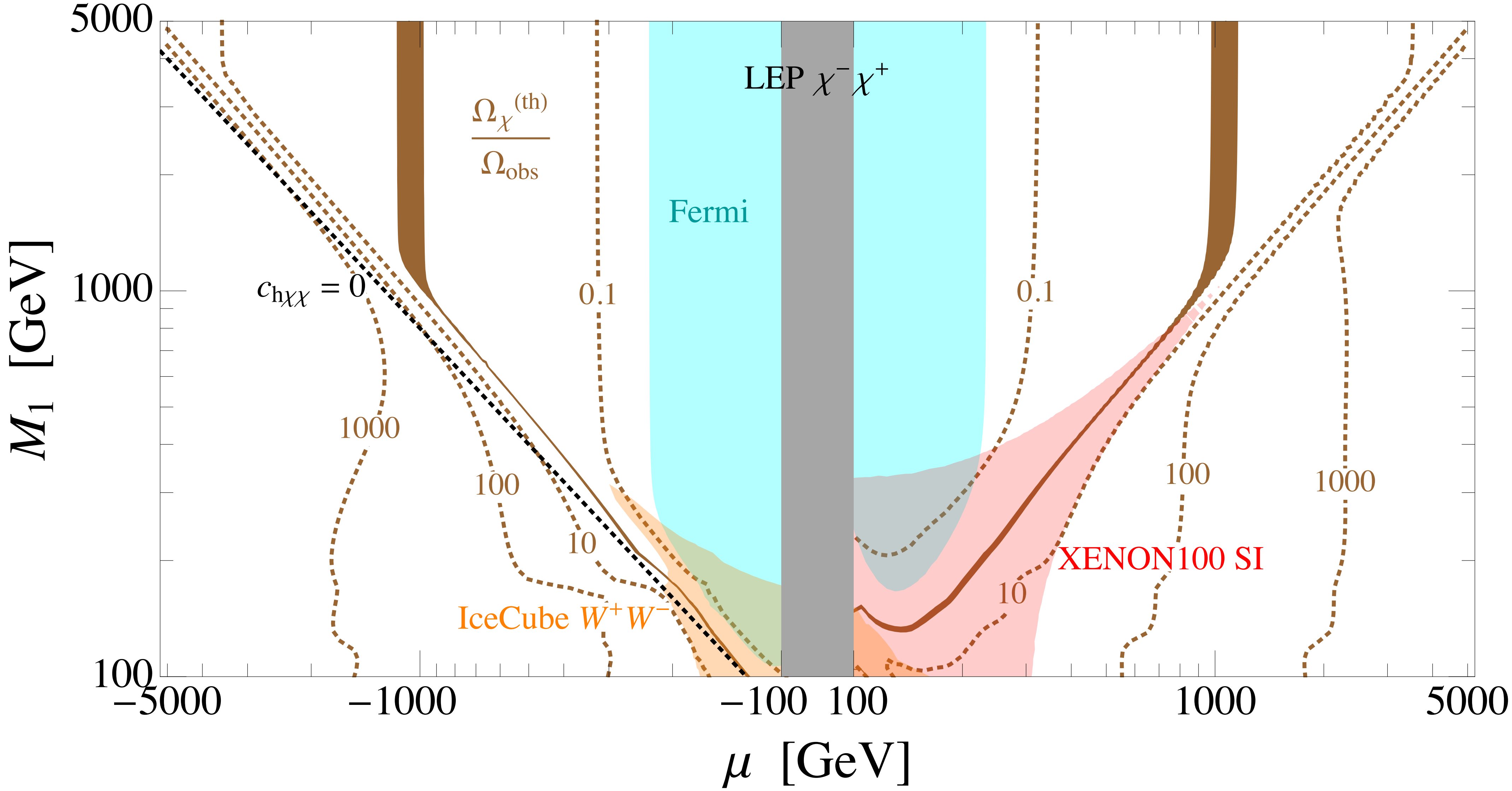
- Understanding dark matter is one of the most pressing questions in physics.
- It is also a very difficult task.
  - Vast range of possibilities, yet we know extremely little.
- After decades of effort, we have several good theories: WIMP + dark sector, axion,...
- Much more is needed to cast a wide net! Huge amount of work left to be done!

extra

Limits from annihilation



Narrowing parameter space.

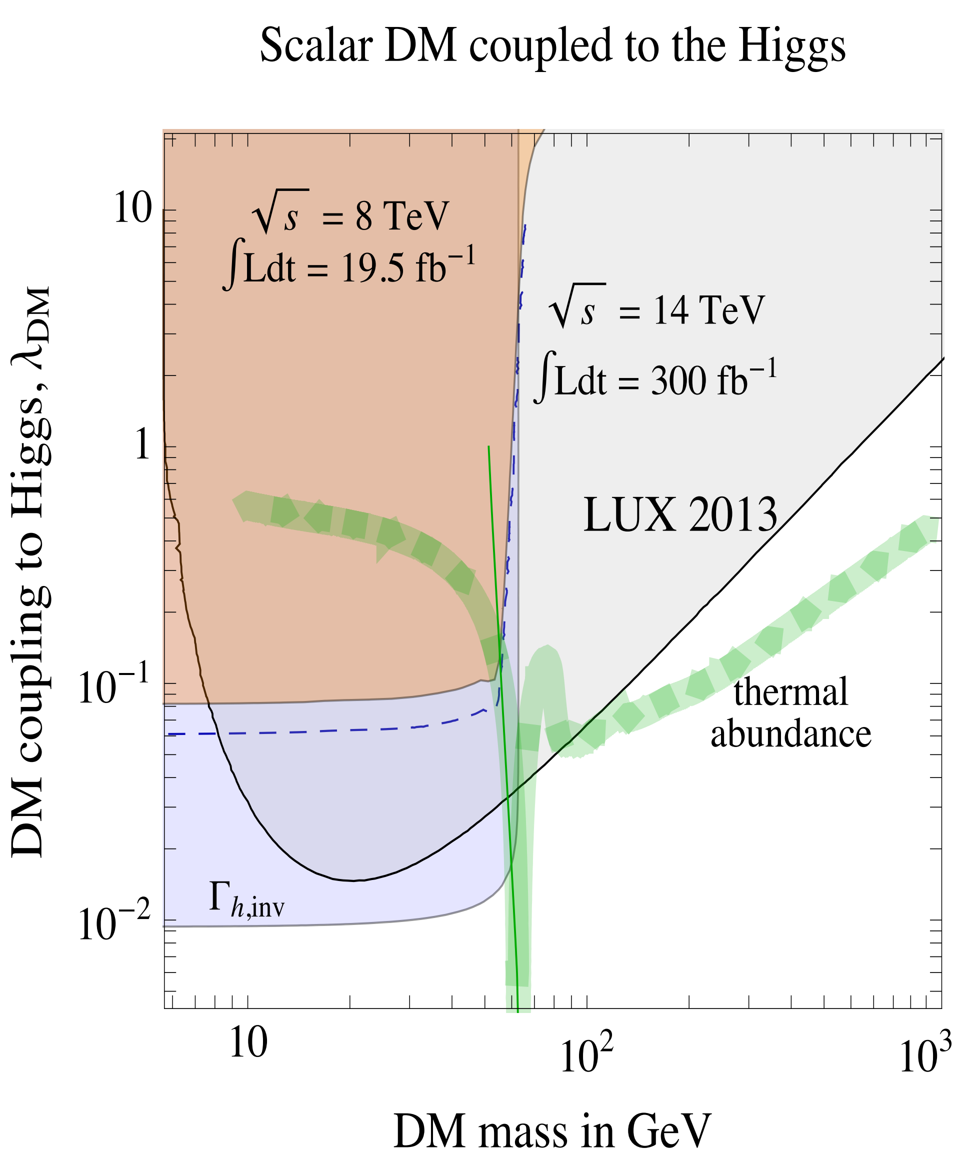


Arkani-Hamed, Delgado, Giudice, 0601041
Cheung, Hall, Pinner, Ruderman, 1211.4873

Example: Higgs portal

$$H^\dagger H \mathcal{O}_{\rm NP}$$

H⁺H lowest dim gauge inv op. Most relevant coupling to SM singlet.



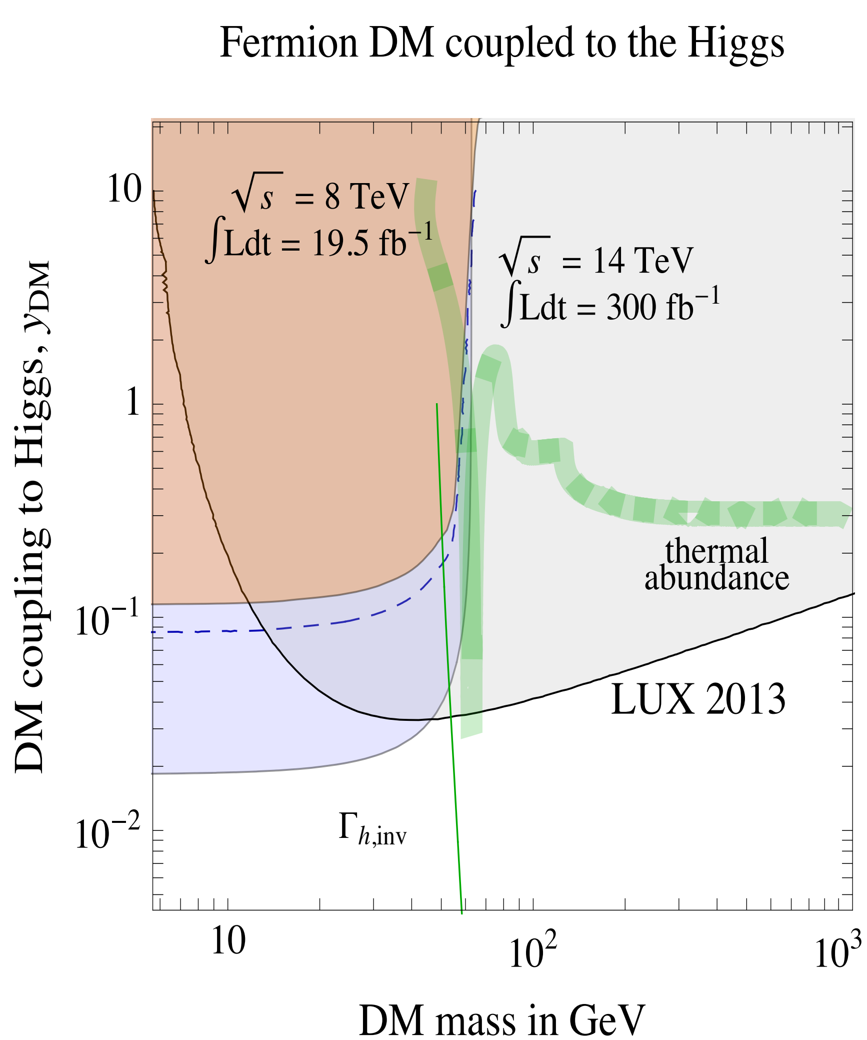
m_{DM} < 1/2 m_h Higgs Γ_{inv}

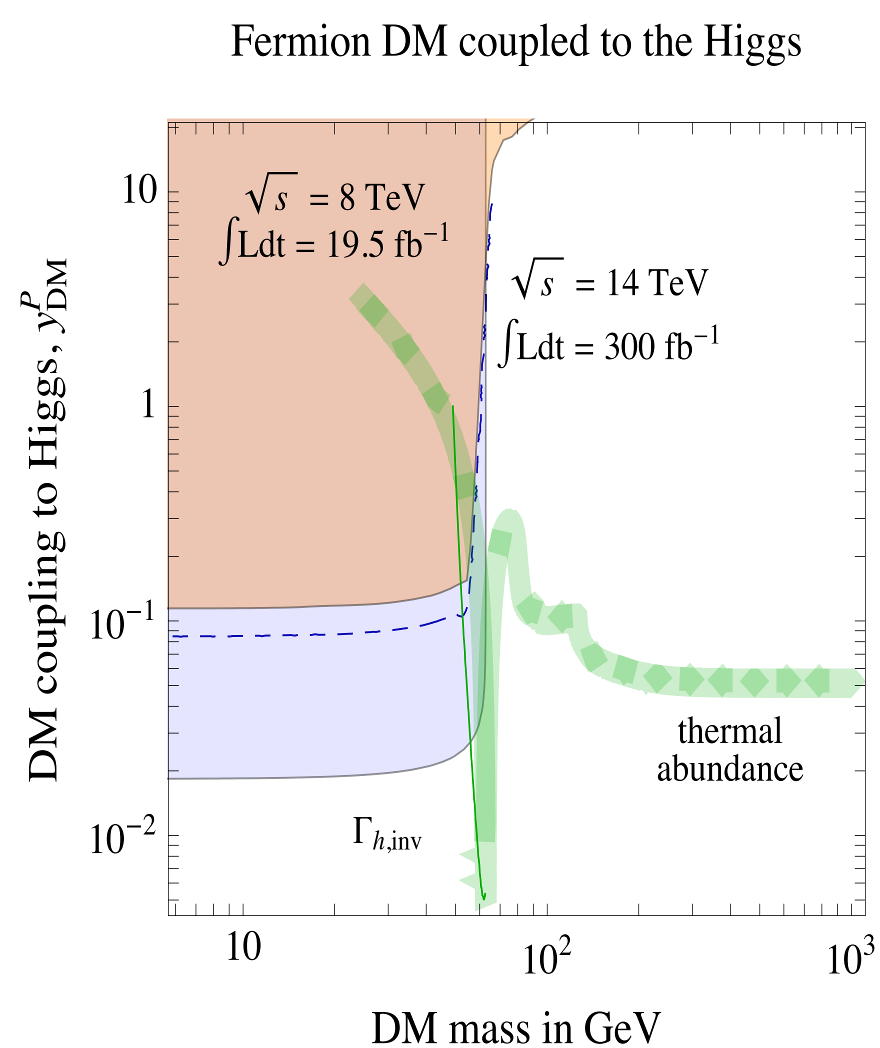
m_{DM} > 1/2 m_h Direct detection

Thermal WIMP
will be completely covered
in this scenario.

De Simone, Giudice, Strumia, 2014

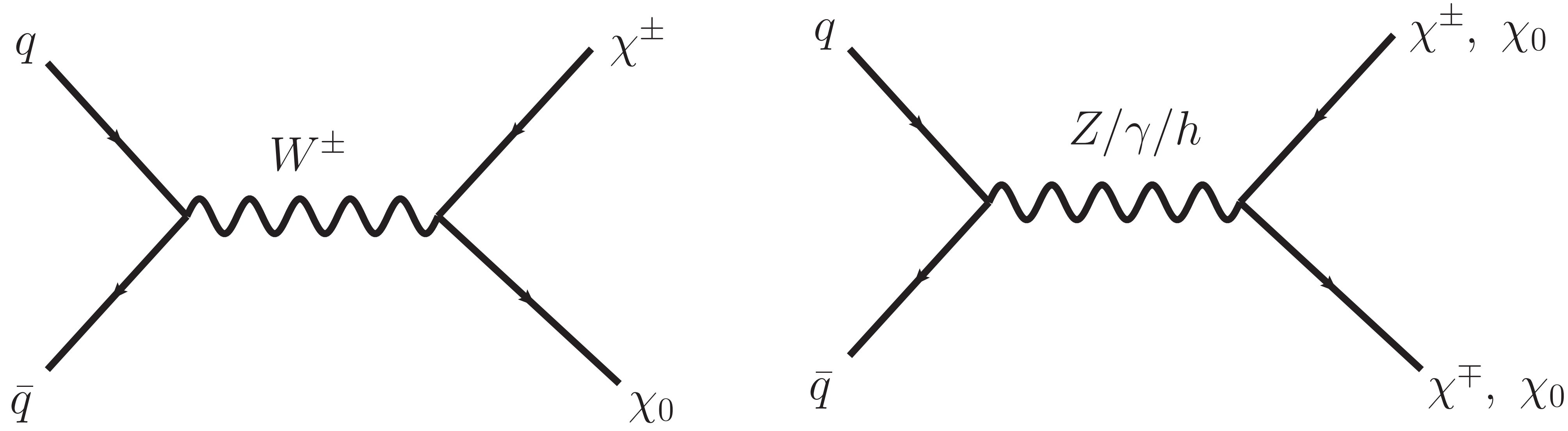
Fermionic Higgs portal.





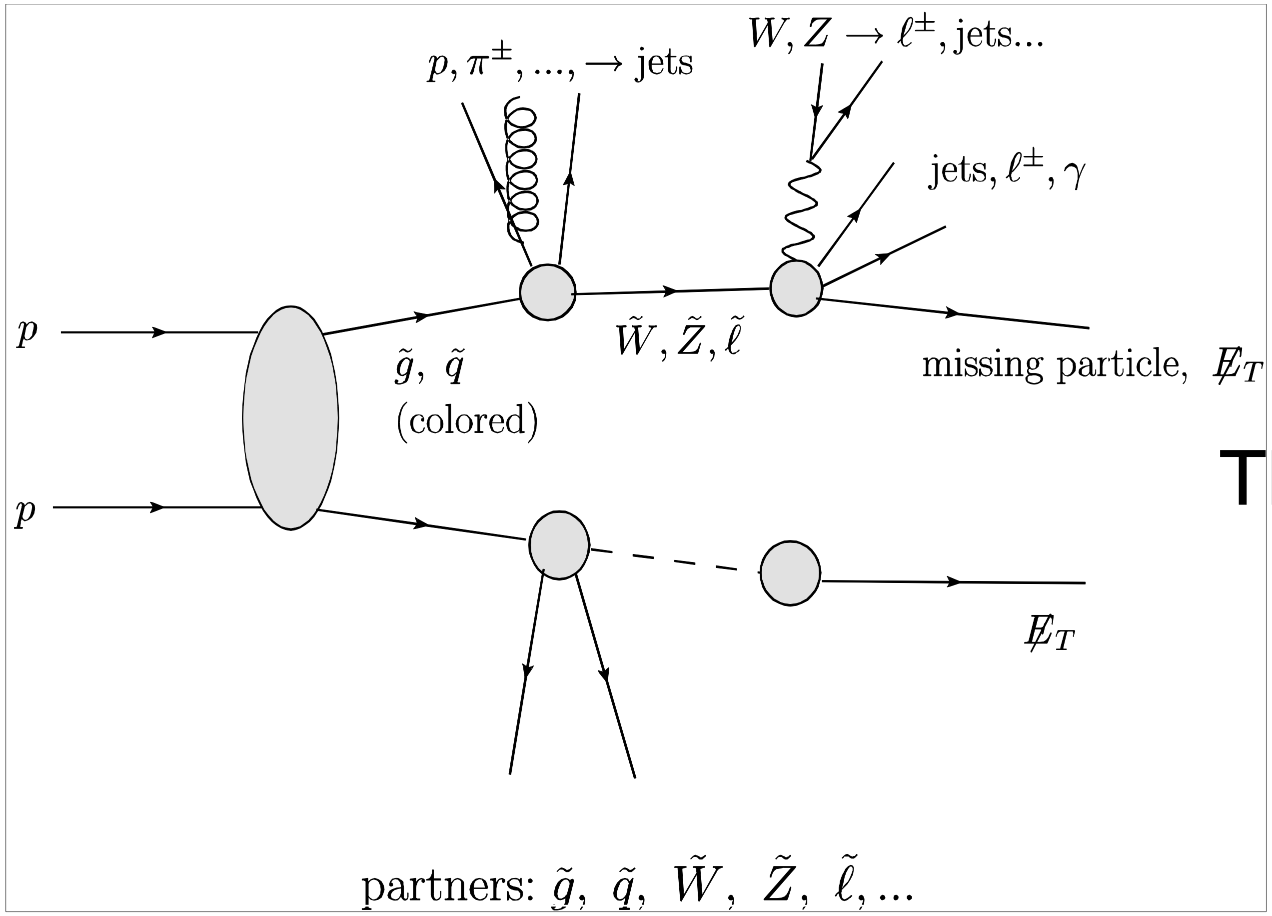
Similar story, strong constraint.
One may evade direct detection

Dark matter part of weak multiplet



- Mediated by W/Z/h.
- Predictive, no unkown particle as mediator.

SUSY DM signal in the compressed case

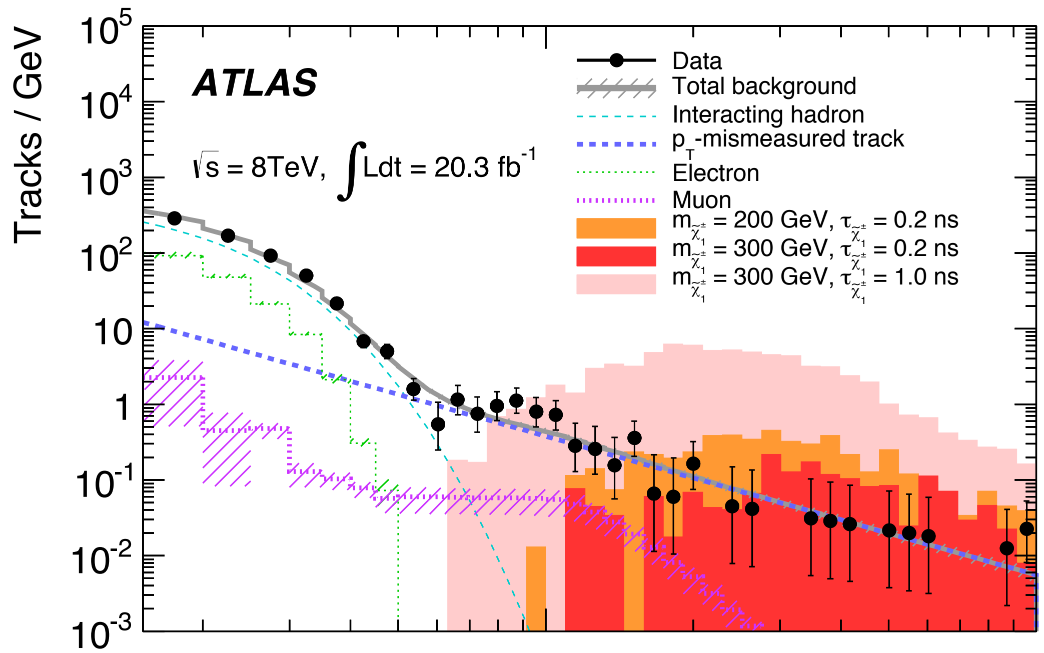


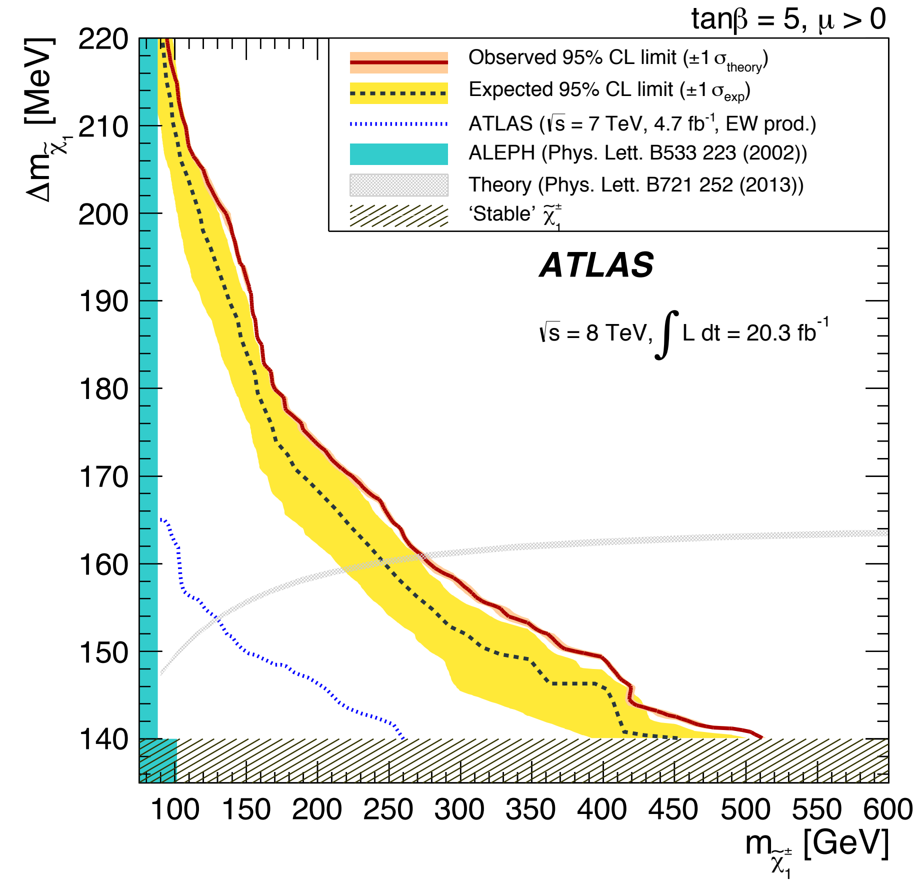
The "usual" story

SUSY DM signal in the compressed case

ATLAS search

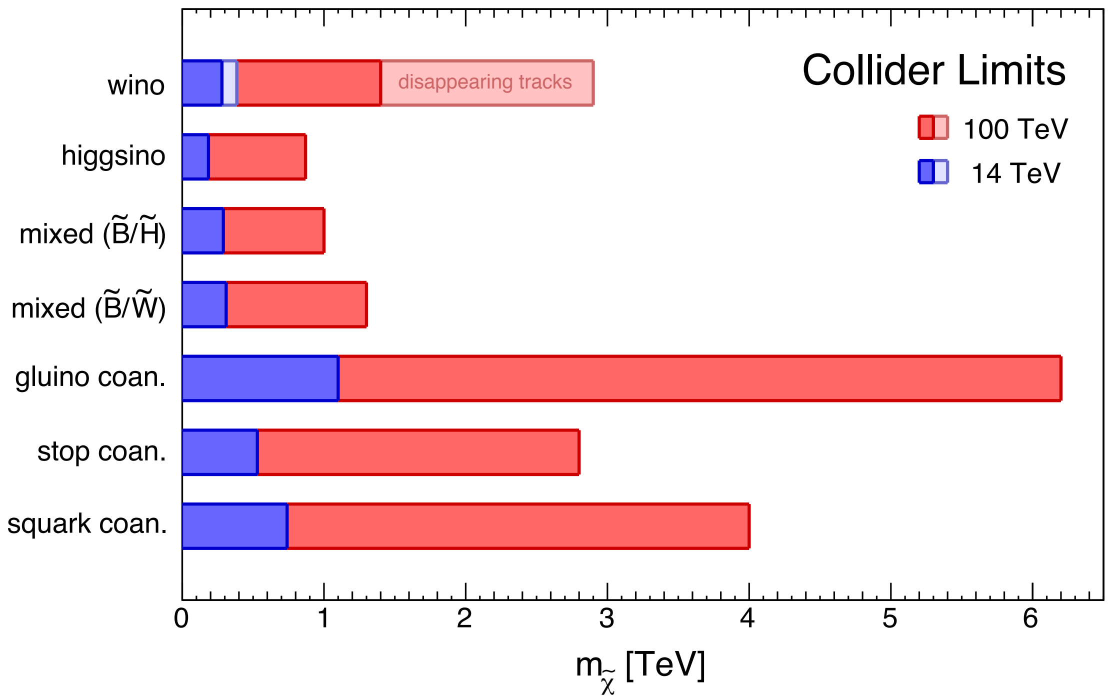
ATLAS, 1310.3675

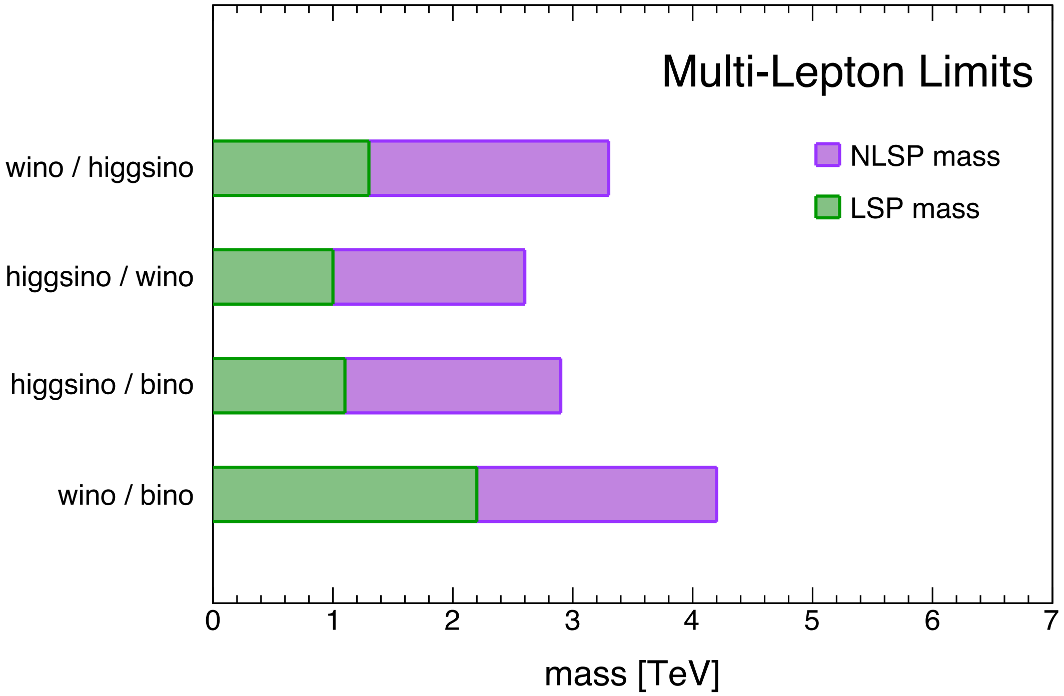




- Essentially free of physics background.
- Dominated by p_T mis-measured tracks.
- Very promising reach, much better than mono-jet

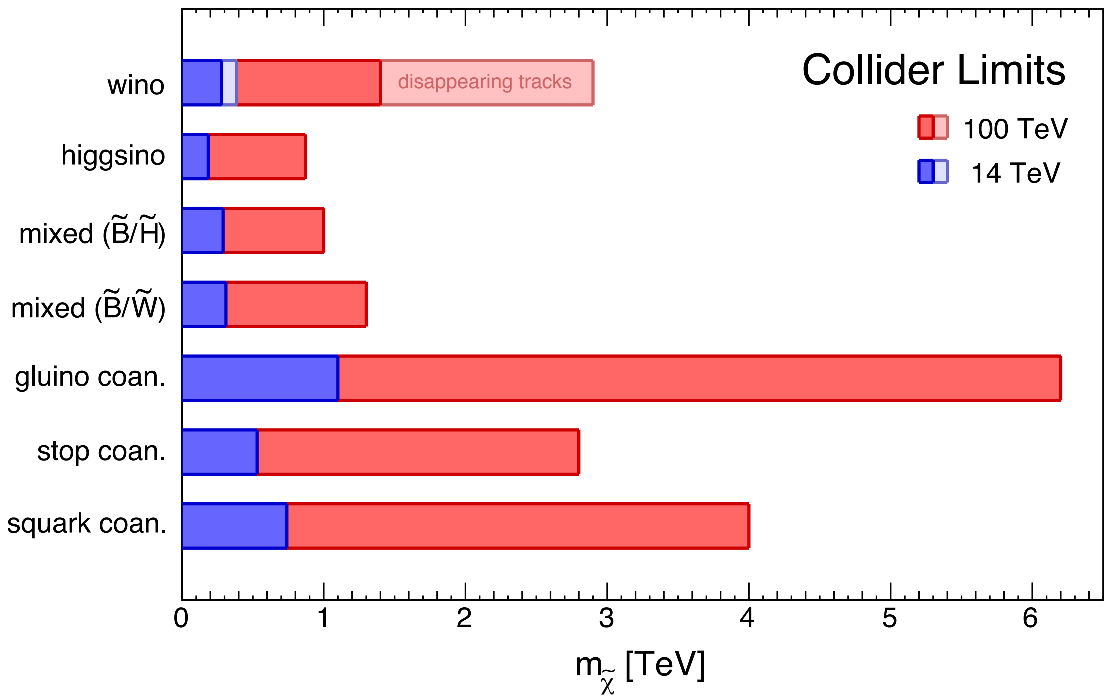
WIMP searches at colliders

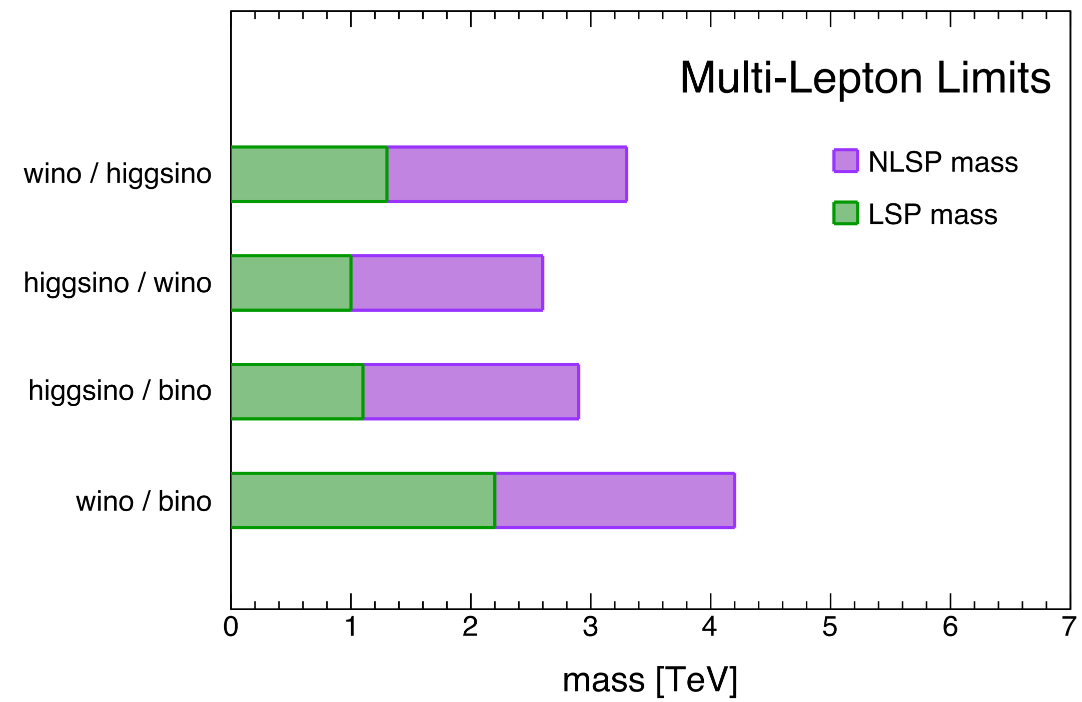




$$M_{\rm WIMP} \leq 1.8 \text{ TeV } \left(\frac{g^2}{0.3}\right)$$

WIMP searches at colliders





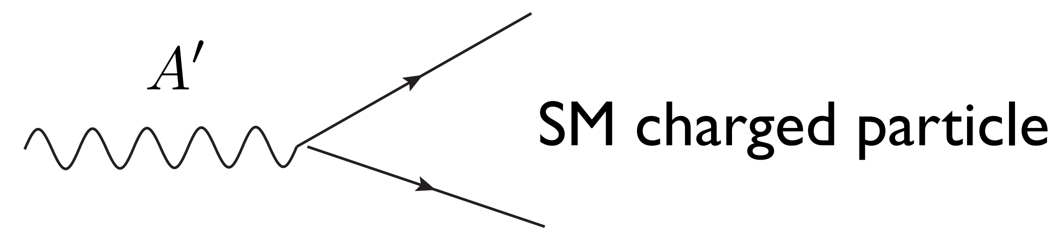
$$M_{\rm WIMP} \leq 1.8 \text{ TeV} \left(\frac{g^2}{0.3}\right)$$

Difficult searches.
In the context of future collider, opportunities of new detector/strategies.

dark photon with kinetic mixing

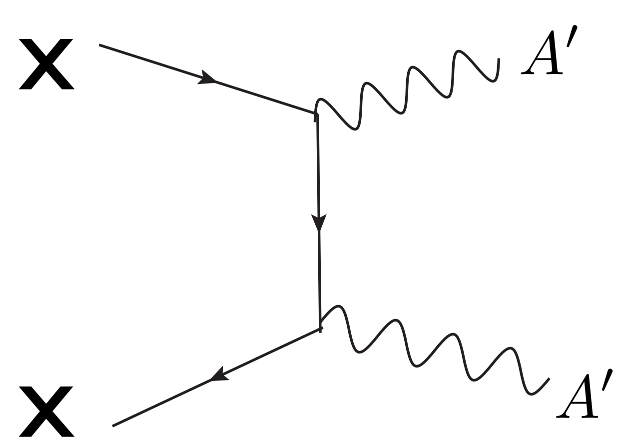
$$\epsilon F'_{\mu\nu} F^{\mu\nu} \quad \Longrightarrow \quad \epsilon e A'_\mu J^\mu_{\rm EM} + g_d A'_\mu J^\mu_{\rm dark}$$

decay before BBN:

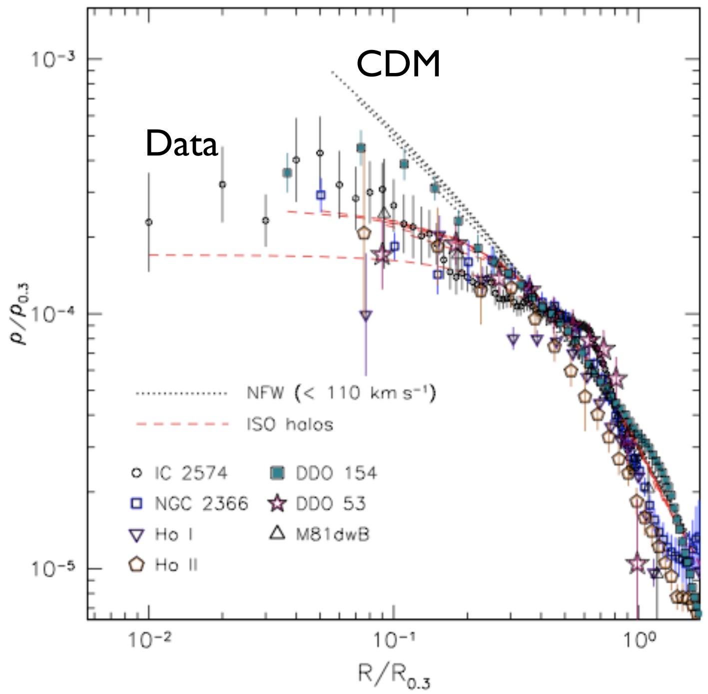


$$\epsilon > 10^{-10} \sqrt{m_{A'}/10 \text{ MeV}}$$

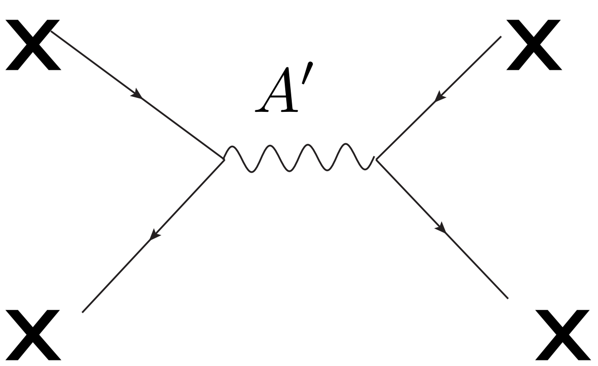
can be part of thermal relic story:



What can dark photon do?

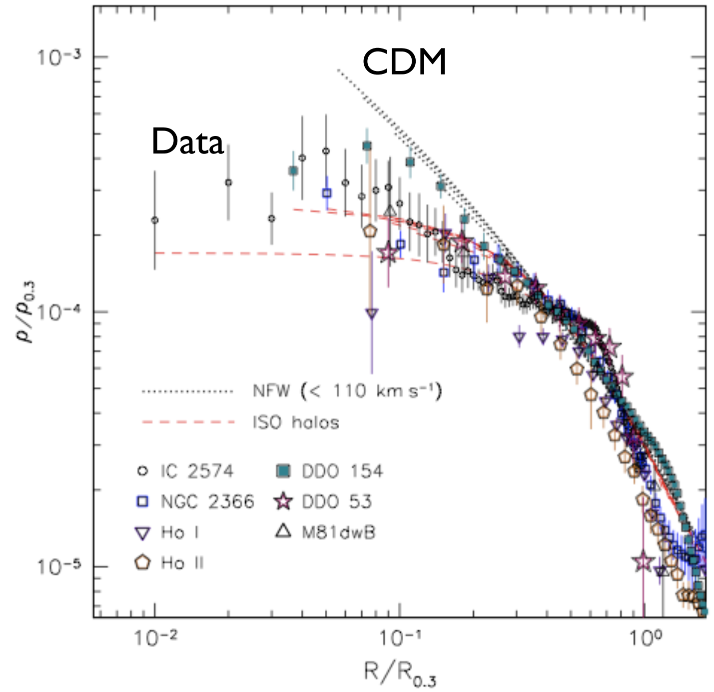


Core vs cusp, too big to fail

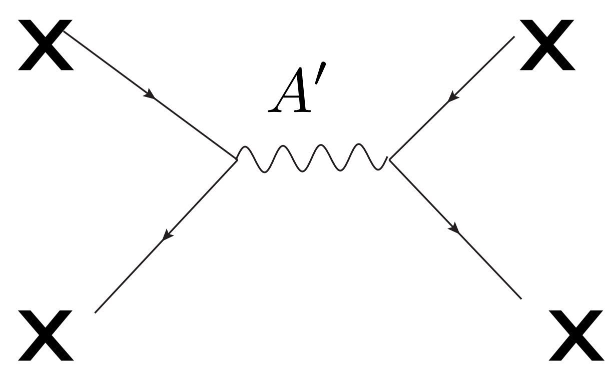


DM self interaction

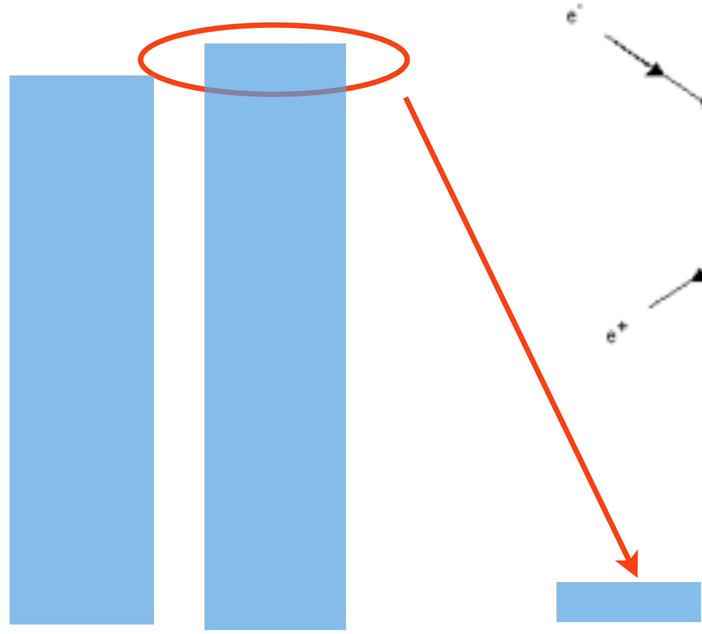
What can dark photon do?

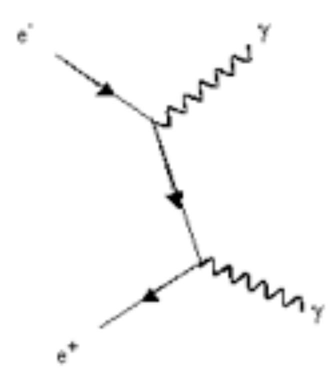


Core vs cusp, too big to fail



DM self interaction





Asymmetric DM

$$n_{\rm B} \sim n_{\rm DM}$$

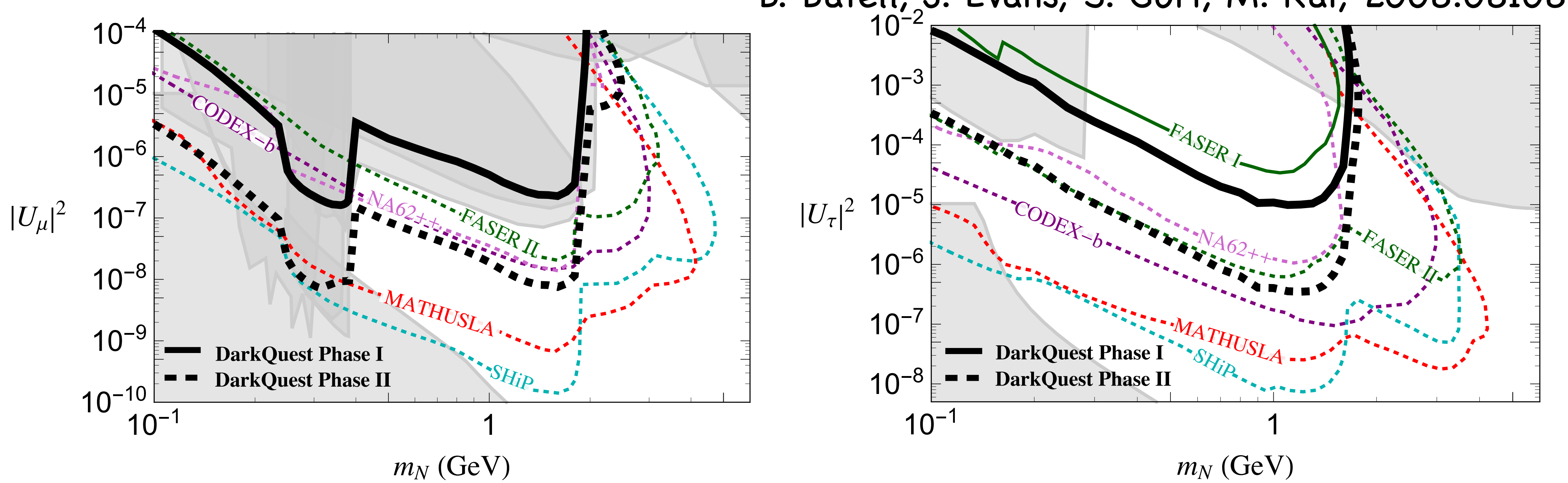
Deplete the symmetric part

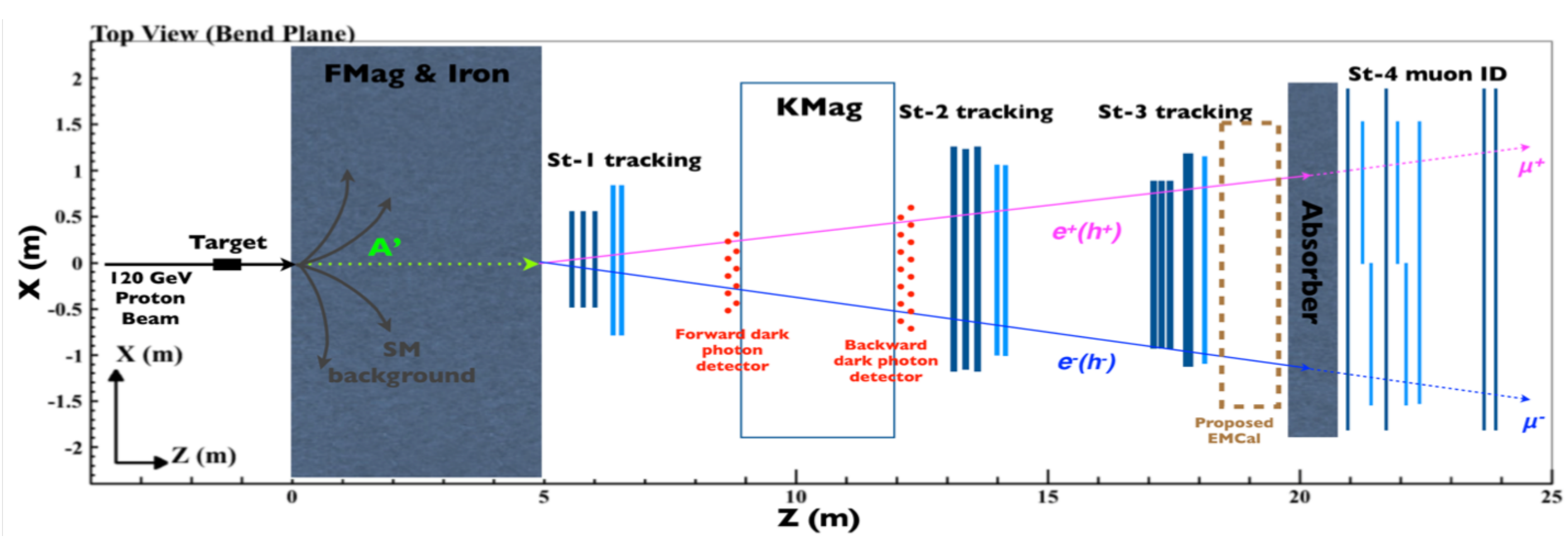
Neutrino portal example: HNL

Heavy Neutral Lepton (HNL):

$$\mathscr{L} \supset (HL)N + h.c.$$

B. Batell, J. Evans, S. Gori, M. Rai, 2008.08108





DarkQuest